

On the dynamics of vortices in the solar atmosphere

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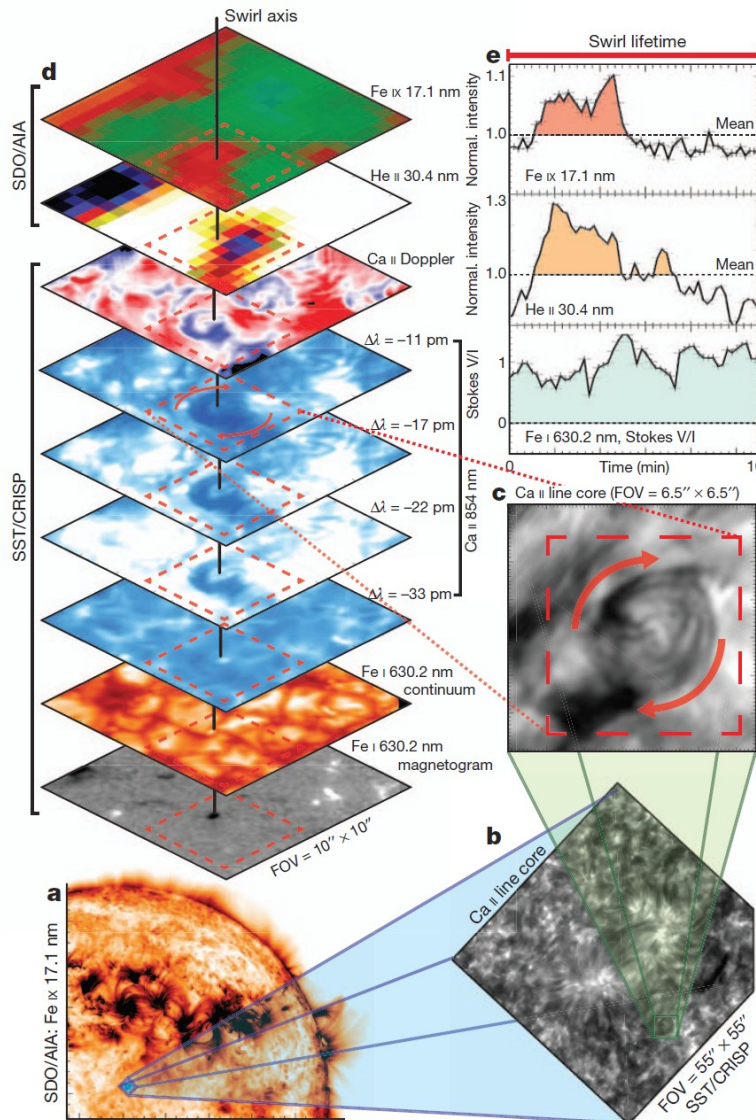
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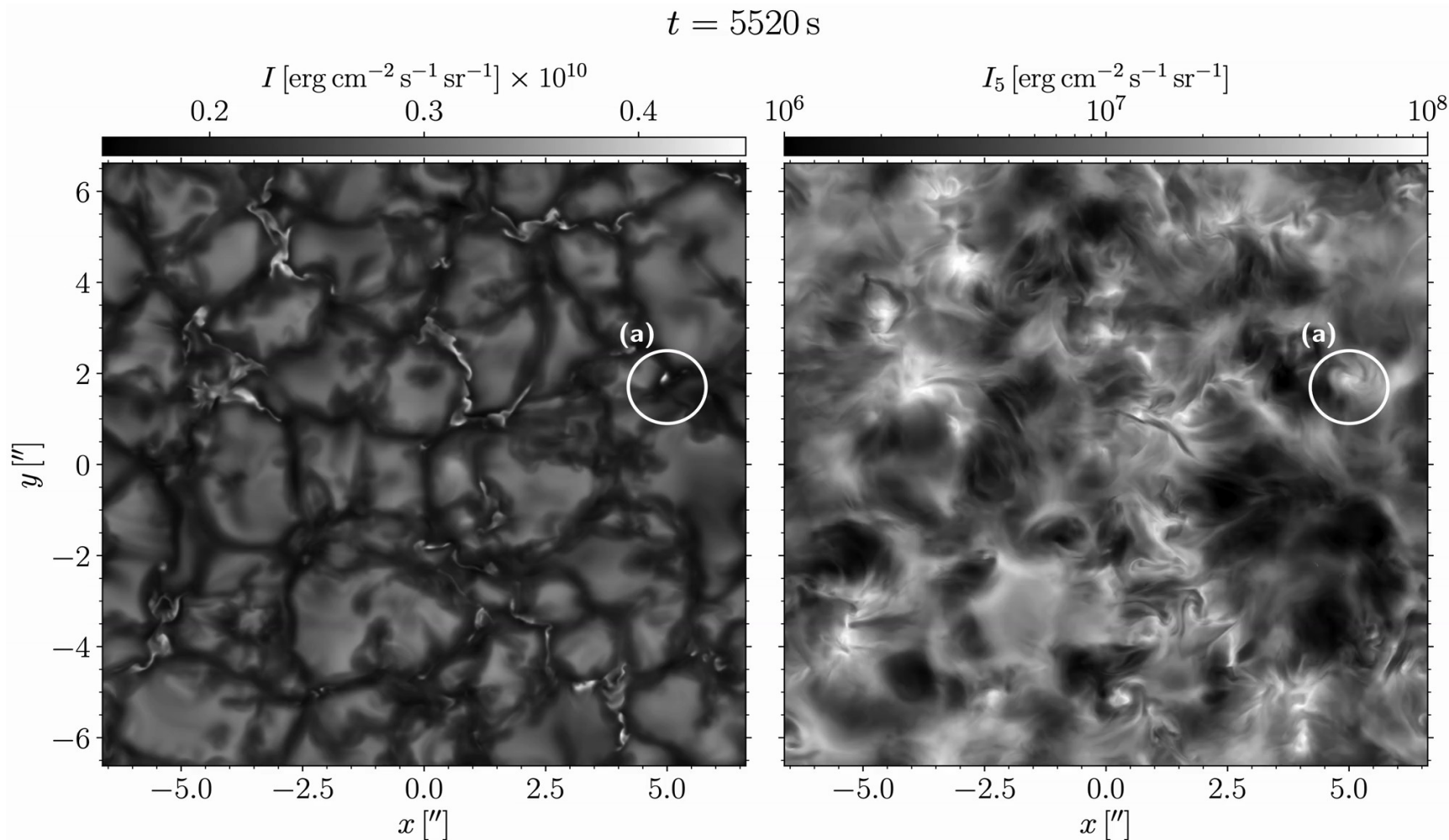
Introduction



- Small-scale swirling motions in the quiet solar atmosphere observations
- Footpoints in intergranular lanes
- Correlated with magnetic fields (Bright points)
 - “Magnetic Tornadoes”

Energy channels from the convection zone to the upper layers?

Simulations



Source: Battaglia et al., 2021

Synthetic intensity maps from a CO5BOLD simulation ($9.6 \times 9.6 \times 2.8 \text{ Mm}^3$)

Detection methods

■ Vortices

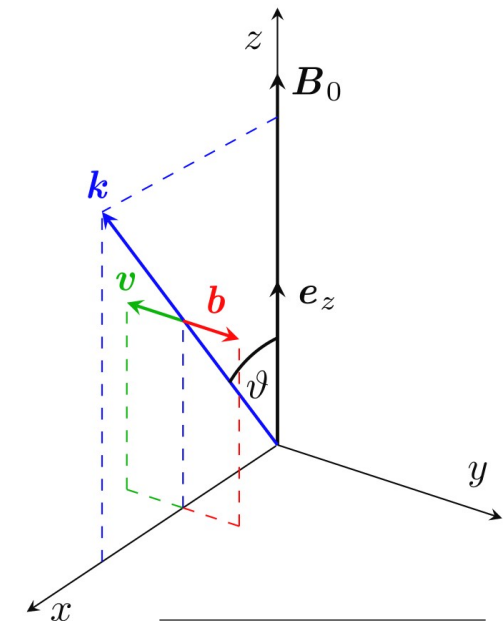
- Swirling strength λ
 - “Better version” of vorticity
 - Evolution equation

$$\begin{aligned}
 \frac{d}{dt}\lambda &= -2\lambda\lambda_{\text{cr}} + 2\text{Im}\left(\mathcal{P}^{-1}\mathcal{M}\mathcal{P}\right)_{22}, \\
 &= -2\lambda\lambda_{\text{cr}} & T_{\lambda}^1 \\
 &\quad - 2\text{Im}\left\{\mathcal{P}^{-1}\left[\nabla\left(\frac{1}{\rho}\nabla p_g\right)\right]\mathcal{P}\right\}_{22} & T_{\lambda}^2 \\
 &\quad - 2\text{Im}\left\{\mathcal{P}^{-1}\left[\nabla\left(\frac{1}{\rho}\nabla p_m\right) - \left(\nabla\frac{1}{\rho}\right)(\mathbf{B}\cdot\nabla)\mathbf{B}\right]\mathcal{P}\right\}_{22} & T_{\lambda}^3 \\
 &\quad + 2\text{Im}\left\{\mathcal{P}^{-1}\left[\frac{1}{\rho}\nabla((\mathbf{B}\cdot\nabla)\mathbf{B})\right]\mathcal{P}\right\}_{22} & T_{\lambda}^4 \\
 &\quad - 2\text{Im}\left\{\mathcal{P}^{-1}\left[\nabla(\nabla\Phi)\right]\mathcal{P}\right\}_{22}. & T_{\lambda}^5
 \end{aligned}$$

Source: Canivete Cuissa & Steiner, 2020

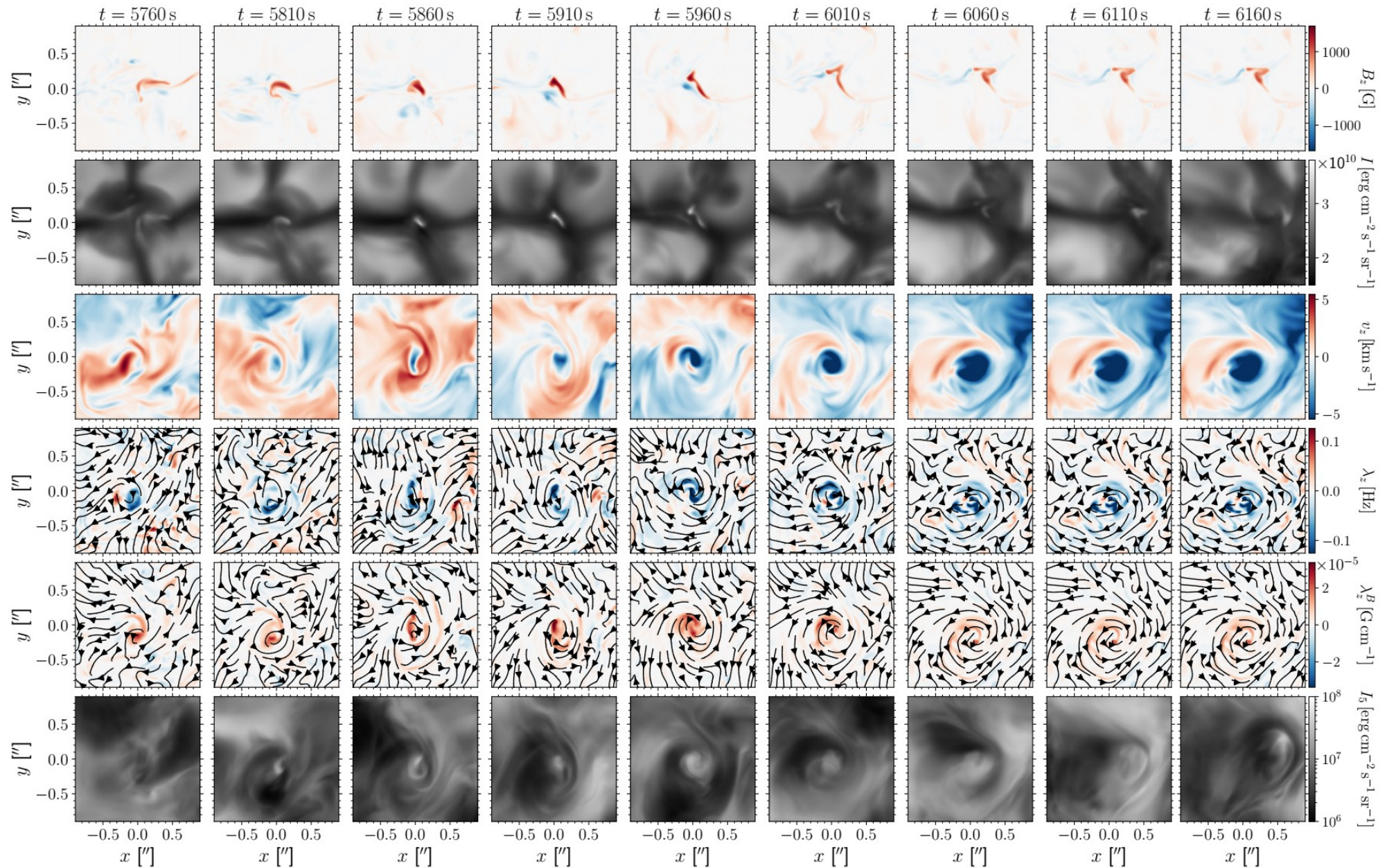
■ Alfvén waves

- Orientation of perturbations (Magnetic swirling strength λ^B)
- Propagate at Alfvén speed $v_A = B_0/\sqrt{4\pi\rho}$
- Driven by magnetic tension



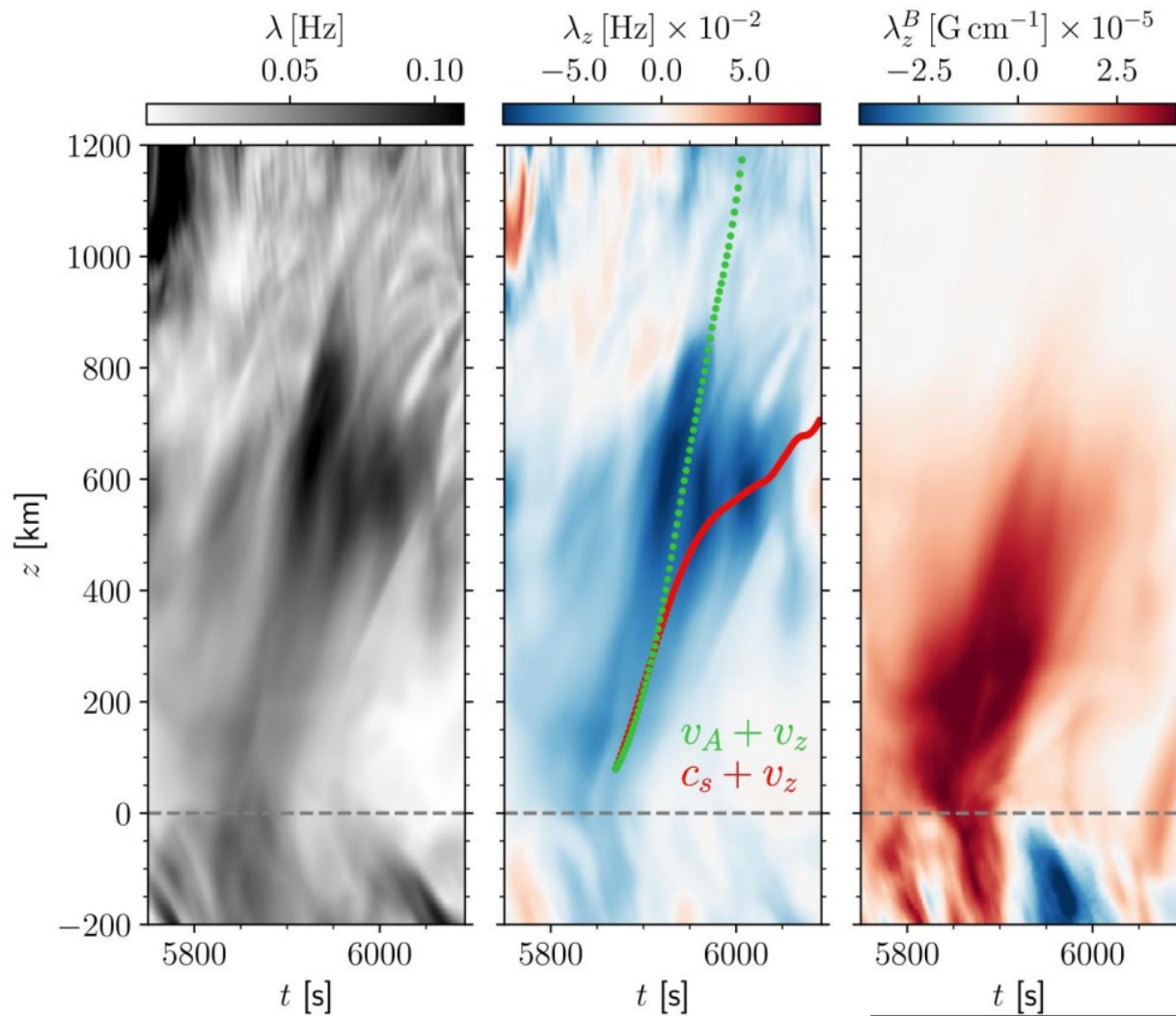
Source: Battaglia et al., 2021

Swirl event



Time sequence of the swirl event at $z=0$ km (B_z) and $z=700$ km ($v_z, \lambda_z, \lambda_z^B$)

Swirl event – Perturbation propagation



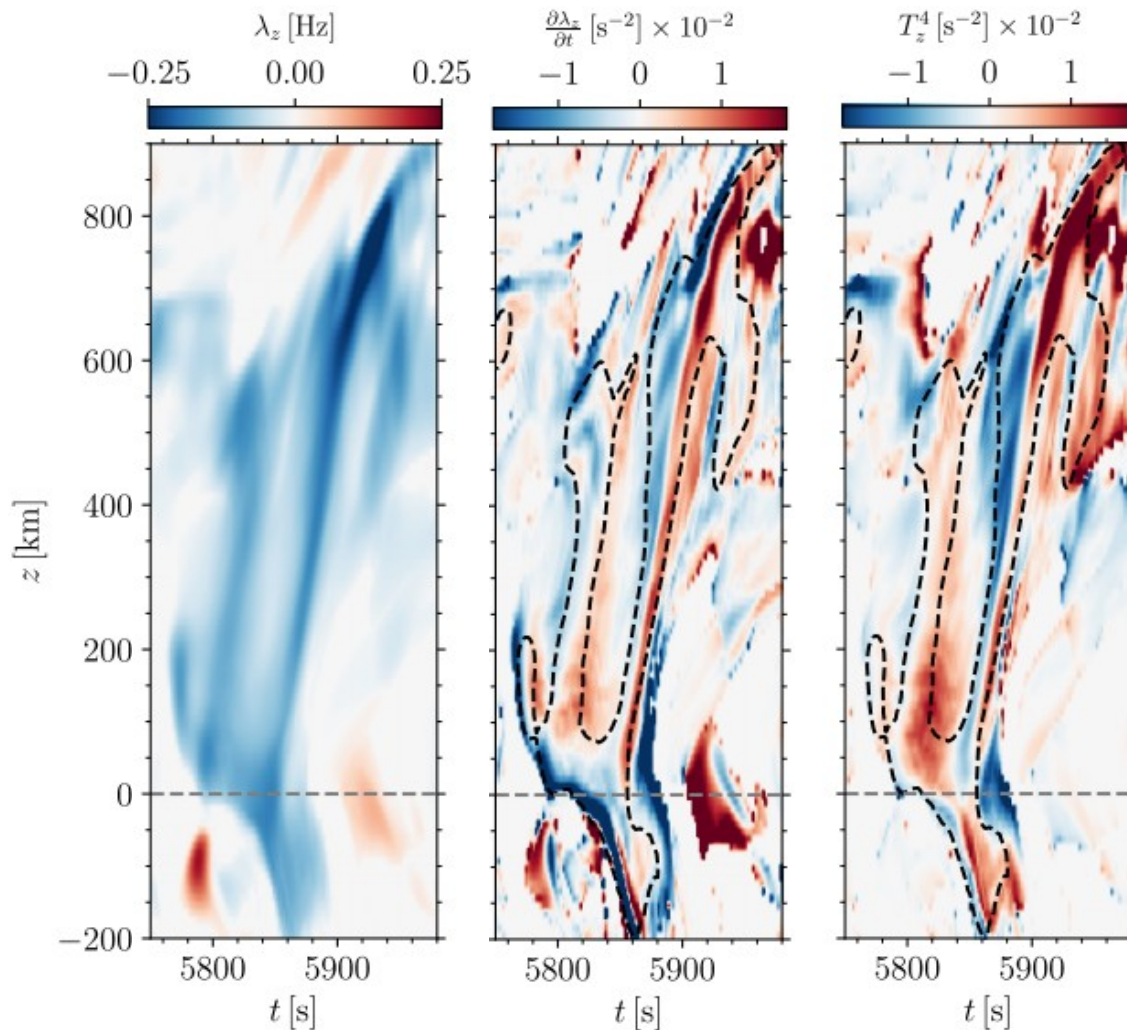
Source: Battaglia et al., 2021

Averaged time-distance diagrams over $150 \times 150 \text{ km}^2$

- Vertical swirl
- Photospheric origin
- **Upward propagation**
- **“Pulse”** (not oscillatory)

- ✓ Propagation at Alfvén speed
- ✓ Magnetic field twist opposite to vortex rotation
- ? Driven by magnetic tension

Swirl event - Dynamics



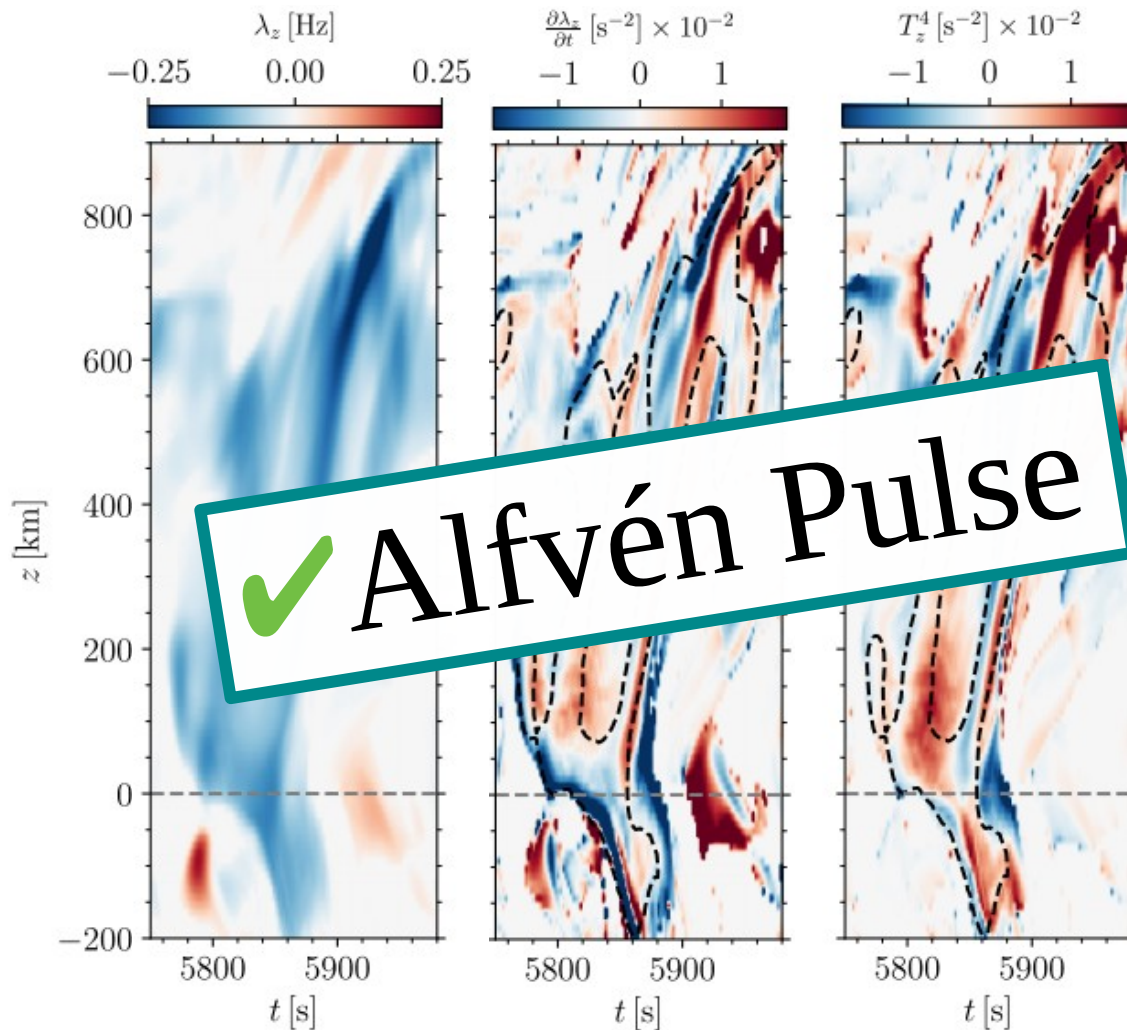
Source: Battaglia et al., 2021

- **Local production** of λ corresponds to propagation of pulses
- Propagation:
Magnetic tension

- ✓ Propagation at Alfvén speed
- ✓ Magnetic field twist opposite to vortex rotation
- ✓ Driven by magnetic tension

Single line-of-sight time-distance diagrams

Swirl event - Dynamics



Source: Battaglia et al., 2021

- **Local production** of λ corresponds to propagation of pulses
- Propagation: **Magnetic tension**

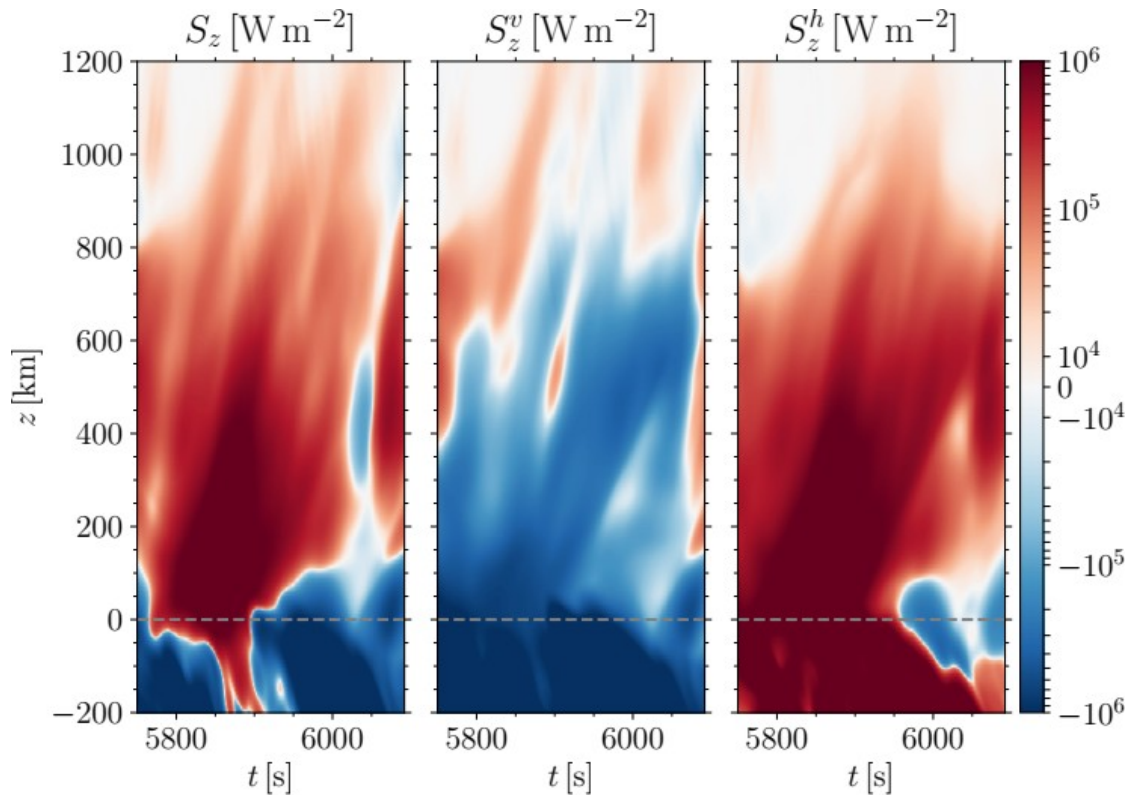
- ✓ Propagation at Alfvén speed
- ✓ Magnetic field twist opposite to vortex rotation
- ✓ Driven by magnetic tension

Single line-of-sight time-distance diagrams

Swirl event - Energetics

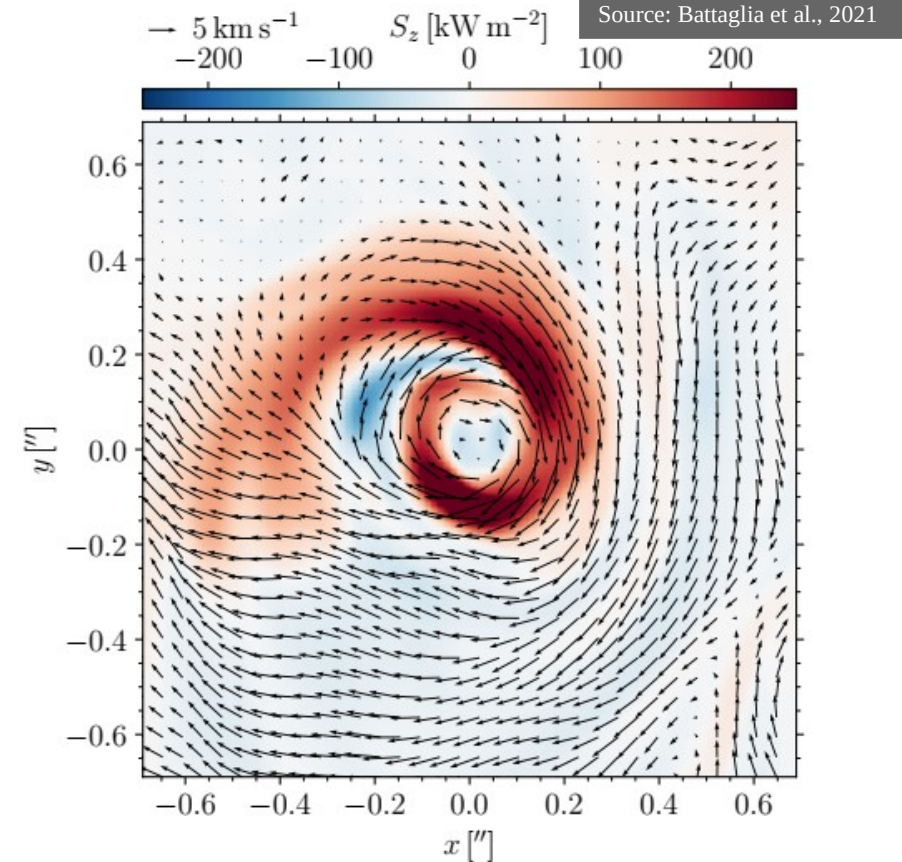
- Poynting flux :

$$S_z = \underbrace{\frac{1}{4\pi} v_z (B_x^2 + B_y^2)}_{S_z^v} - \underbrace{\frac{1}{4\pi} B_z (v_x B_x + v_y B_y)}_{S_z^h}$$



Source: Battaglia et al., 2021

Averaged time-distance diagrams



Source: Battaglia et al., 2021

Horizontal section at $z = 700$ km

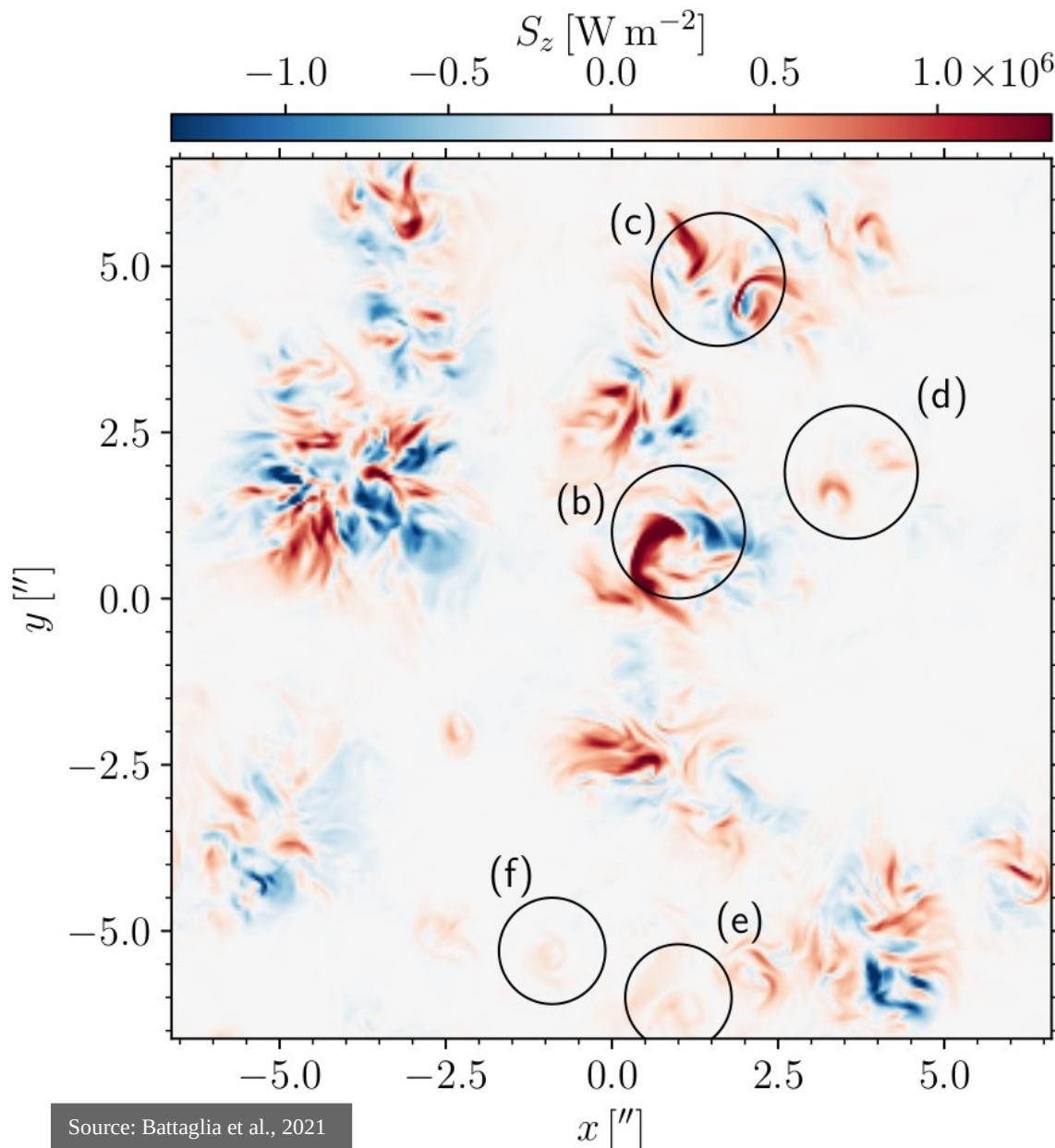
- At base of chromosphere :

$$\bar{S}_z = 1.2 \text{ kW m}^{-2}$$

- ✗ Not enough to compensate radiative losses!

- Swirling motions carry up energy

Energetics – the whole picture



Source: Battaglia et al., 2021

Horizontal section at $z = 700 \text{ km}$

- Contribution of single, isolated swirls (e,d,f) is small
- Large contributions come from **large and complex magnetic footpoints** (b,c)
- Multiple swirls can be seen contemporaneously in these regions

“Superposition of swirls”

➤ At base of chromosphere:

$$\bar{S}_z = 12.8 \pm 6.5 \text{ kW m}^{-2}$$

✓ Enough to compensate radiative losses!

Conclusions

- The event studied has **Alfvénic nature**:
 - Magnetic twist has opposite orientation to plasma rotation
 - Upward propagation with local Alfvén speed
 - Magnetic tension is the driving force**“Alfvén pulse”**
- Vertical positive Poynting flux due to swirling motions
 - Most contributions from large and complex magnetic regions
 - Potentially contribute to **heat the upper layers**

➤ Canivete Cuissa, J. R. & Steiner, O. 2020, A&A, 639, A118

➤ Battaglia, A. F., Canivete Cuissa, J. R., Calvo, F., Bossart, A. A., & Steiner, O. 2021, A&A, 649, A121

