

Spectral modelling of kilonovae in 3D NLTE

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Supervisor: **Anders Jerkstrand**

Co-supervisor: **Stephan Rosswog**

SirEN conference, 25. 06. 11.



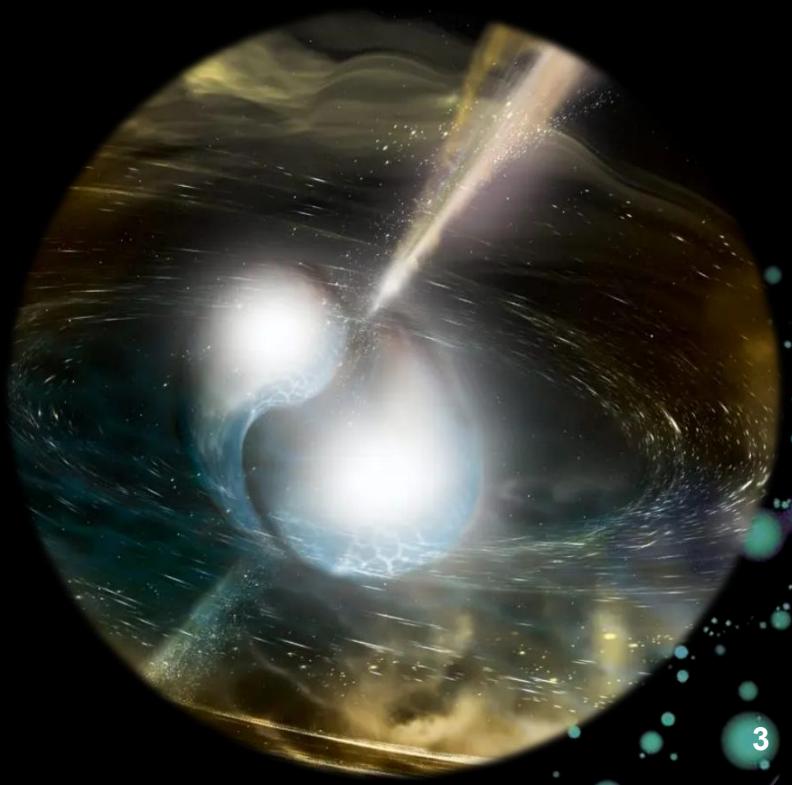
*Knut and Alice
Wallenberg
Foundation*

Spectral modelling of kilonovae in 3D NLTE



Kilonovae (KN)

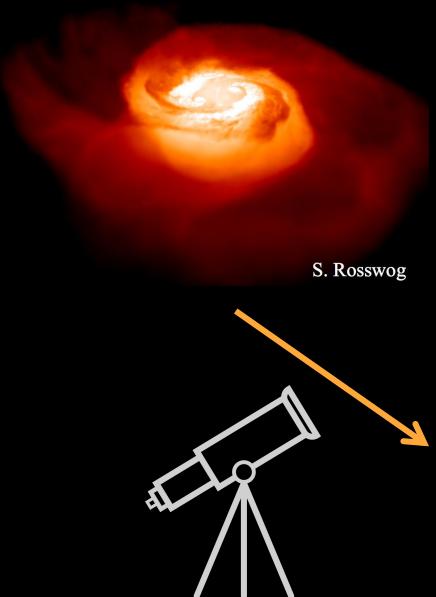
- Transients powered by the radioactive r-process elements
- Compact object mergers (NS-NS, NS-BH)



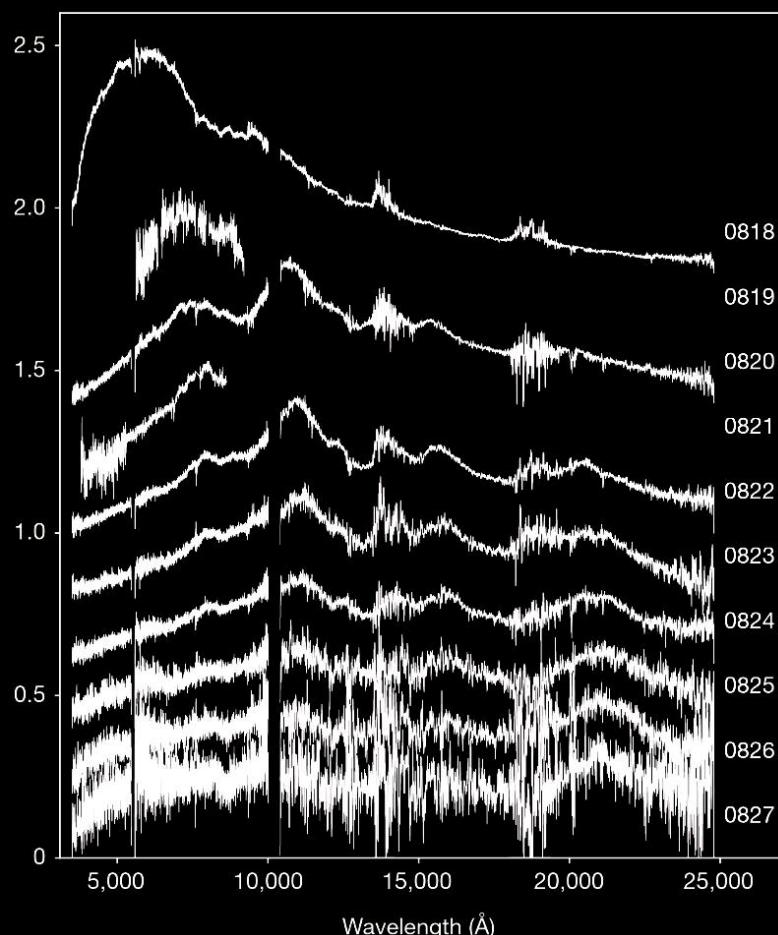
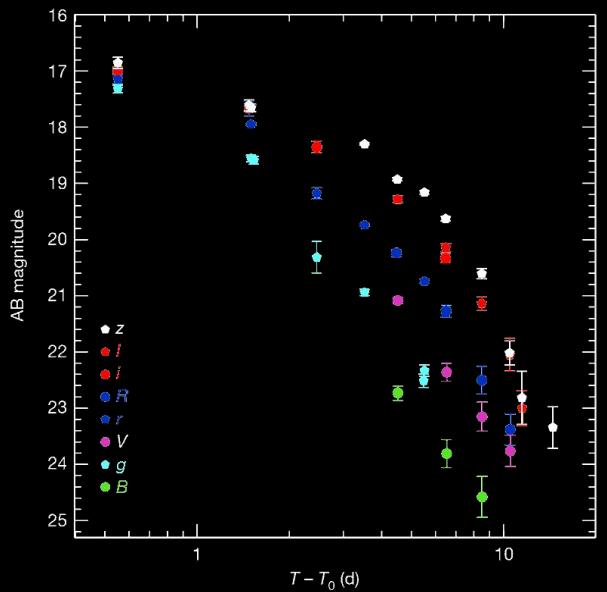
NLTE spectral modeling

Spectral modelling

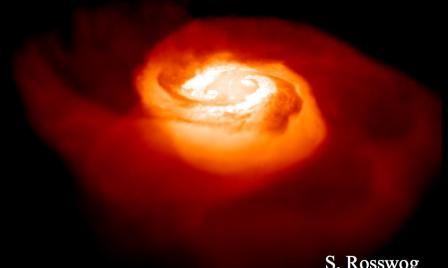
- Reproduce light curves, spectra
- Identify elements (+abundances)



S. Rosswog



Ingredients of spectral modeling



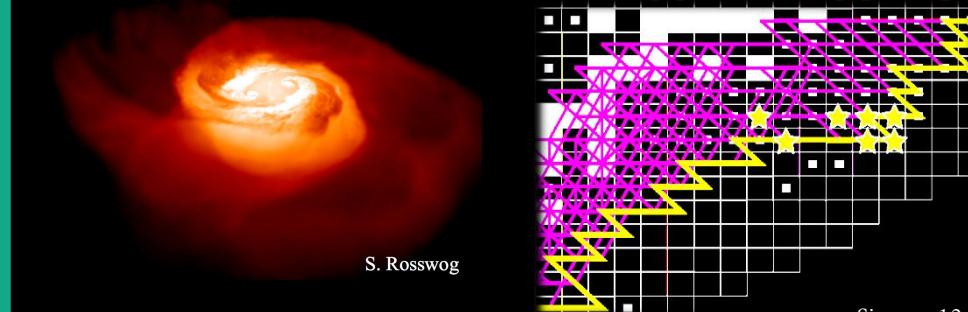
S. Rosswog

NSM model
(hydrodynamics -
numerical relativity)



Grid, ρ , v

Ingredients of spectral modeling



NSM model
(hydrodynamics -
numerical relativity)



Grid, ρ , v

Nucl. reaction
network



Composition (Y_e)
Radioactive heat

Ingredients of spectral modeling

NSM model
(hydrodynamics -
numerical relativity)



Grid, ρ , v

Nucl. reaction
network

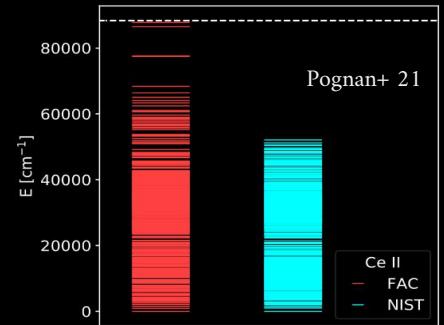
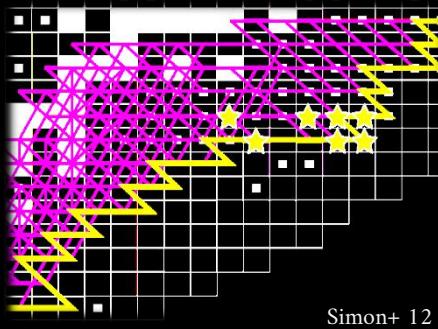


Composition (Y_e)
Radioactive heat

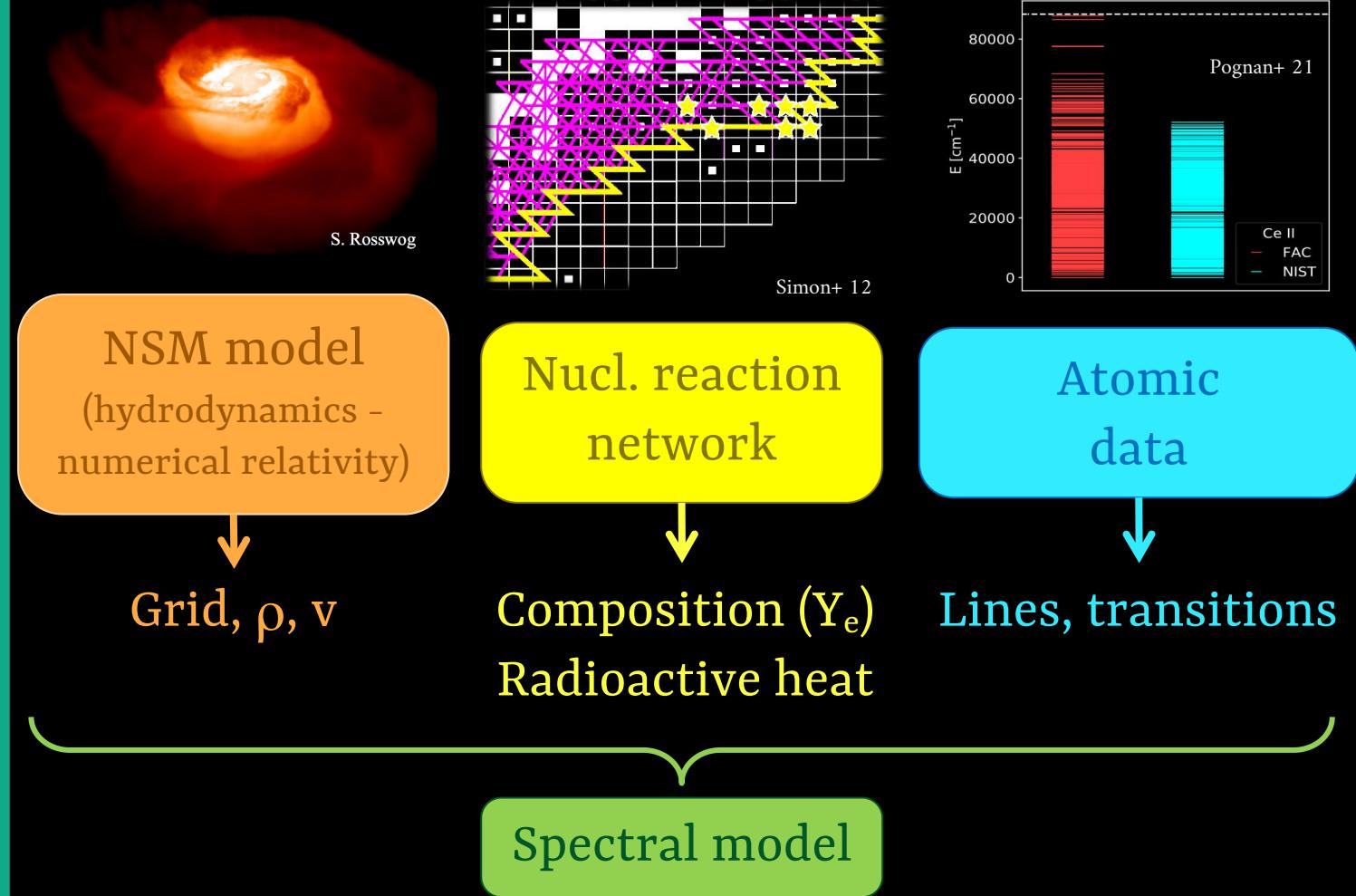
Atomic
data



Lines, transitions



Ingredients of spectral modeling



Nebular phase

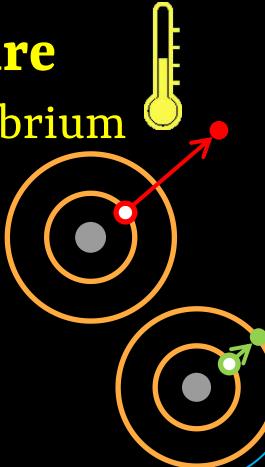
- **Nebular phase** = low density, emission lines
- Important for...
 - Nucleosynthesis (bulk of material, emission lines)
 - Physical conditions (e.g. morphology, velocity)
 - For KNe, after \sim a week
- **NLTE** (non-local thermodynamic equilibrium)
 - Collisional excitations and de-excitations \ll radiative processes
 - Full rate equations

MATTER

Physical conditions:

- Temperature

Thermal equilibrium



- Ionisation

Saha eq.

- Excitation

Boltzmann eq.

Emit

LTE

PHOTONS

Radiation field

Density, spectrum

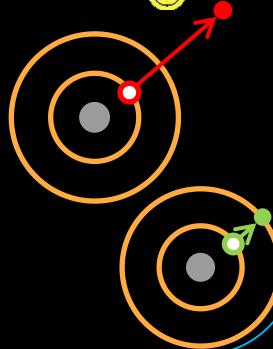
MATTER

Physical conditions:

- Temperature



- Ionisation



- Excitation

Emit

NLTE

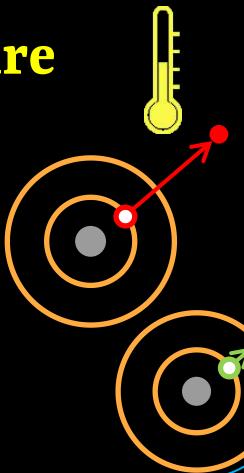
PHOTONS

Radiation field

Density, spectrum

MATTER

- Physical conditions:**
- Temperature
 - Ionisation
 - Excitation



Emit

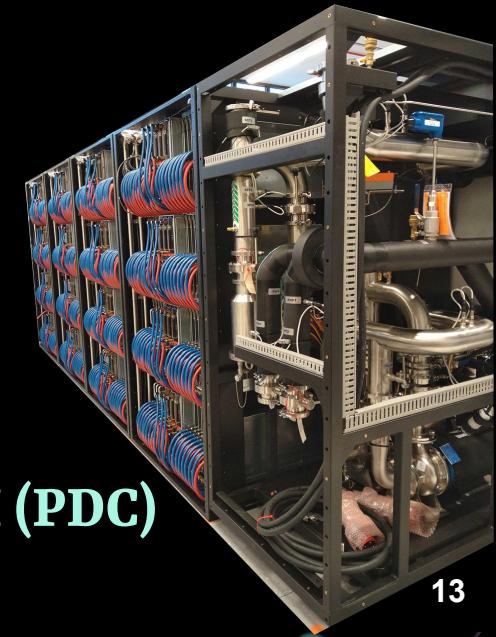
NLTE

Excite
(+ non-local)

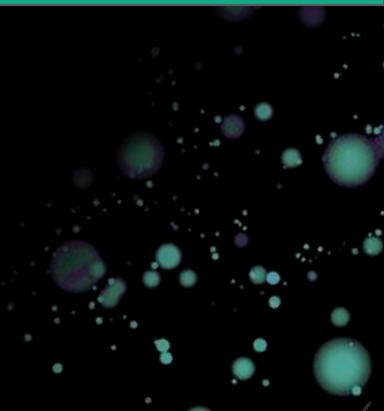
PHOTONS

Radiation field
Density, spectrum

Dardel, KTH (PDC)



ExTraSS

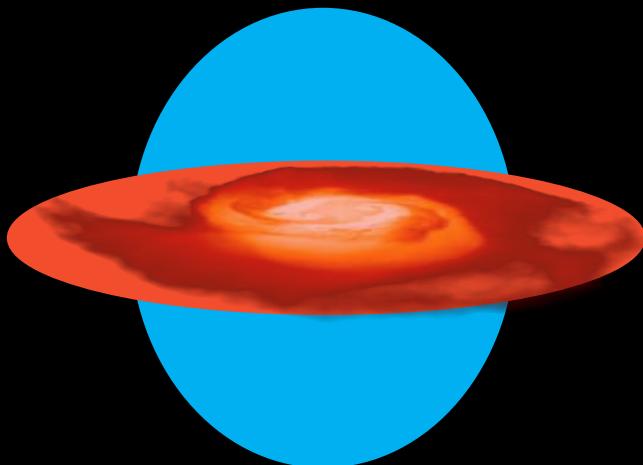


Stockholm codes for spectral modeling

- NLTE spectral codes for transients
- 1D: **SUMO** (SUpernova MOnte carlo code) - *A. Jerkstrand, Q. Pognan*
3D: **ExTraSS** (EXplosive TRAnsient Spectral Simulator) - *AJ, B. van Baal*
- My PhD: add KN physics to ExTraSS

	SUMO	EXTRASS
NLTE	✓	✓
3D	✗	✓
SNe	✓	✓
KNe	✓	...

Ejecta components



**Disk wind / NS driven
outflows (polar)**

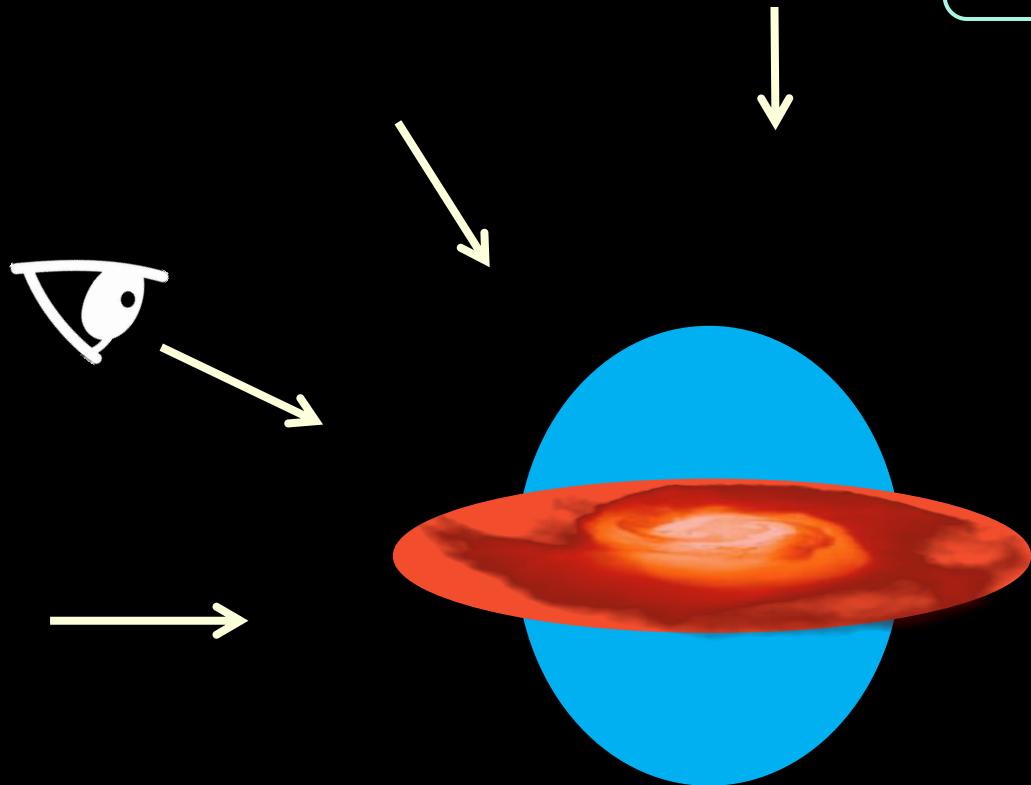
lighter r-nuclei, high Y_e
 $v \sim 0.1 c$

**Dynamical
ejecta (equatorial)**

heavy r-nuclei, low Y_e
 $v \sim 0.3 c$

Ejecta components

3D needed!!



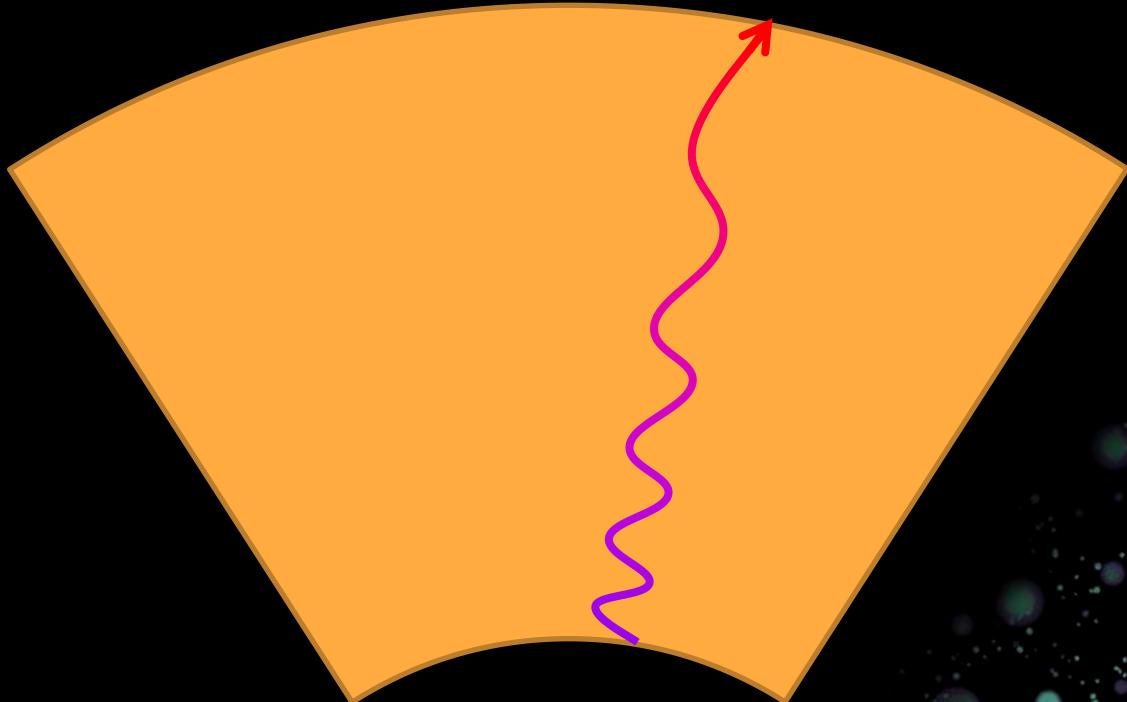
Disk wind / NS driven
outflows (polar)

lighter r-nuclei, high Y_e
 $v \sim 0.1 c$

Dynamical
ejecta (equatorial)

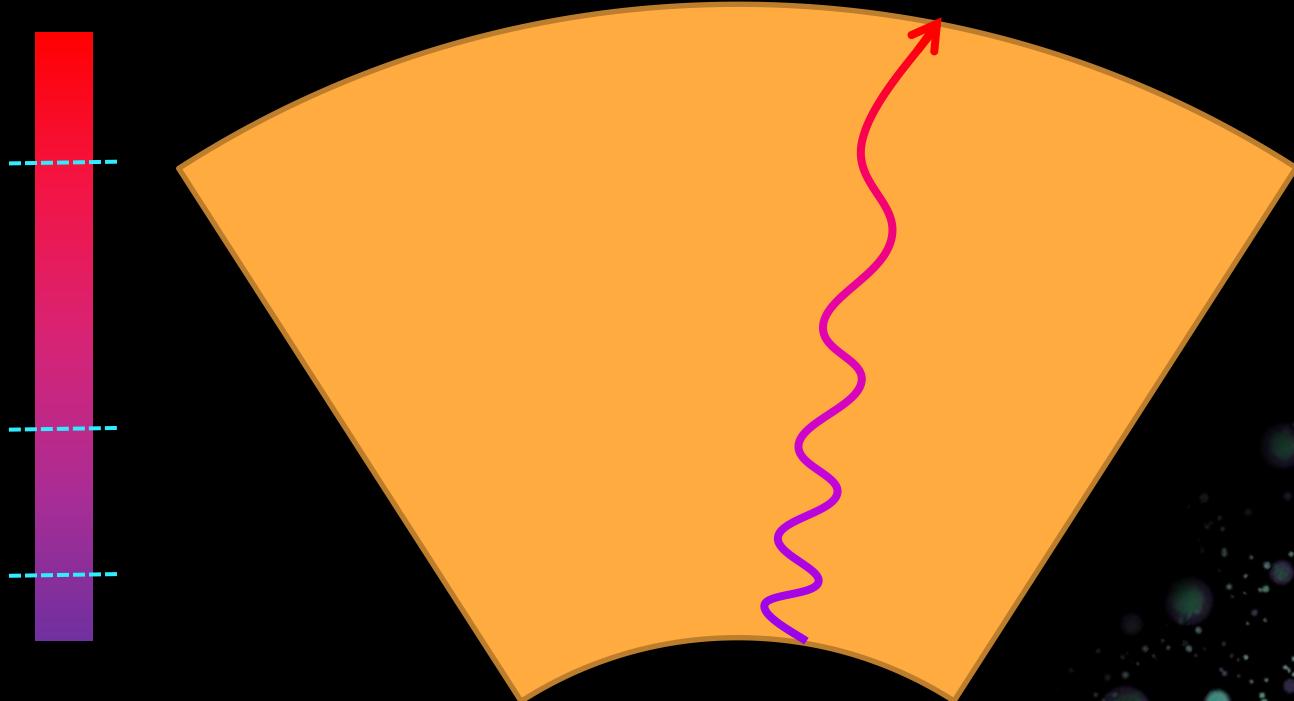
heavy r-nuclei, low Y_e
 $v \sim 0.3 c$

Line-to-line photon transfer



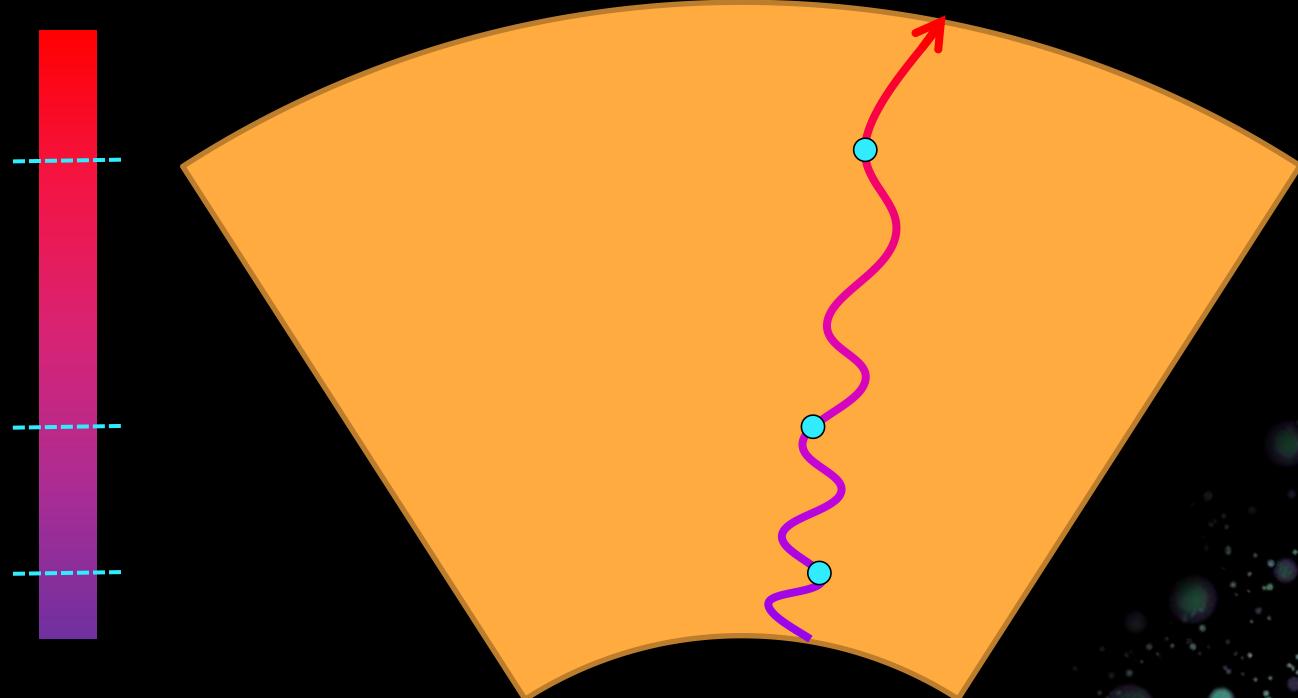
Line-to-line photon transfer

Potential
atomic lines:

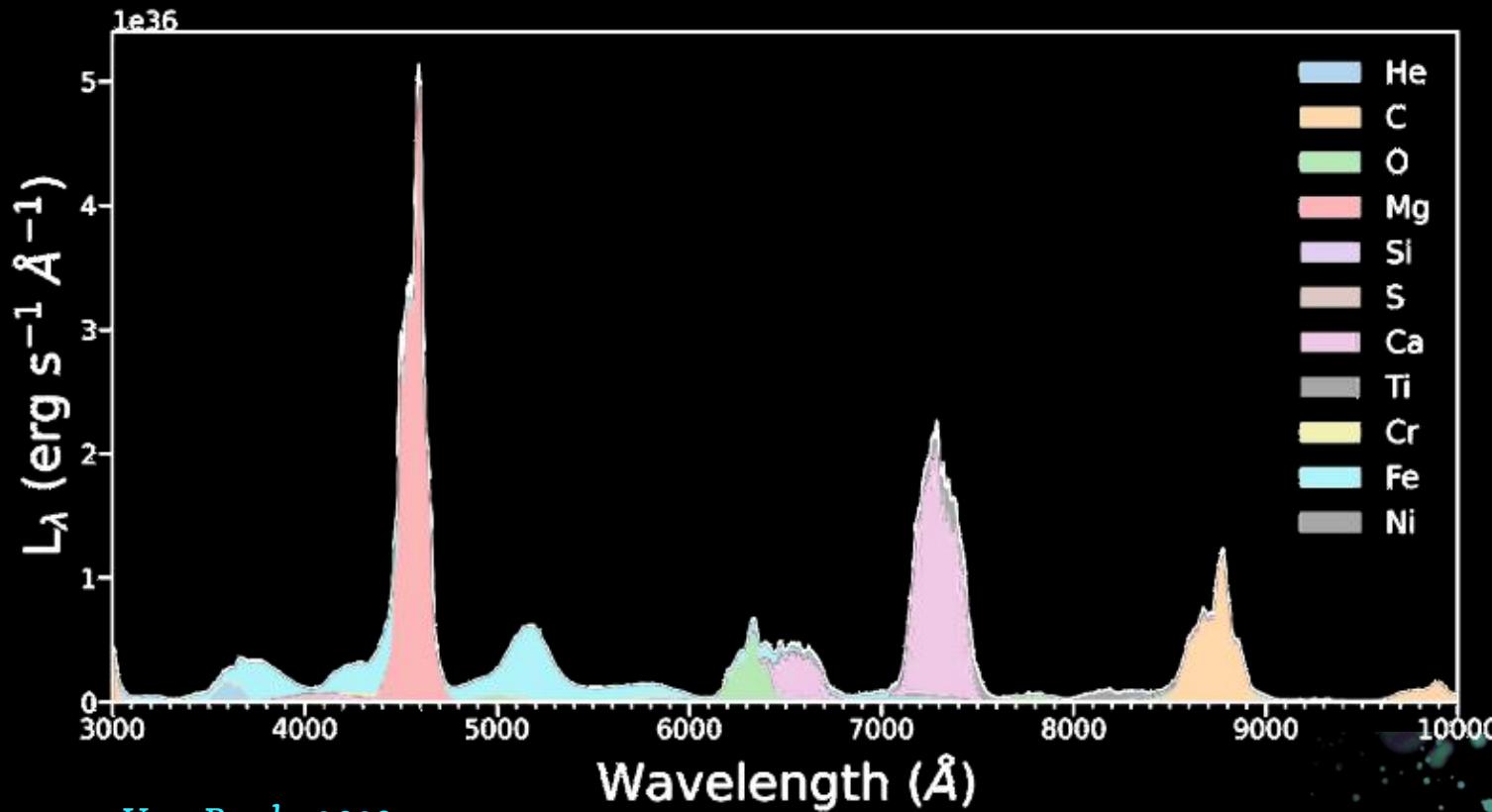


Line-to-line photon transfer

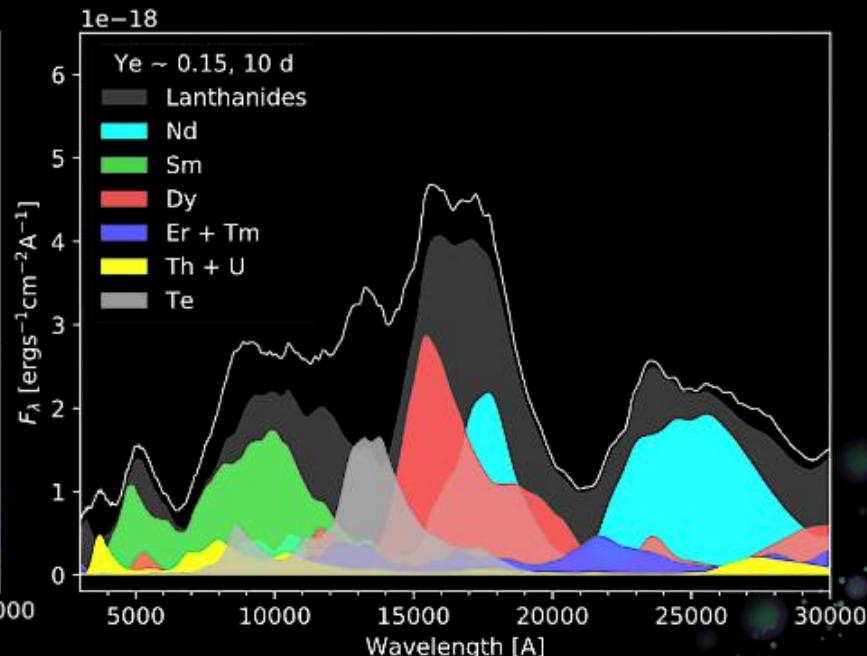
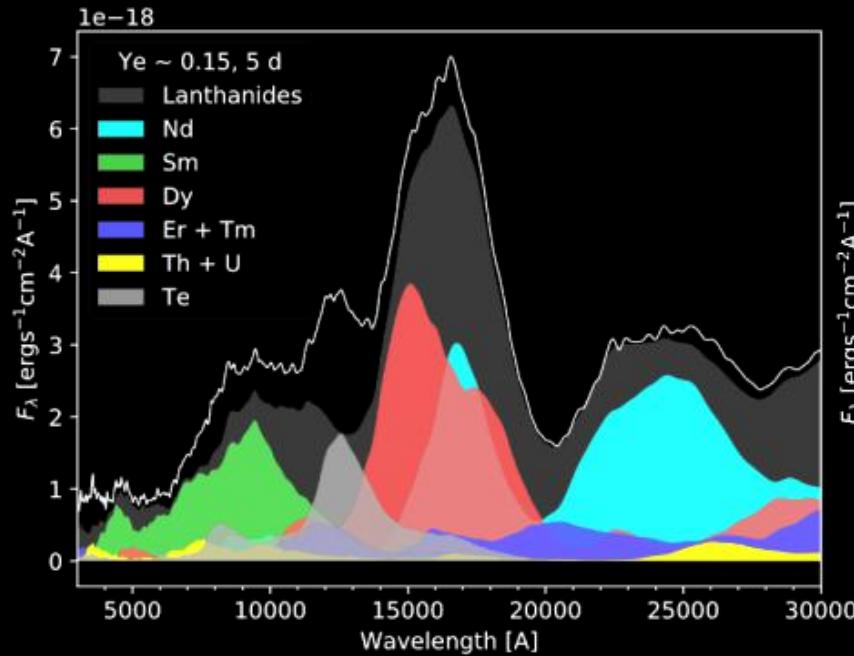
Potential
atomic lines:



SN 3D spectrum (ExTraSS)



KN 1D spectra (SUMO)



Pognan+ 2023

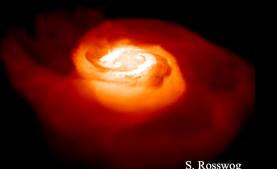
Towards KNe

What is different for KNe (vs SNe)?

- Higher velocity, lower ejecta mass
 - Very quickly diluted → NLTE effects
- Composition: heavy *r*-process elements
 - Radioactive heating: $\alpha + \beta + \gamma +$ fission
(SNe: only γ -rays of Ni-56)
 - Atoms: many levels and line rich
→ heavier transfer
 - Lines blended, no distinct lines

	SN	KN
M	$5 M_{\odot}$	$0.05 M_{\odot}$
V	$0.01c$	$0.1c$
t_{peak}	20d	2d
ρ_{peak}	10^{-11}	10^{-13}
$\frac{L(10t_{peak})}{L(t_{peak})}$	0.16	0.05
N_{lines}	$\sim 10^6$	$\sim 10^8$
% r.-a.	5%	100%

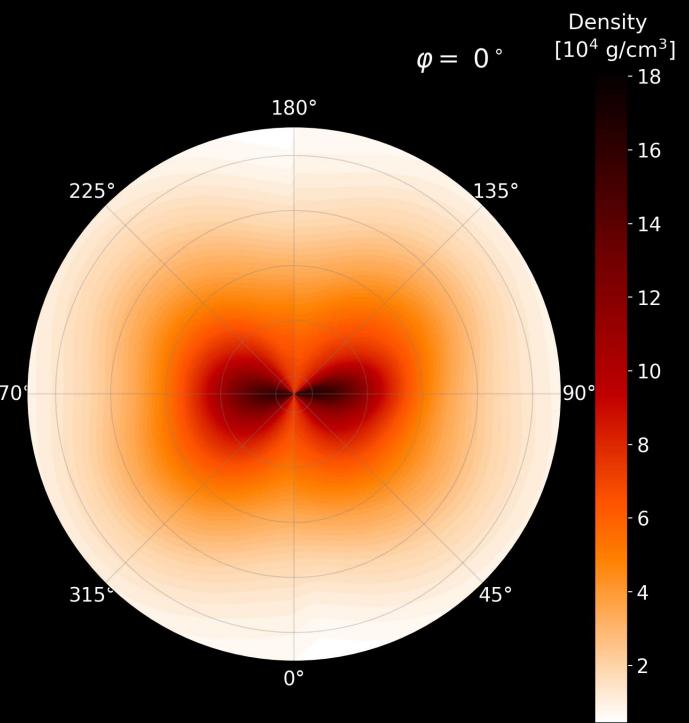
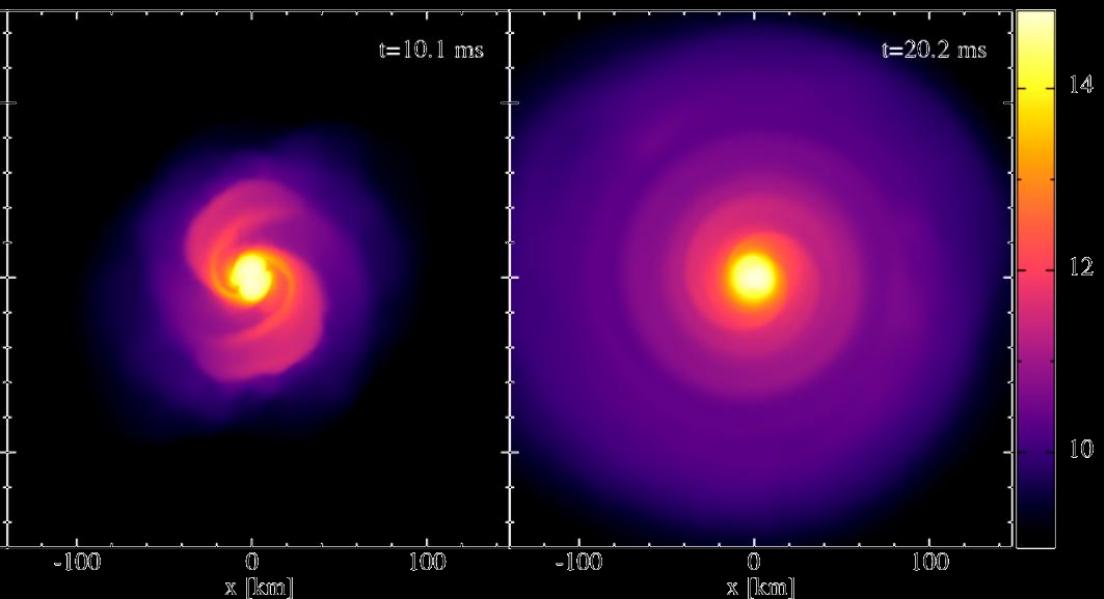
Input model



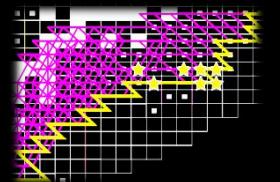
S. Rosswog

- SPHINCS_BSSN Lagrangian Numerical Relativity model
(only dynamical ejecta) *Rosswog+ 2025*
- $1.3 + 1.3 M_{\odot}$ NS; $M_{\text{ej}} 7 \cdot 10^{-4} M_{\odot}$

NSM model



Input model – composition

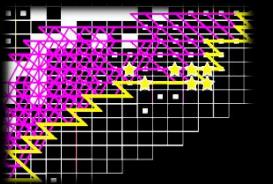


Simo

Nucl. network

1 H Hydrogen 1.008	2 Be Beryllium 9.012	3 Li Lithium 6.941	4 Mg Magnesium 24.305	5 B Boron 10.811	6 C Carbon 12.011	7 N Nitrogen 14.007	8 O Oxygen 15.999	9 F Fluorine 18.998	10 Ne Neon 20.180
11 Na Sodium 22.990	12 Mg Magnesium 24.305	3 Sc Scandium 44.956	4 Ti Titanium 47.867	5 Cr Chromium 50.942	6 Mn Manganese 54.938	7 Fe Iron 55.845	8 Co Cobalt 58.933	9 Ni Nickel 58.693	10 Cu Copper 63.546
19 K Potassium 39.098	20 Ca Calcium 40.078	21 Sc Scandium 44.956	22 V Vanadium 50.942	23 Cr Chromium 51.996	24 Mn Manganese 54.938	25 Fe Iron 55.845	26 Co Cobalt 58.933	27 Ni Nickel 58.693	28 Cu Copper 63.546
37 Rb Rubidium 85.468	38 Sr Strontium 87.62	39 Y Yttrium 88.906	40 Zr Zirconium 91.224	41 Nb Niobium 92.905	42 Mo Molybdenum 95.95	43 Tc Technetium 98.907	44 Ru Ruthenium 101.07	45 Rh Rhodium 102.906	46 Pd Palladium 106.42
55 Cs Cesium 132.905	56 Ba Barium 137.328	57-71	72 Hf Hafnium 176.49	73 Ta Tantalum 180.948	74 W Tungsten 183.84	75 Re Rhenium 186.207	76 Os Osmium 190.23	77 Ir Iridium 192.217	78 Pt Platinum 195.085
87 Fr Francium 223.020	88 Ra Radium 226.025	89-103	104 Rf Rutherfordium [261]	105 Db Dubnium [262]	106 Sg Seaborgium [266]	107 Bh Bohrium [264]	108 Hs Hassium [269]	109 Mt Meitnerium [278]	110 Ds Darmstadtium [281]
Lanthanide Series	57 La Lanthanum 138.905	58 Ce Cerium 140.116	59 Pr Praseodymium 140.908	60 Nd Neodymium 144.243	61 Pm Promethium 144.913	62 Sm Samarium 150.36	63 Eu Europium 151.964	64 Gd Gadolinium 157.25	65 Tb Terbium 158.925
Actinide Series	89 Ac Actinium 227.028	90 Th Thorium 232.038	91 Pa Protactinium 231.036	92 U Uranium 238.029	93 Np Neptunium 237.048	94 Pu Plutonium 244.064	95 Am Americium 243.061	96 Cm Curium 247.070	97 Bk Berkelium 247.070
	98 Cf Californium 251.080	99 Es Einsteinium [254]	100 Fm Fermium 257.095	101 Md Mendelevium 258.1	102 No Nobelium 259.101	103 Lr Lawrencium [262]			

Input model – composition

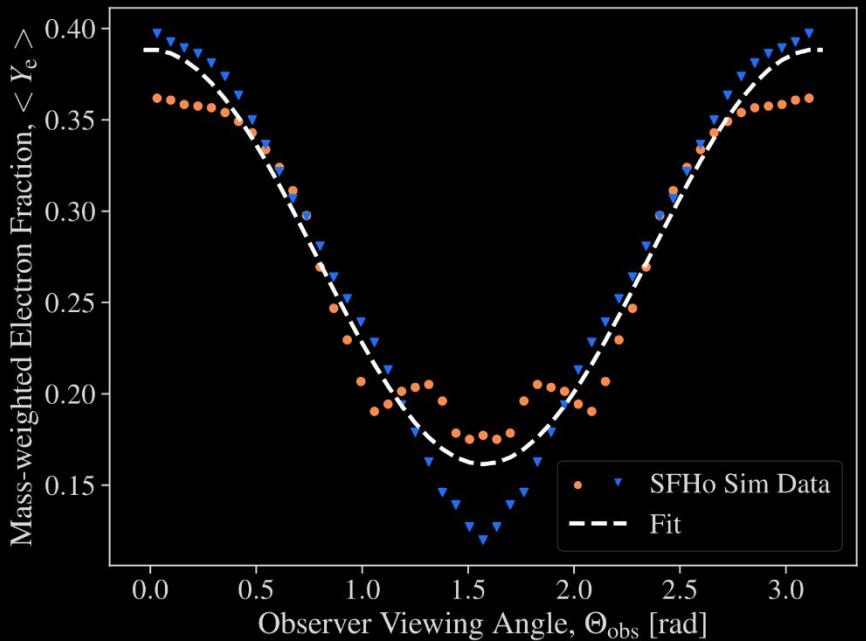


Simo

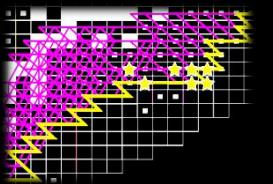
Nucl. network

$n_{e,12}$

- Y_e distribution: *Setzer+ 2023*



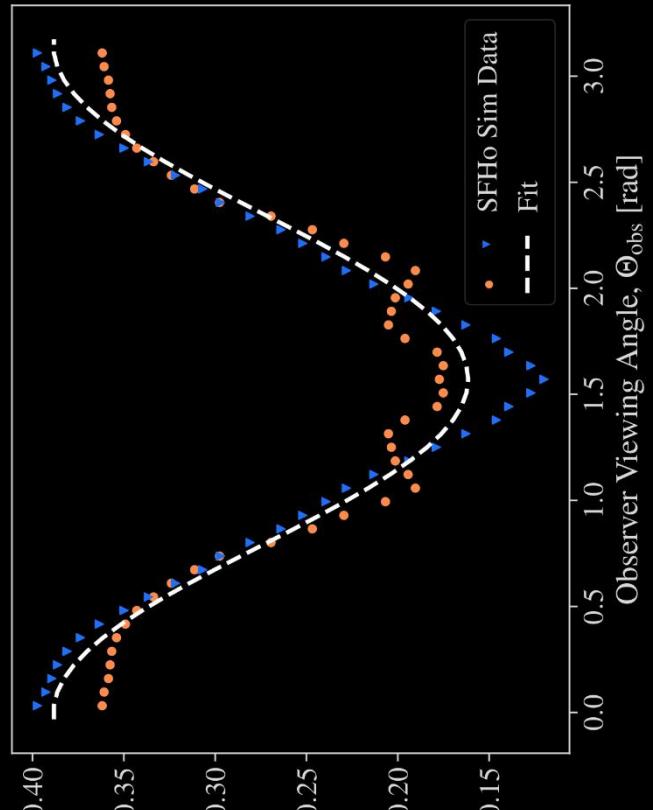
Input model - composition



Simo

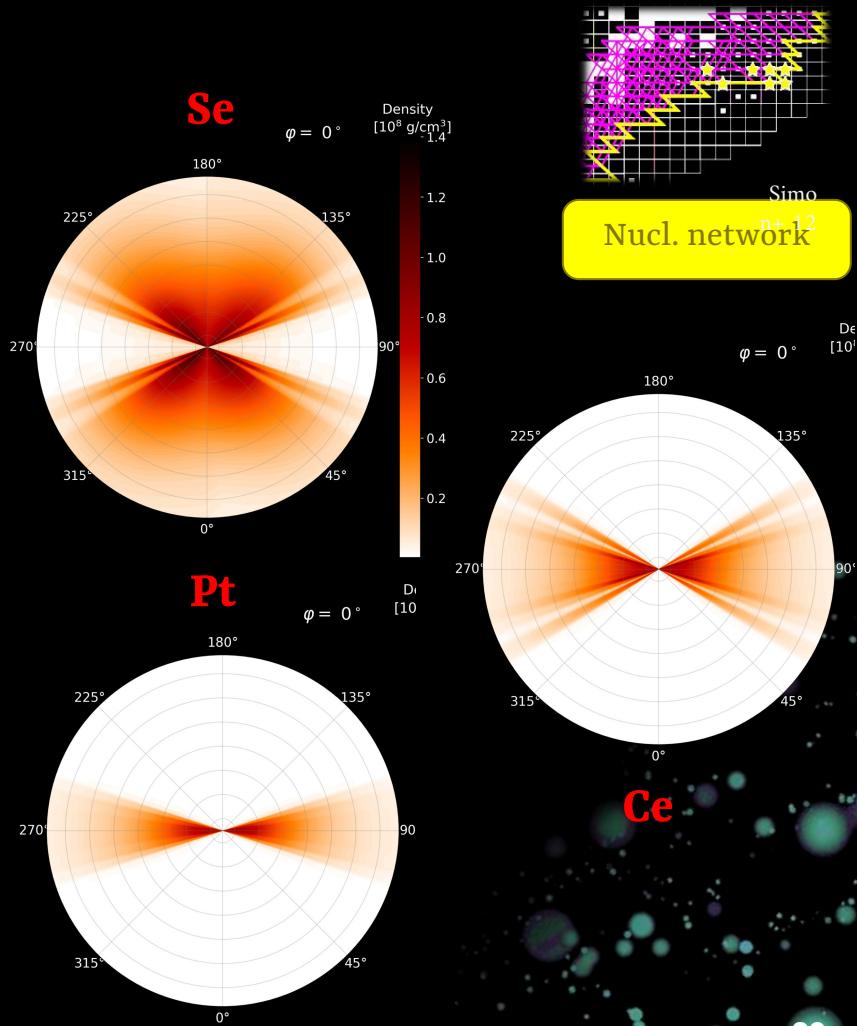
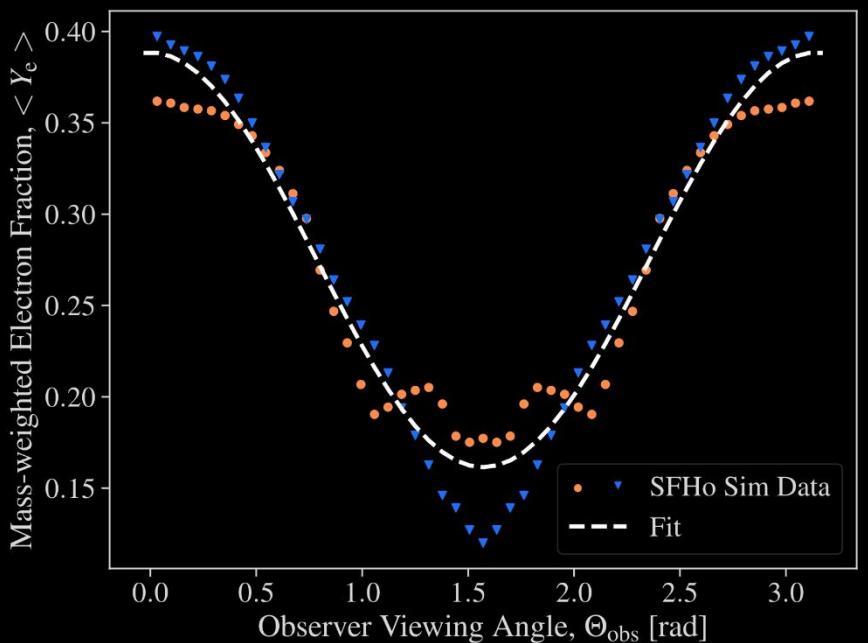
Nucl. networkⁿ⁼¹²

- Y_e distribution: *Setzer+ 2023*



Input model – composition

- Y_e distribution: *Setzer+ 2023*
- $Y_e \rightarrow$ composition, radioactive heating
Wanajo+ 2018



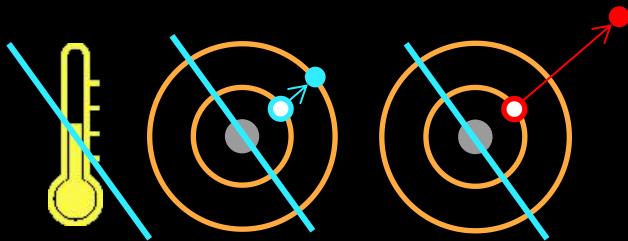


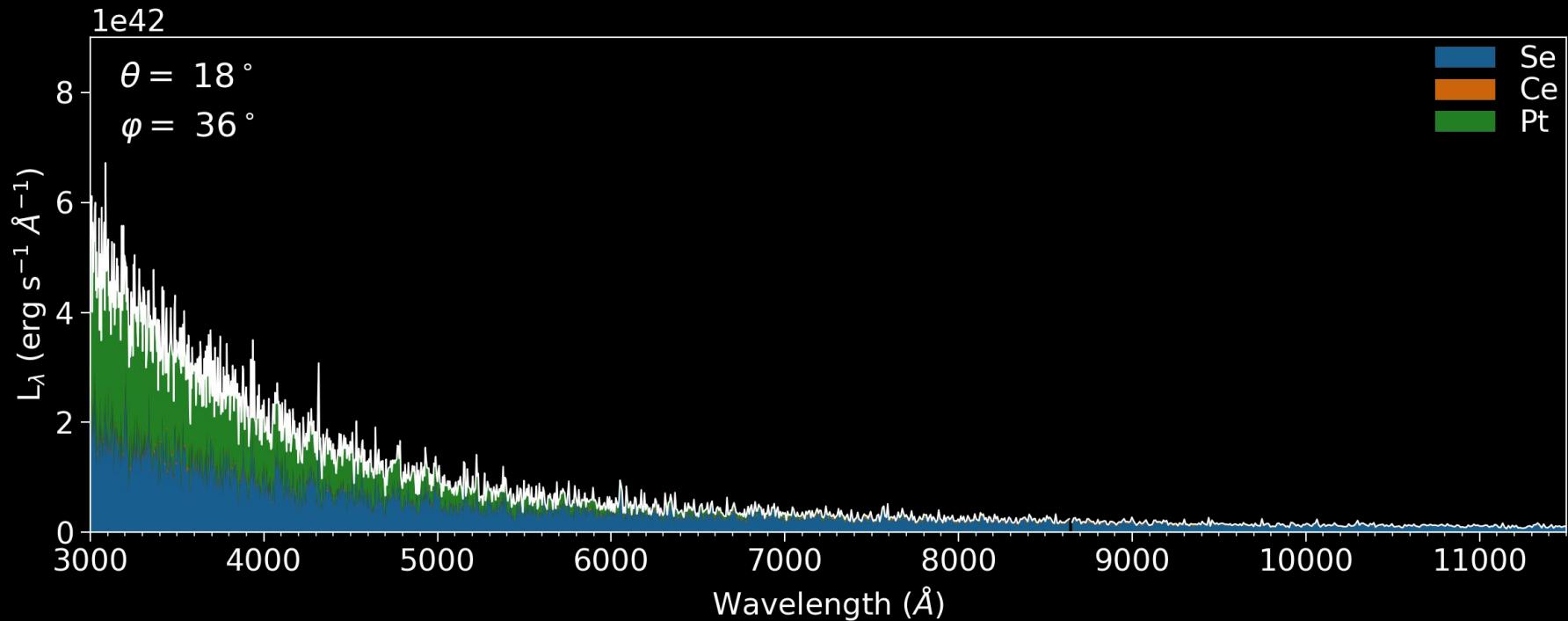
Preliminary spectra

Spectra so far

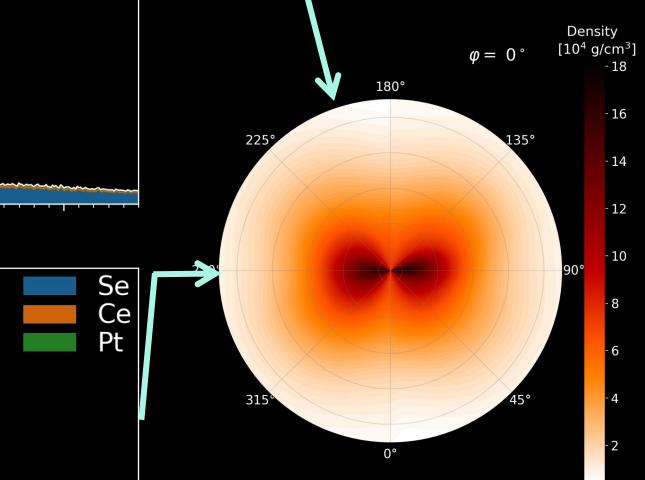
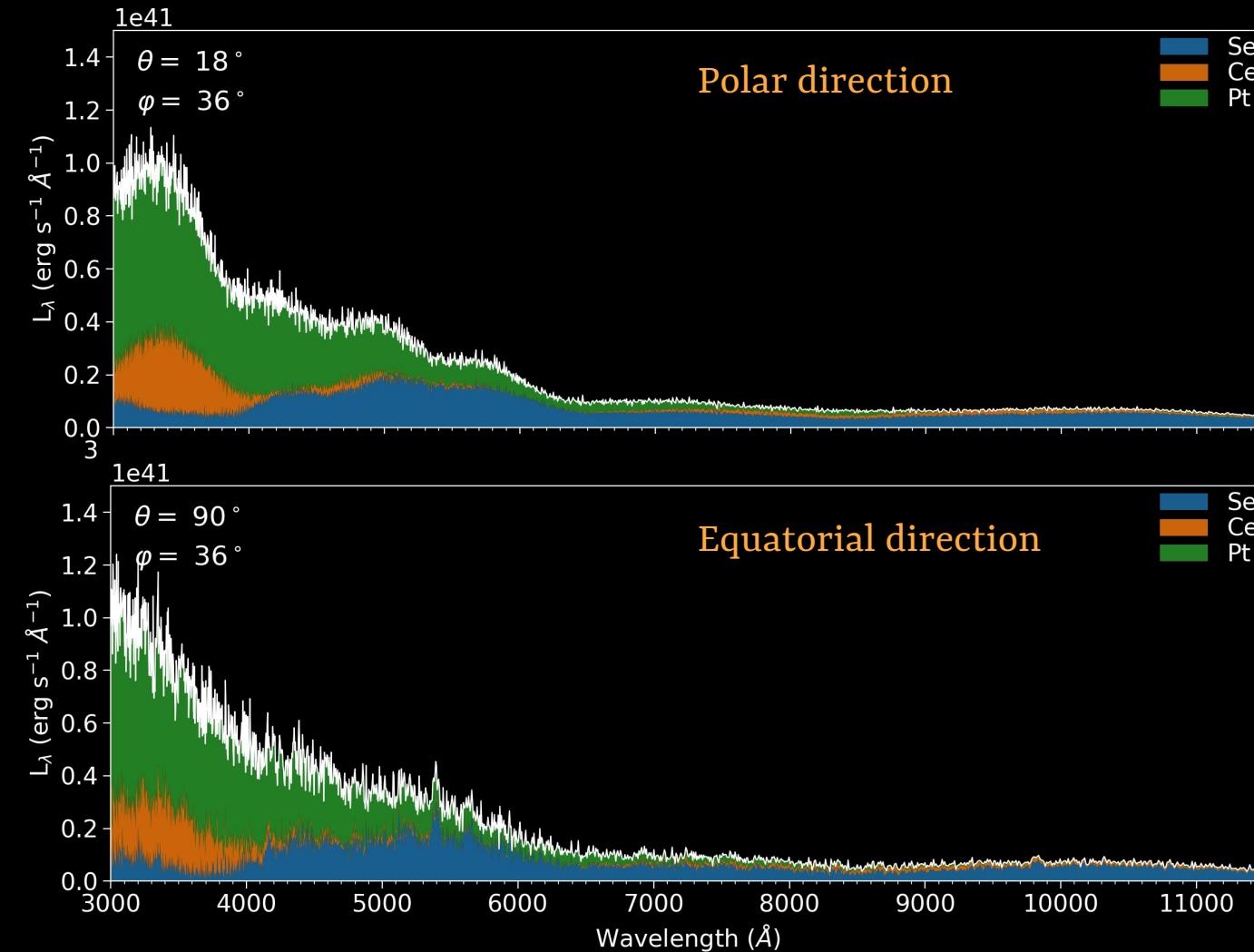
Spectral modelling of kilonovae in 3D NLTE

- Locked physical conditions:
 - T fixed (5000 K)
 - Excitation: LTE (Boltzmann)
 - Ionisation: fixed
- What makes different spectra for diff. viewing angles?
 - Doppler factors
 - Composition





Velocity scaled
down by 10



Future plans

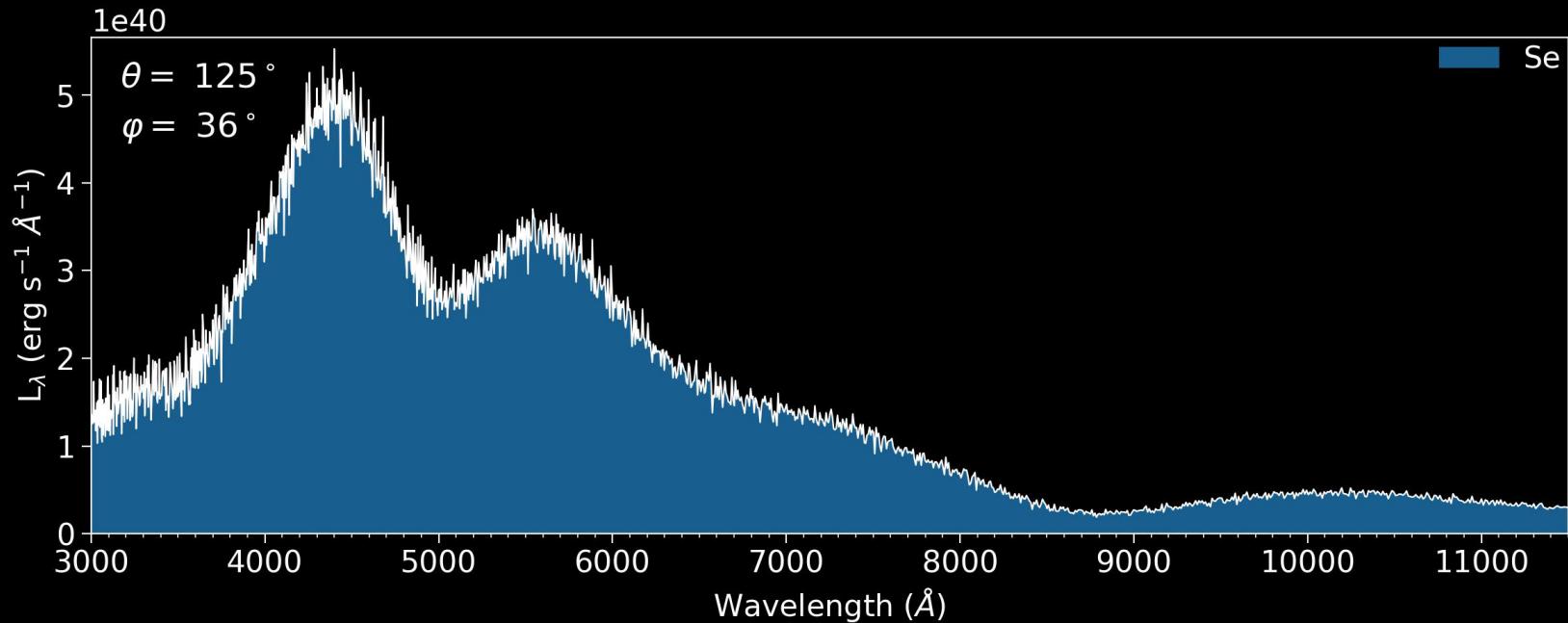
- NLTE solver
- Include multiple r-process elements
- Time dependence, light curves
- Different input models



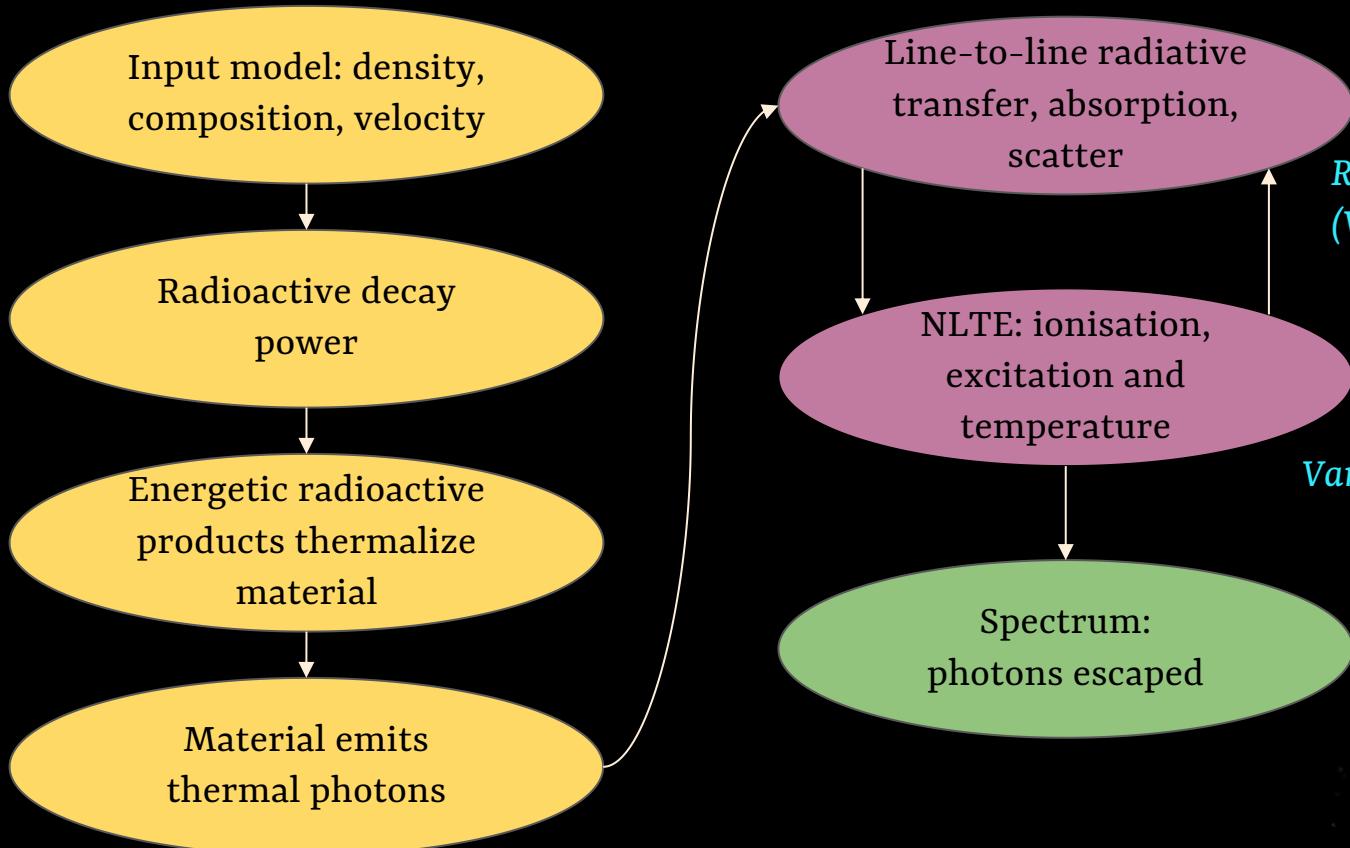
3D NLTE is
important!

Toy model

- Homogenous sphere (1D), $0.1 M_{\odot}$, $v = 0.1 c$, $t = 2d$
- Pure Se



Structure of the code



Recently completed
(Van Baal+ in prep)

Van Baal+ 2023, 2024

Stockholm codes for spectral modeling

- Photons **Doppler shifted** in expanding material
(comoving frame treatment)
- Photon transfer by **Monte Carlo** random sampling
- **Line-to-line radiative transfer**
 - No assumptions of optical thinness or opacity