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Tracing r-process nucleosynthesis in neutron star mergers with long-lived remnants

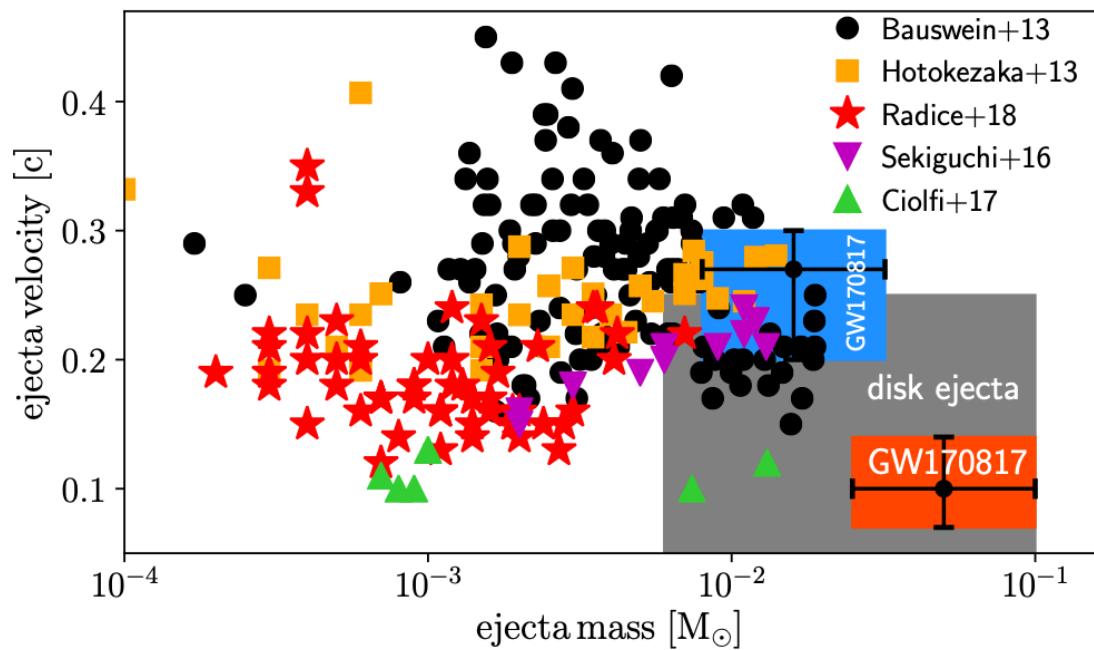
Eleonora Loffredo

INAF OAAB & INFN LNGS

S. Cristallo, F. M. Guercilena, A. Perego, D. Vescovi

BNS & Kilonovae: some challenges

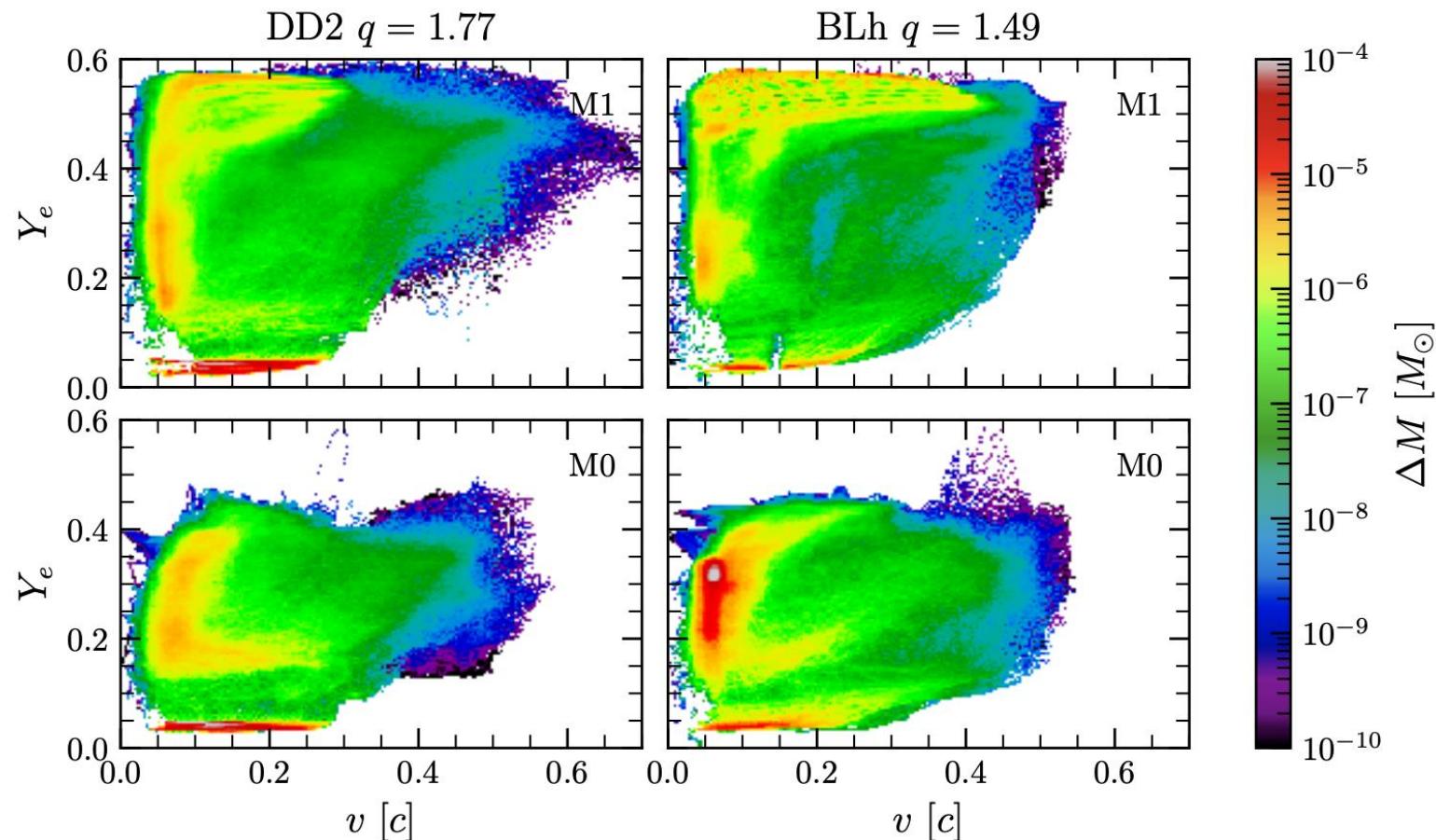
- BNS contribution to r-process
- AT2017gfo: ejecta, nucleosynthesis, spectra
- Kilonovae & long GRBs [Rastinejad+22; Troja+22; Mei+22; Levan+23]
- BNS merger population [Santoliquido+21; Mandel+22; Iorio+23]



Credits: Siegel, EPJA, 2019

Impact of EOS and neutrino treatment

Credits: Bernuzzi et al, 2024



E.g. Foucart+16,20; Radice+22; Kiuchi+23; Just+23; Riciglano+24; Bernuzzi+24; Sneppen+24; Jacobi+25



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Aims

- Investigate impact of neutrino winds on the nucleosynthesis
- Produce publicly available database of r-process yields varying BNS mass ratio and EOS



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Method

- Set of BNS merger simulations with long-lived remnants
- Neutrino treatment: M1 gray scheme [Radice+22]
- Extract tracer particles for ejecta
- Compute nucleosynthesis with WinNet [Reichert+23]



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Method

- Set of BNS merger simulations with long-lived remnants → SFHo EOS and unitary mass ratio
- Neutrino treatment: M1 grey scheme [Radice+22]
- Extract tracer particles for ejecta
- Compute nucleosynthesis with WinNet [Reichert+23]



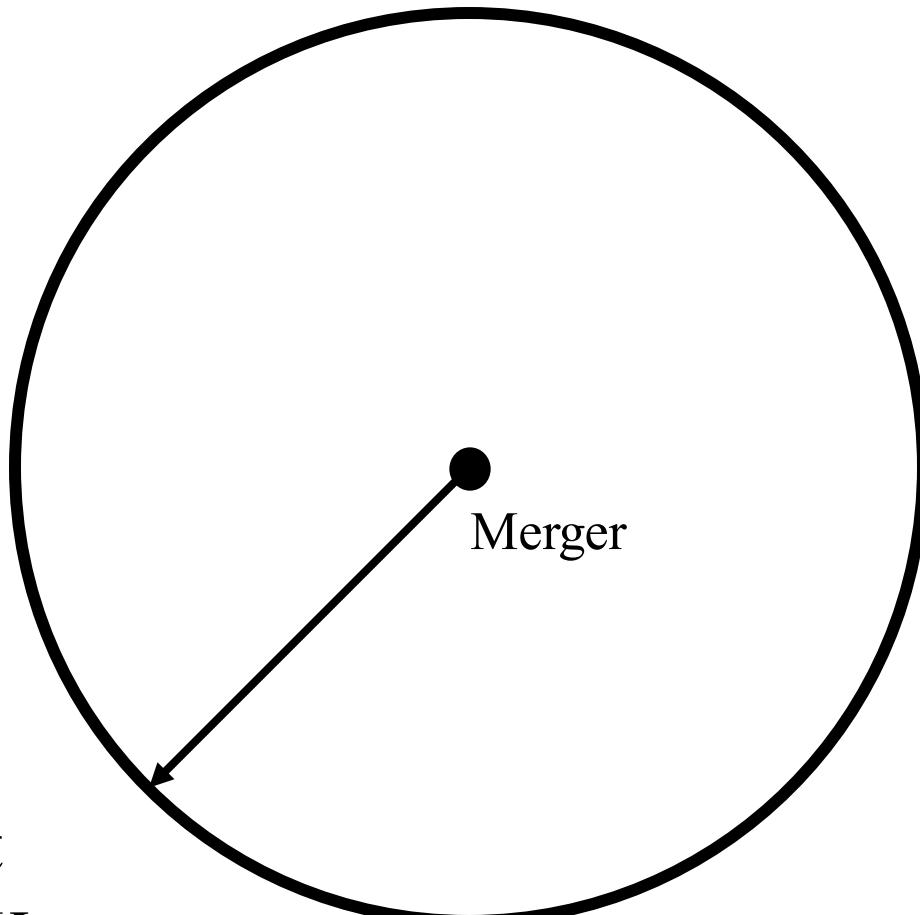
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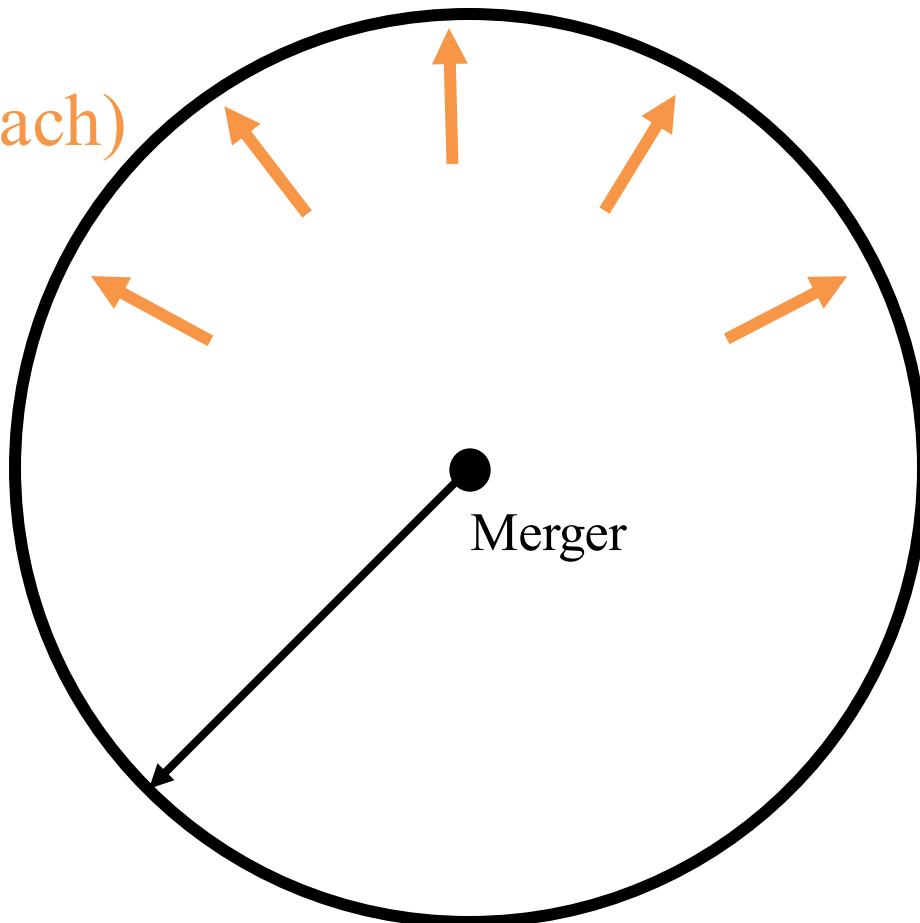


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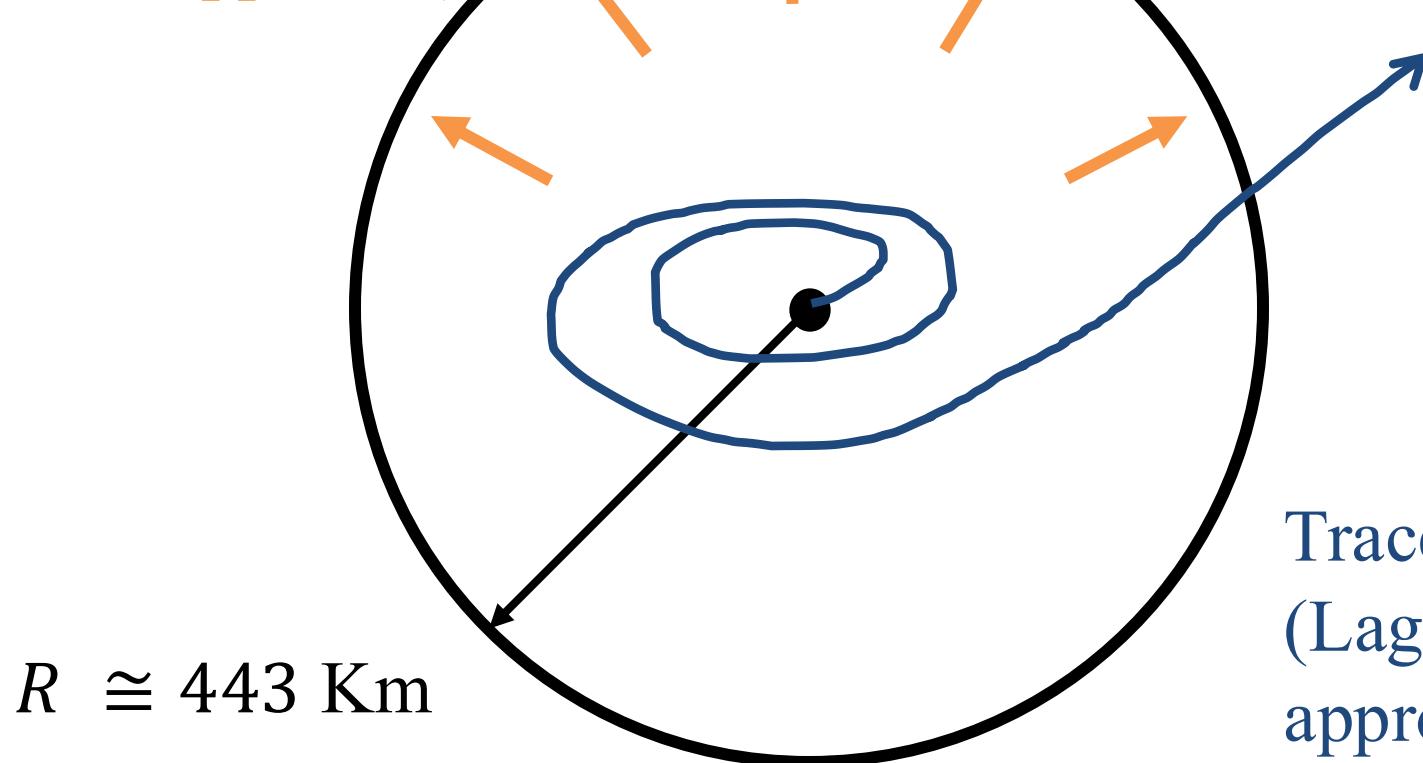


Detector at
 $R \cong 443$ Km

Flux of ejected matter (Eulerian approach)



Flux of ejected
matter
(Eulerian approach)



Tracer particles
(Lagrangian
approach)



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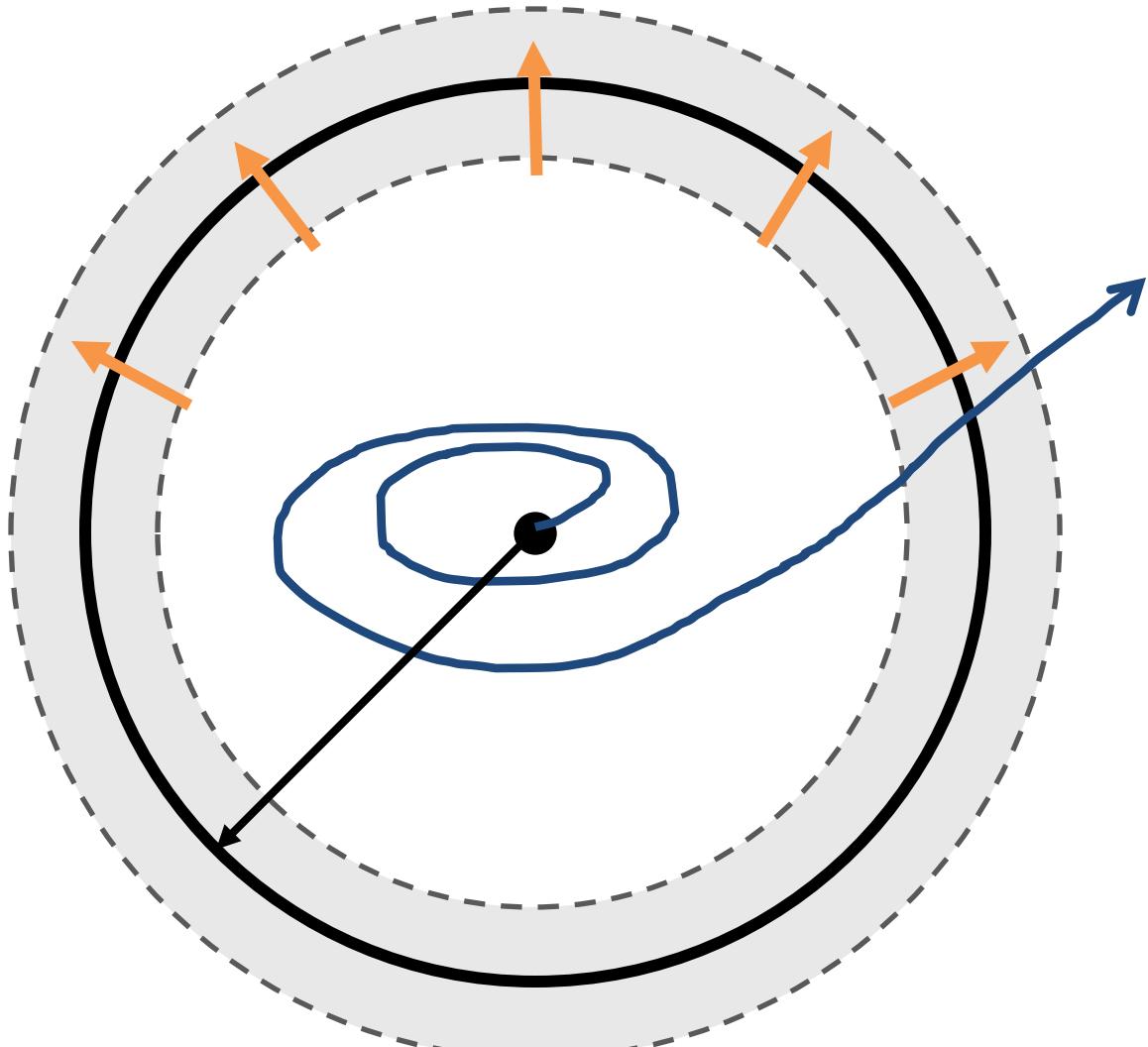
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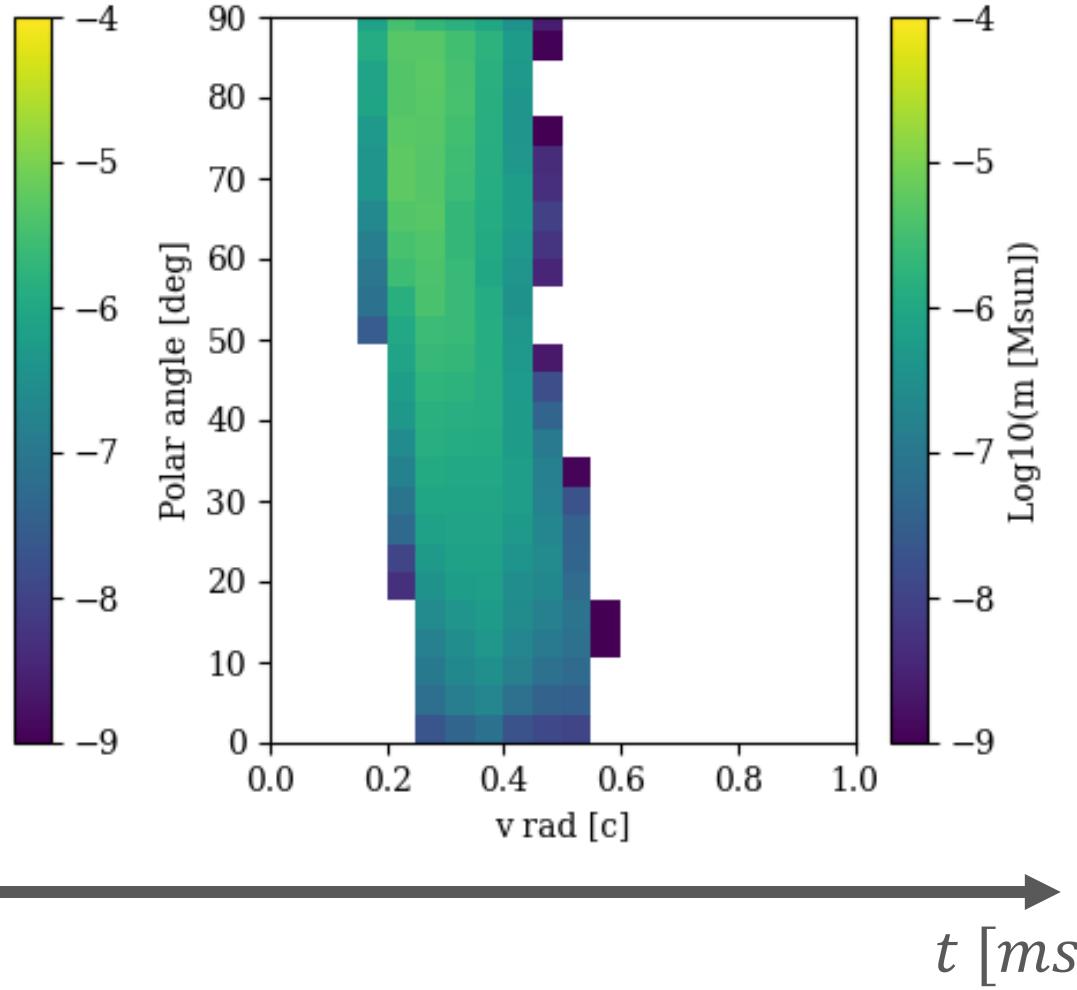
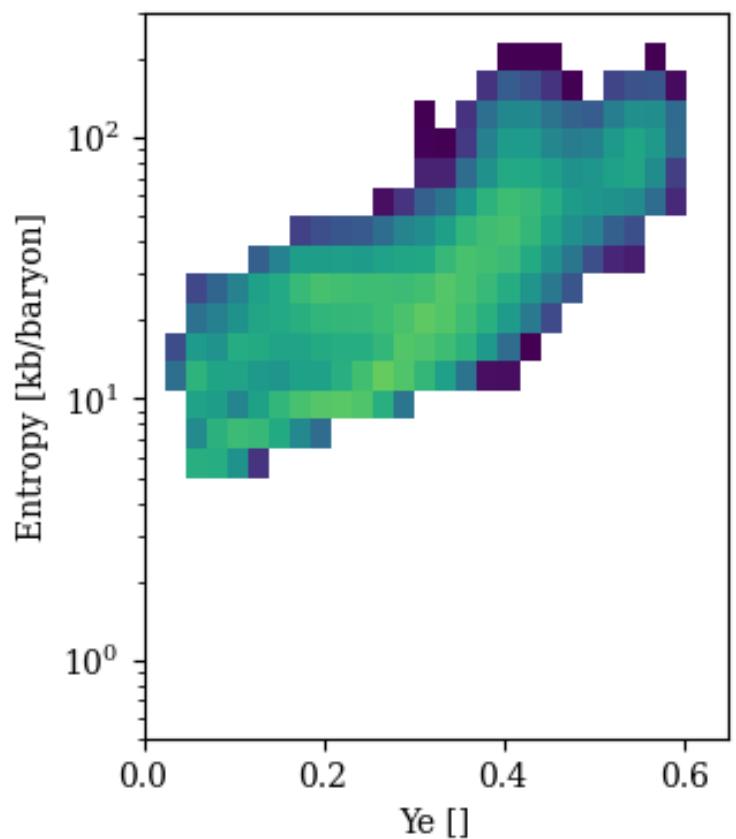
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Time
windows
 $\Delta t_1, \dots, \Delta t_n$



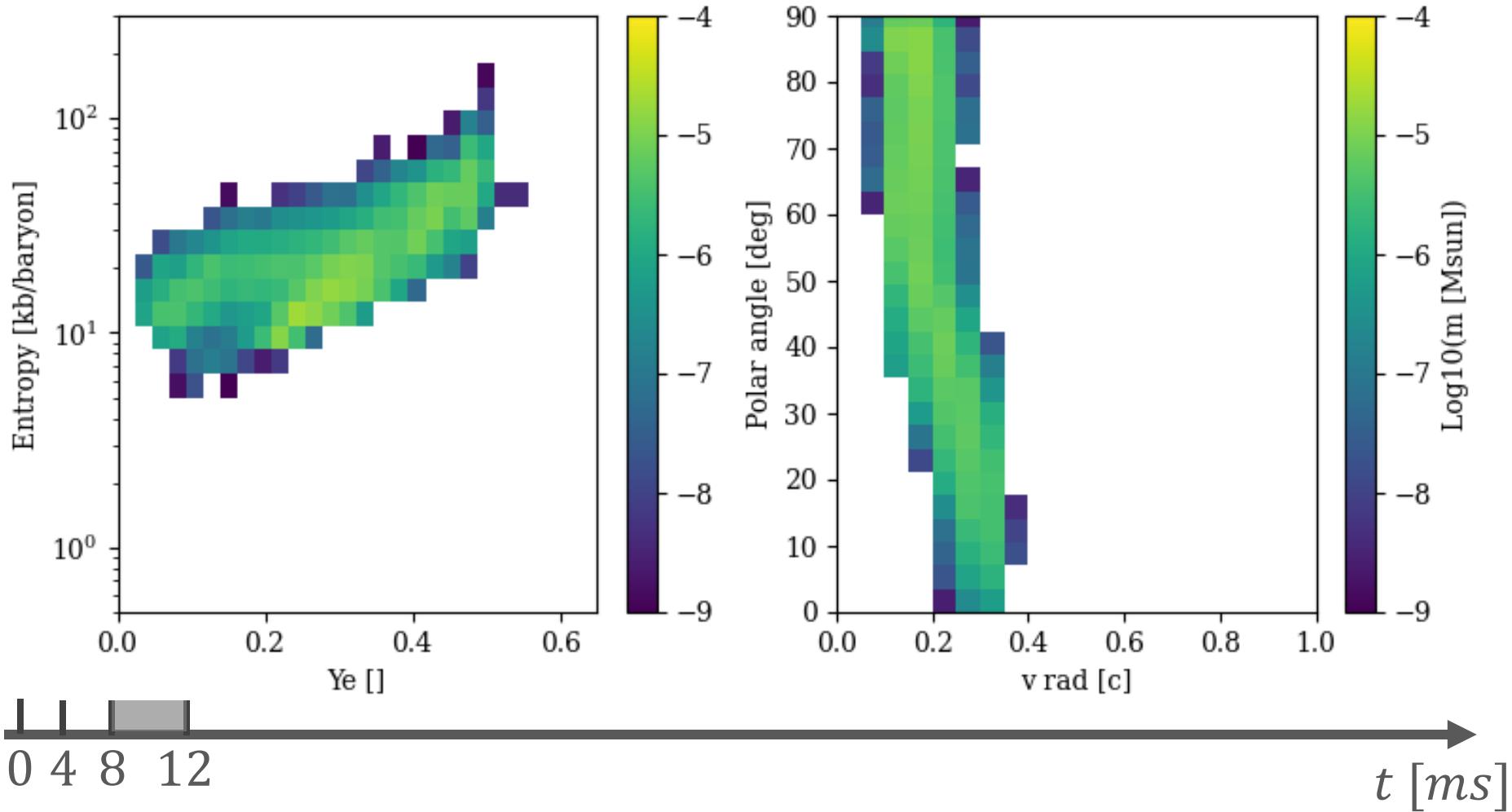
Ejecta analysis



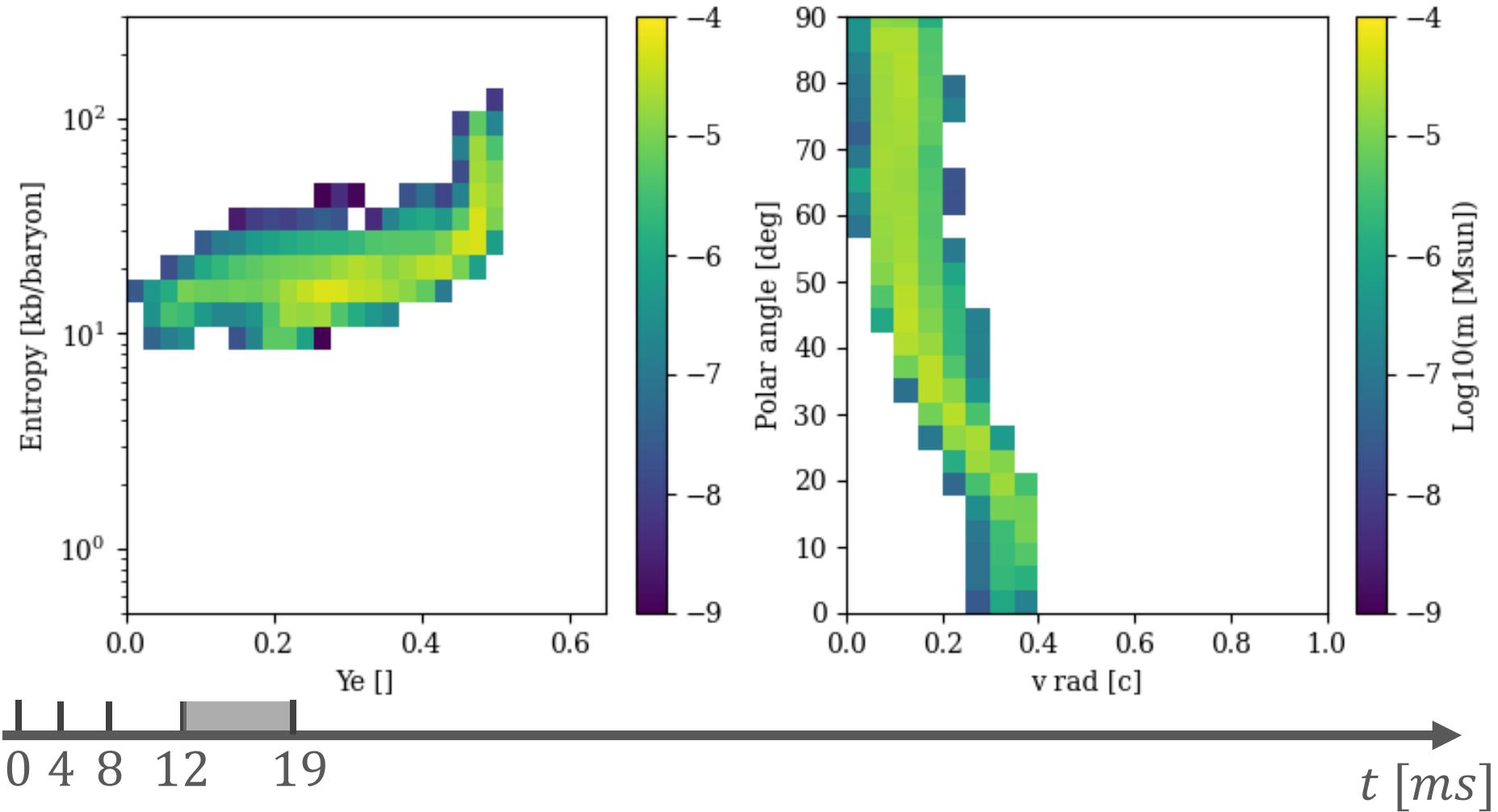
0 4 8

t [ms]

Ejecta analysis



Ejecta analysis





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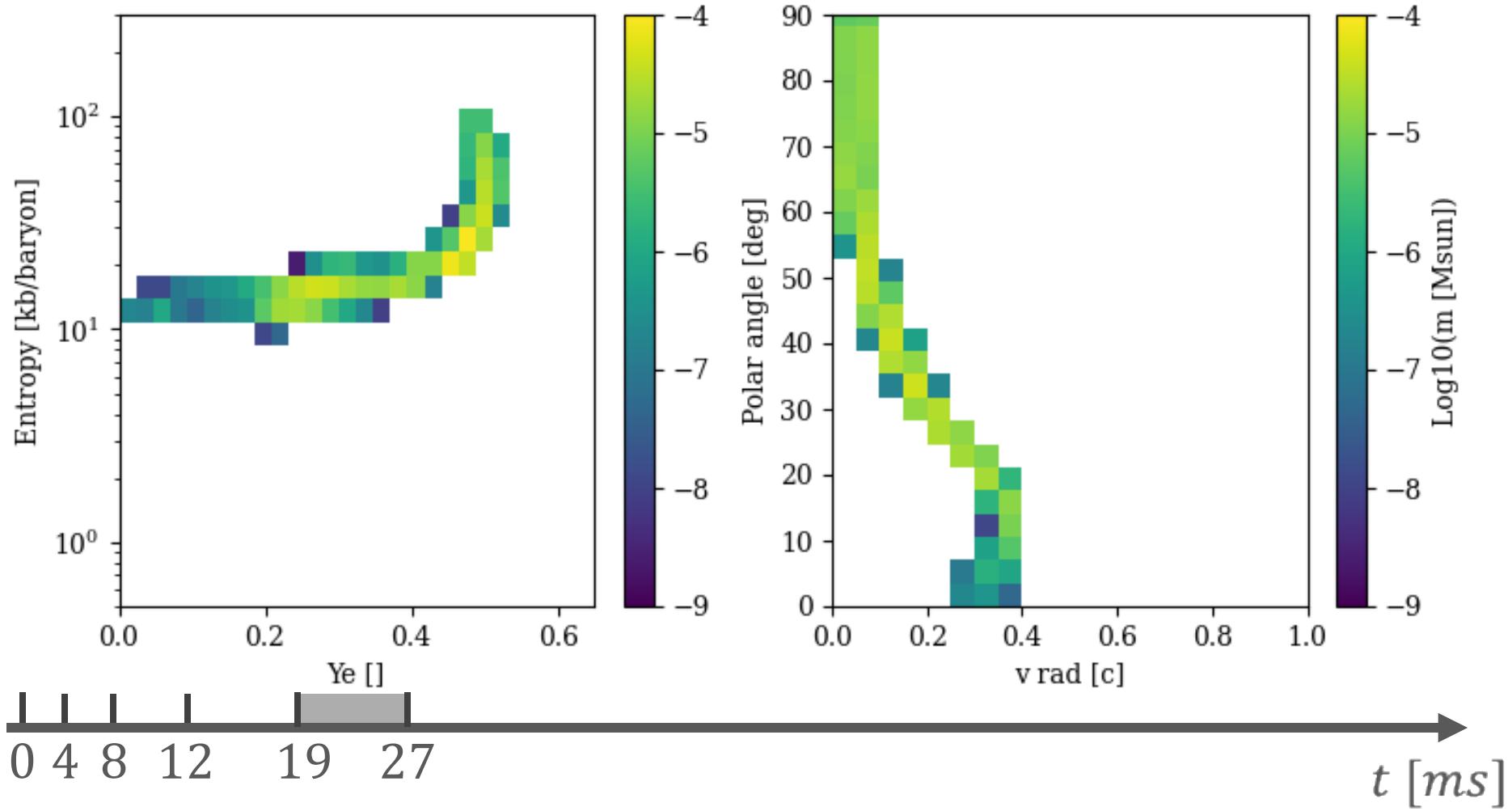
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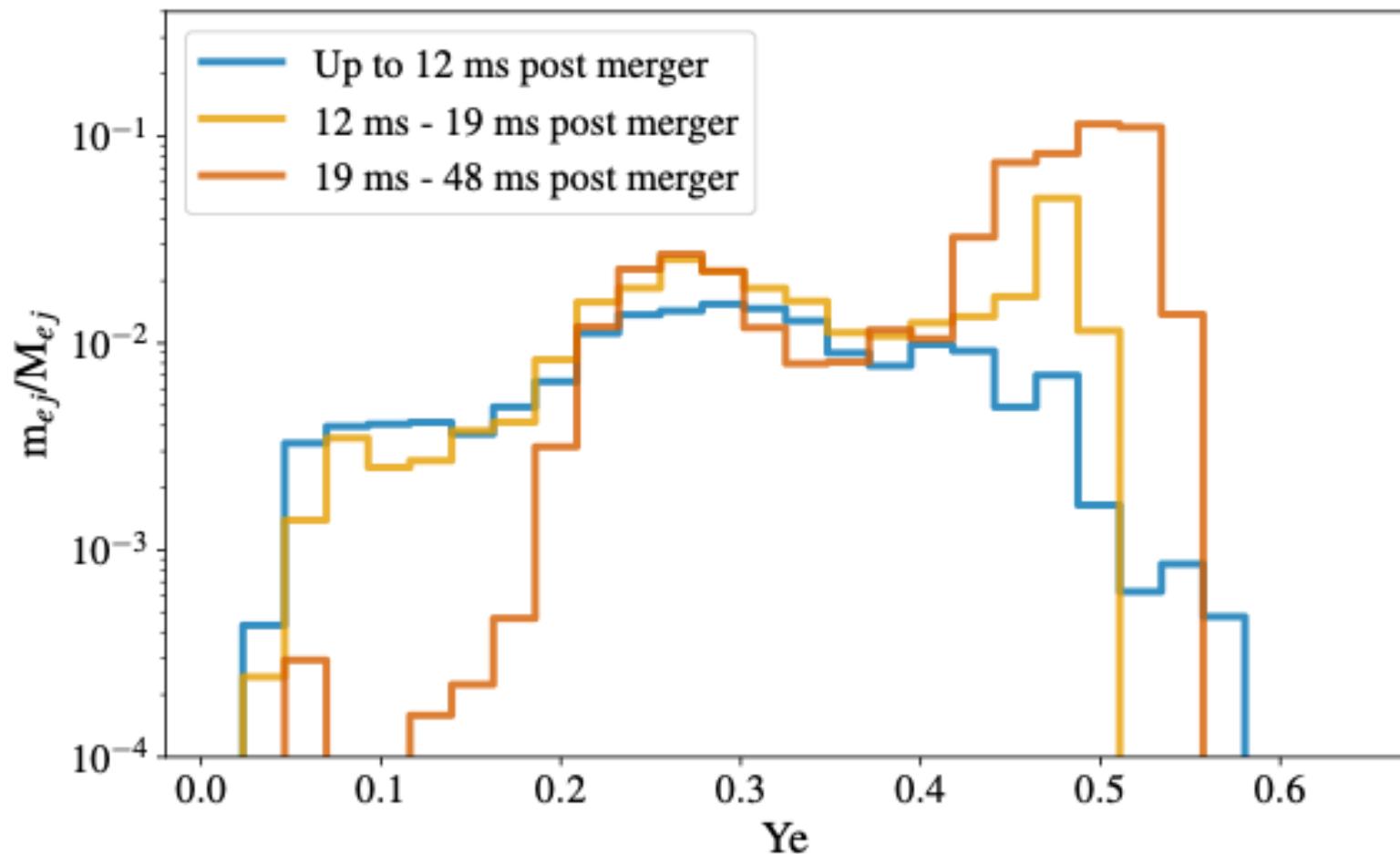
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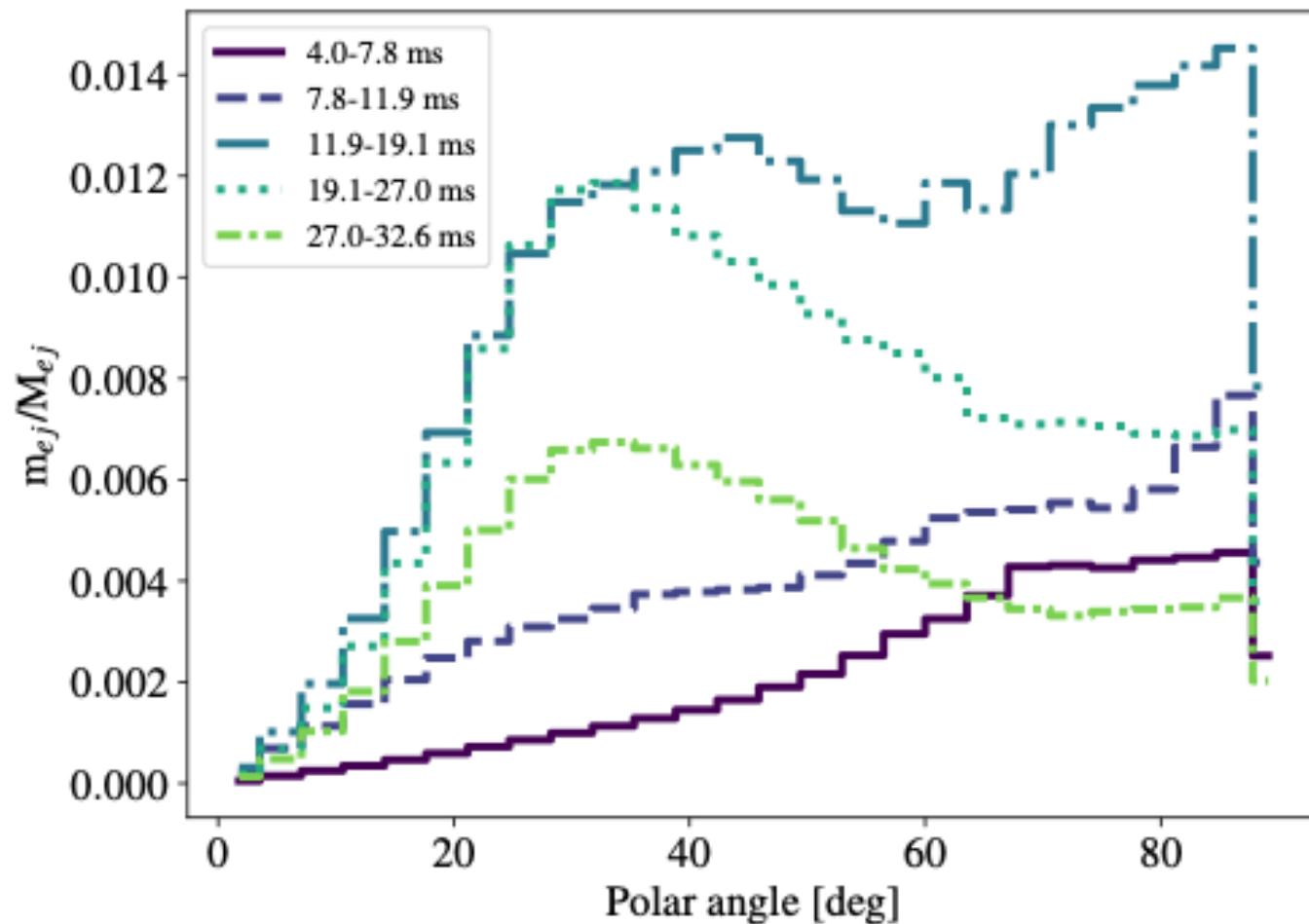
Ejecta analysis



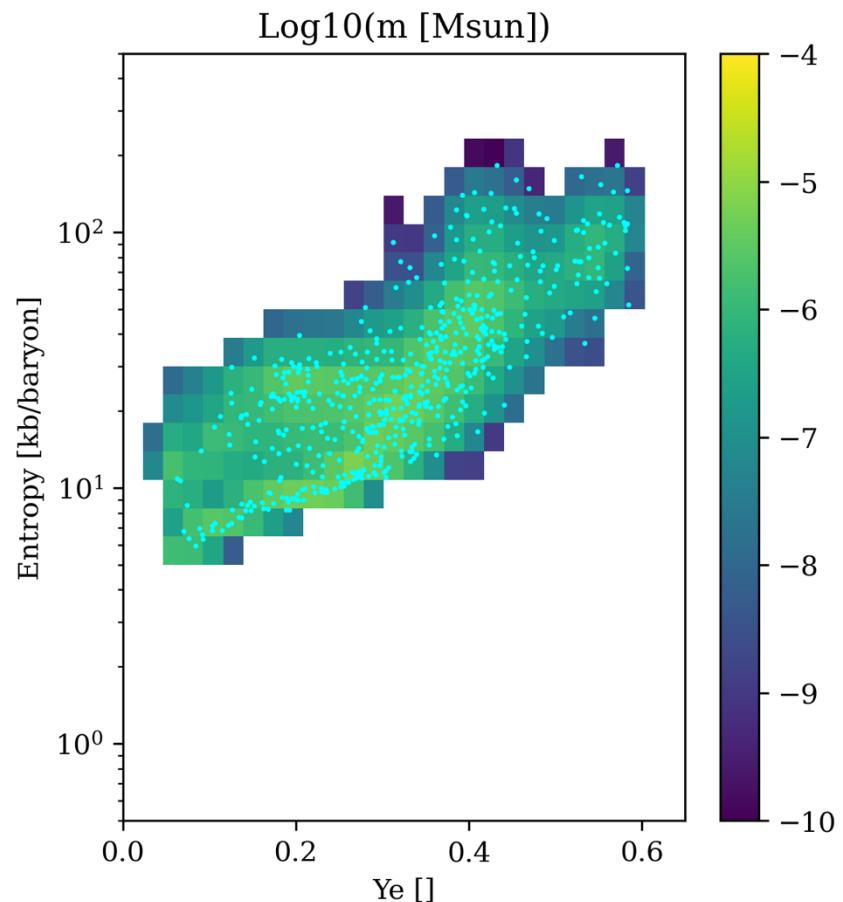
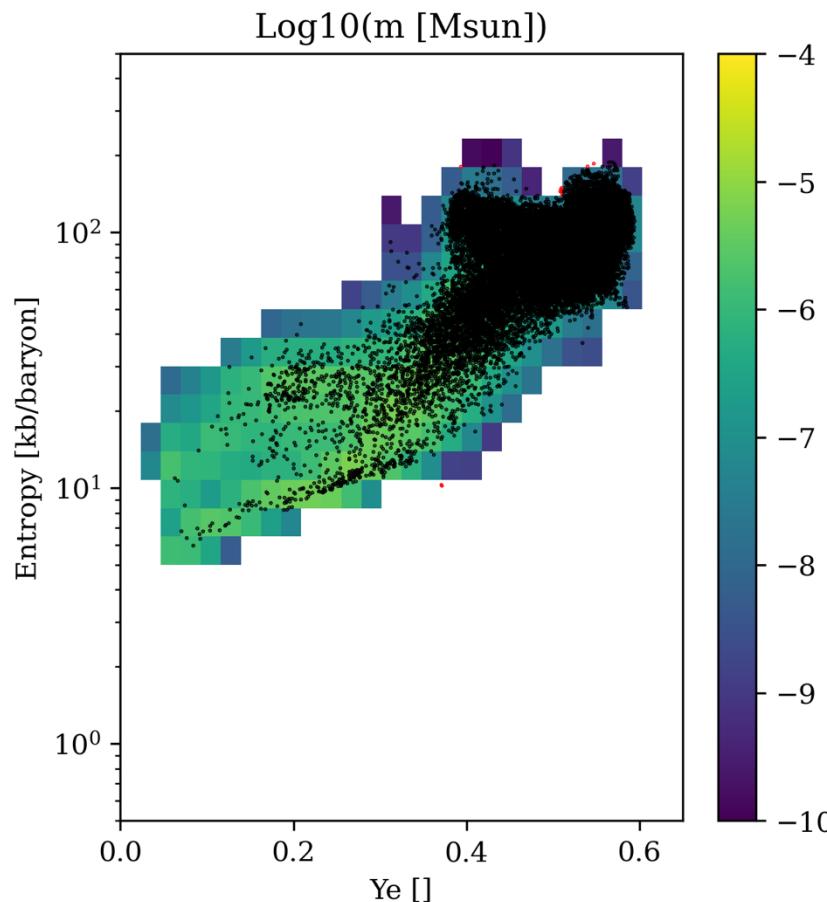
Ejecta analysis: Ye distribution



Ejecta analysis: angular distribution

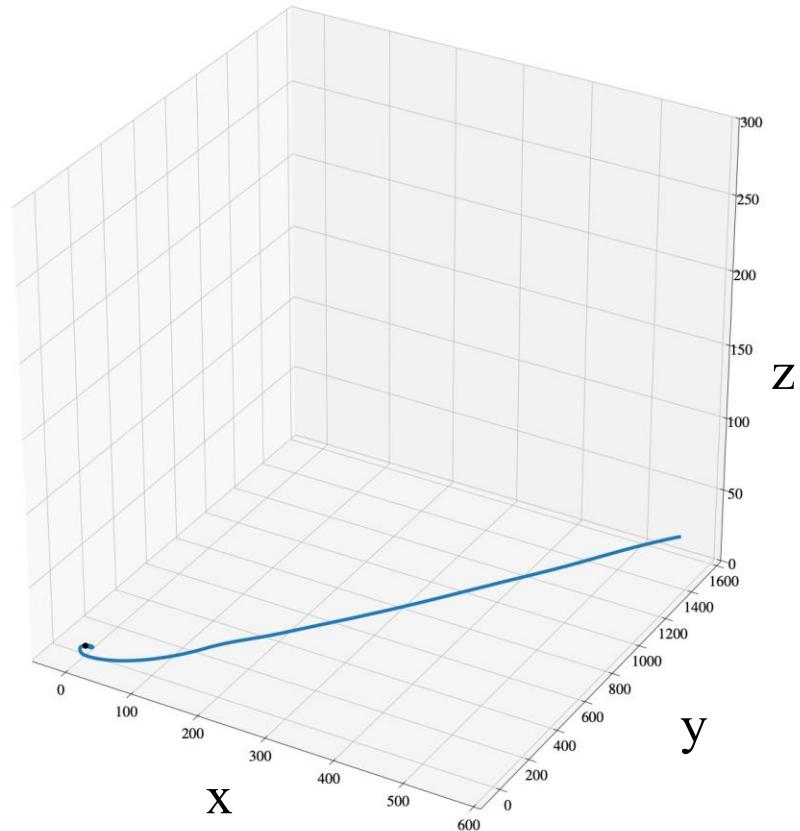
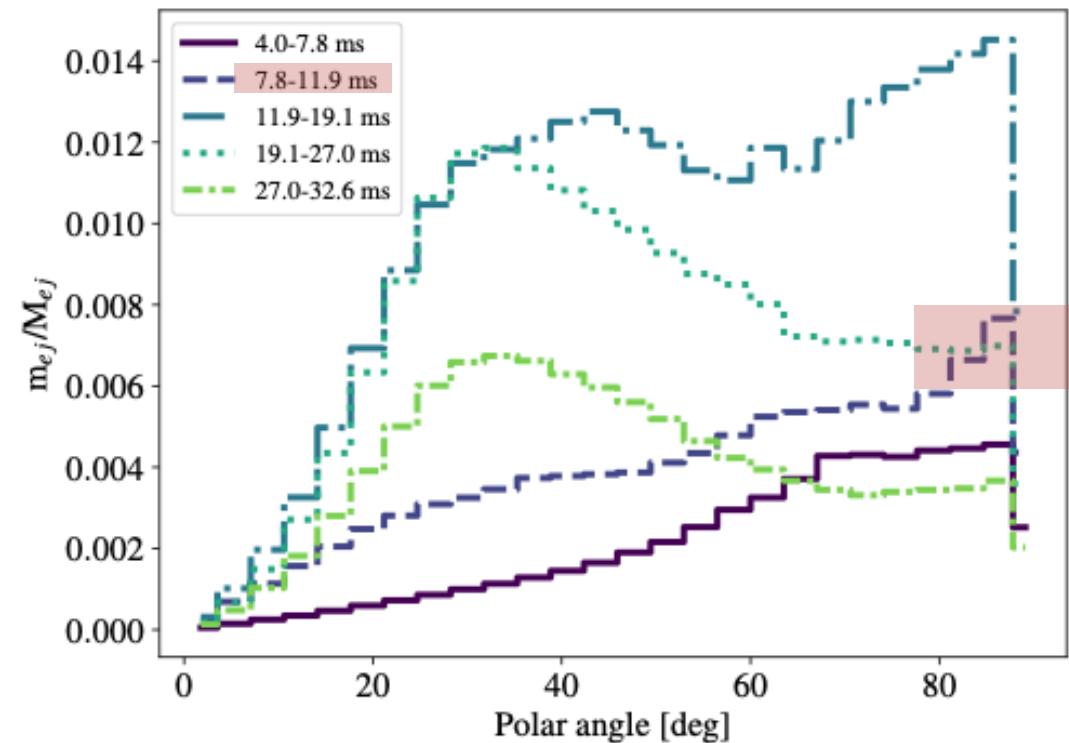


Tracer extraction



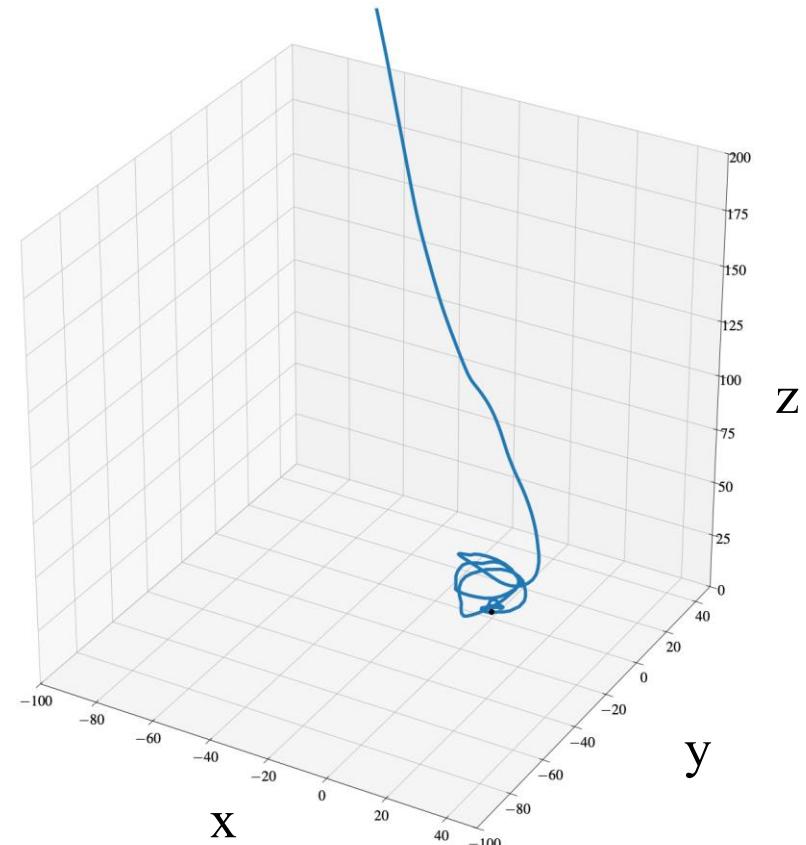
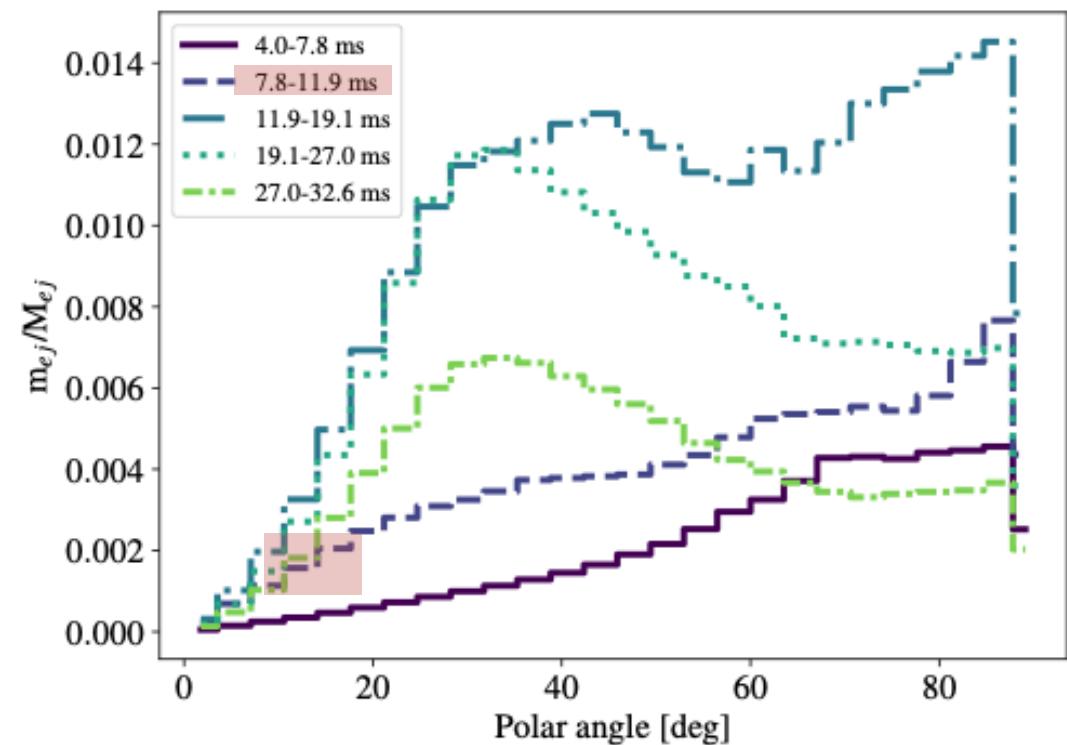
From 2.6×10^5 to 3×10^3 tracers: mass weighted extraction per time window

$$Y_e(t_{merger}) = 0.10$$



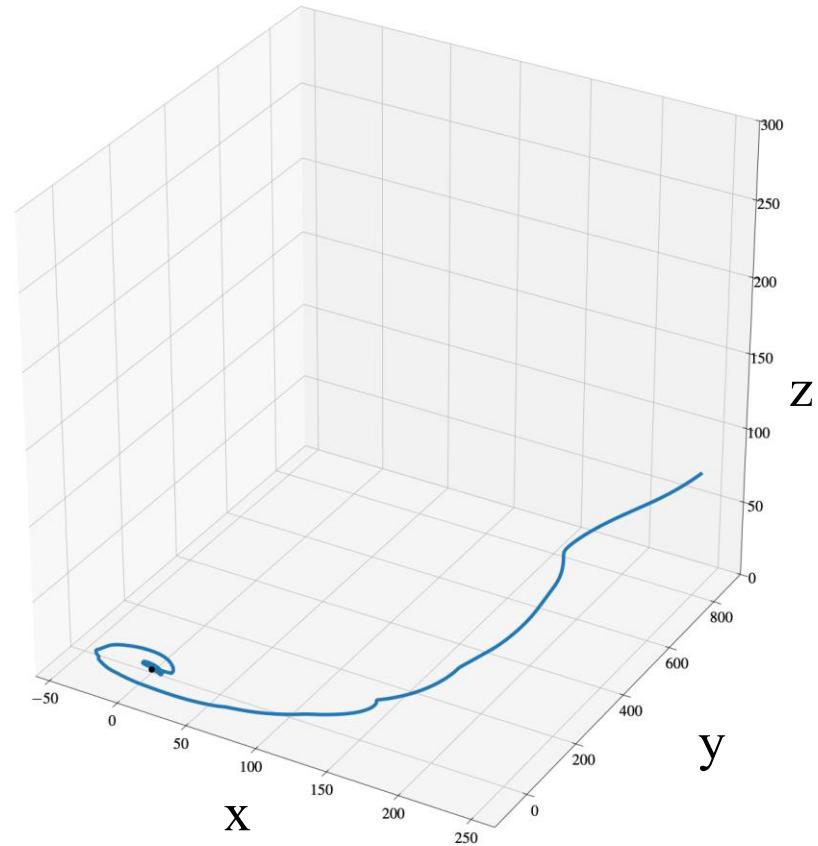
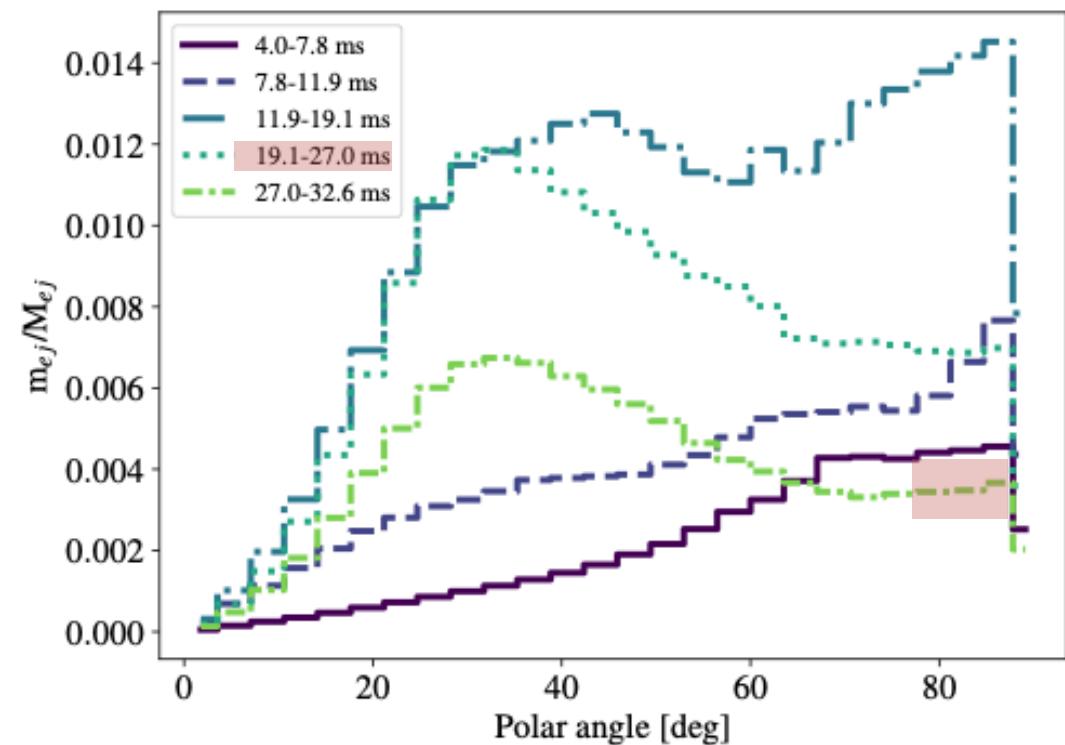
$$Y_e(t_{fin}) = 0.06$$

$$Y_e(t_{merger}) = 0.10$$



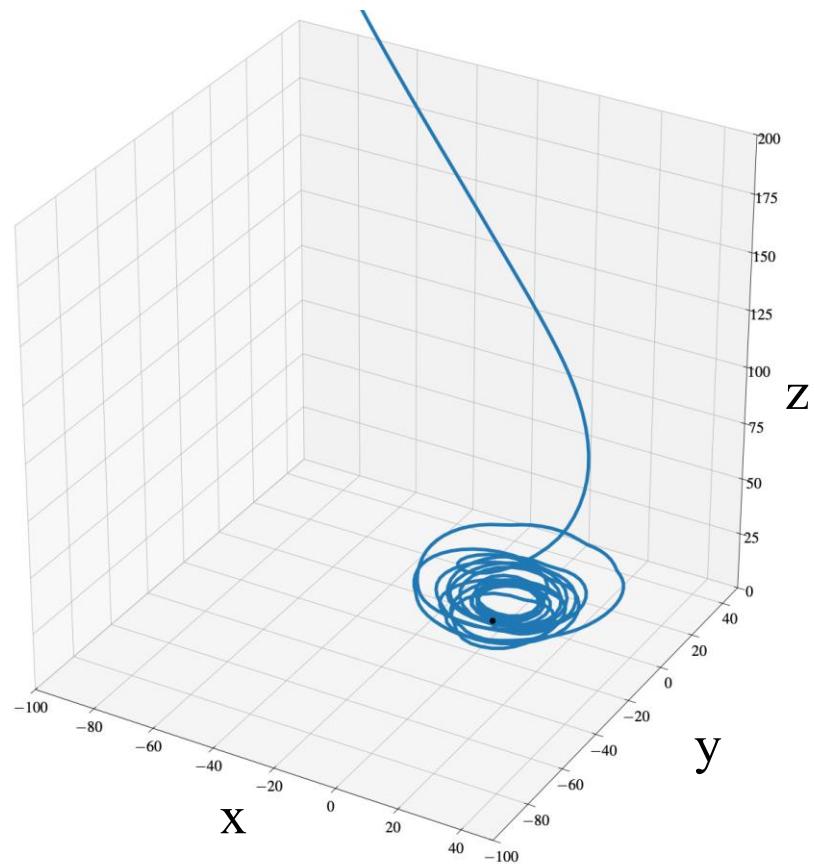
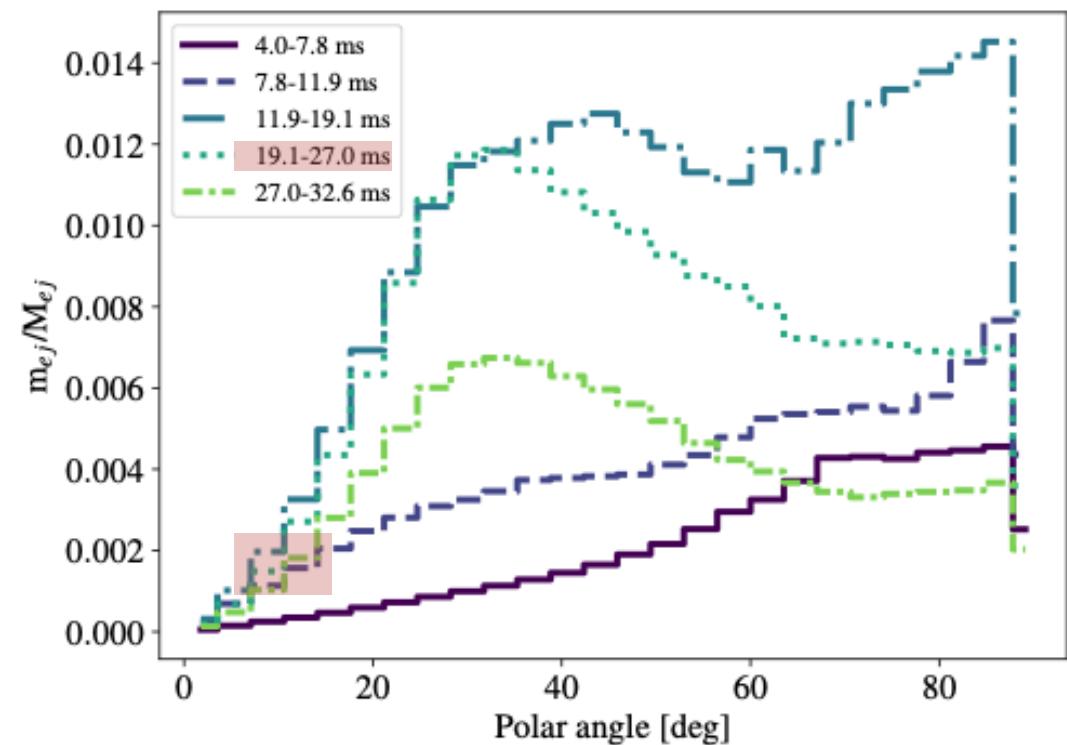
$$Y_e(t_{fin}) = 0.49$$

$$Y_e(t_{merger}) = 0.10$$



$$Y_e(t_{fin}) = 0.29$$

$$Y_e(t_{merger}) = 0.10$$



$$Y_e(t_{fin}) = 0.5$$



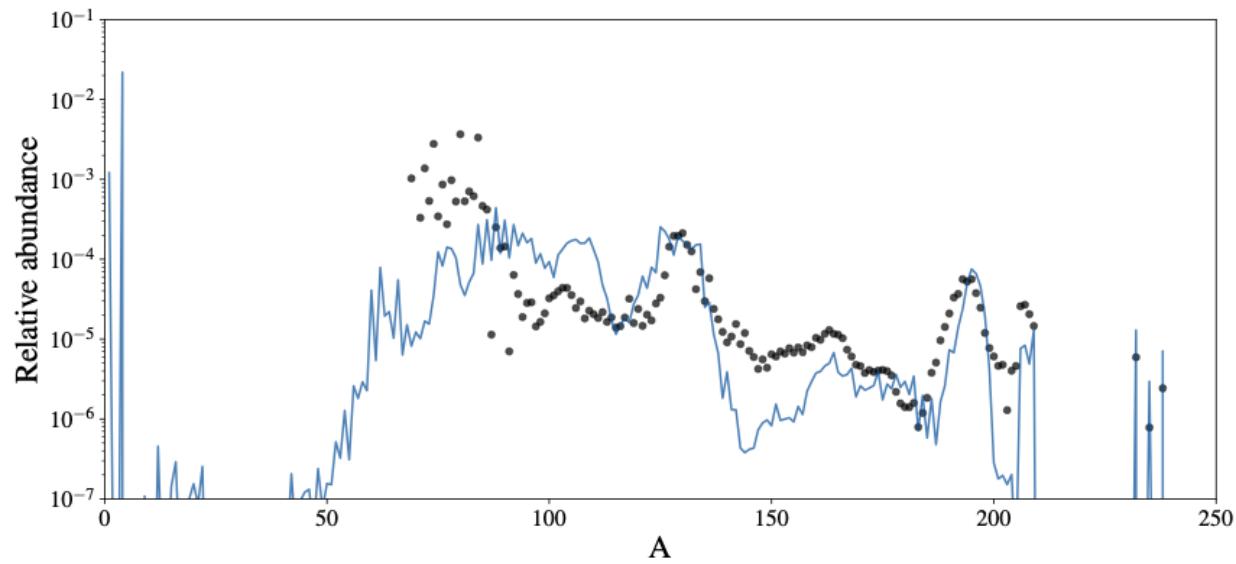
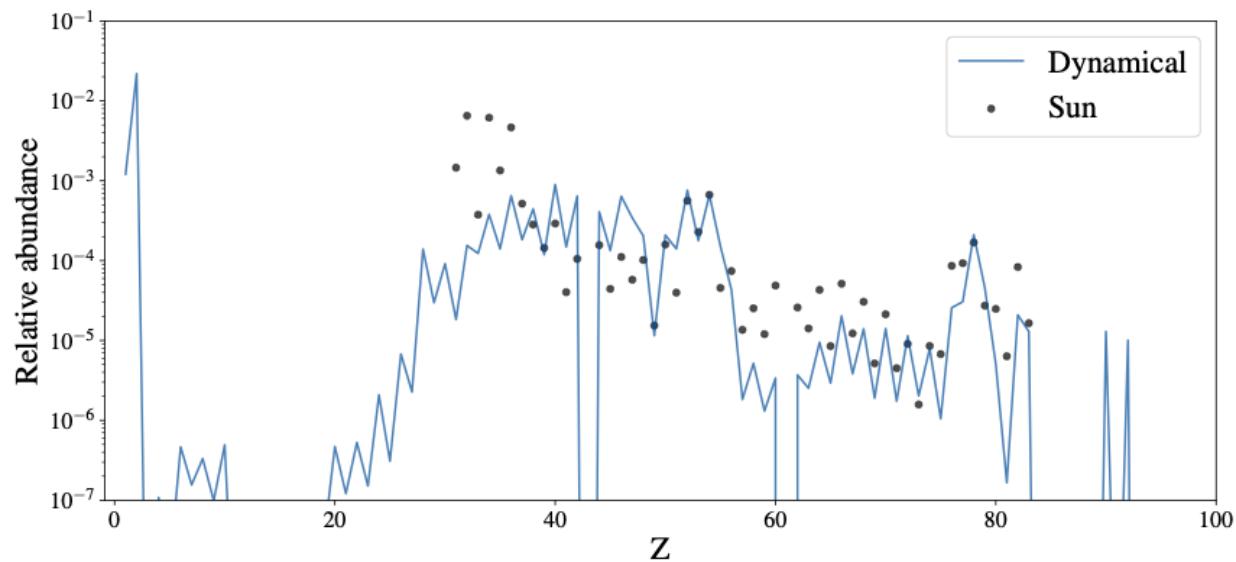
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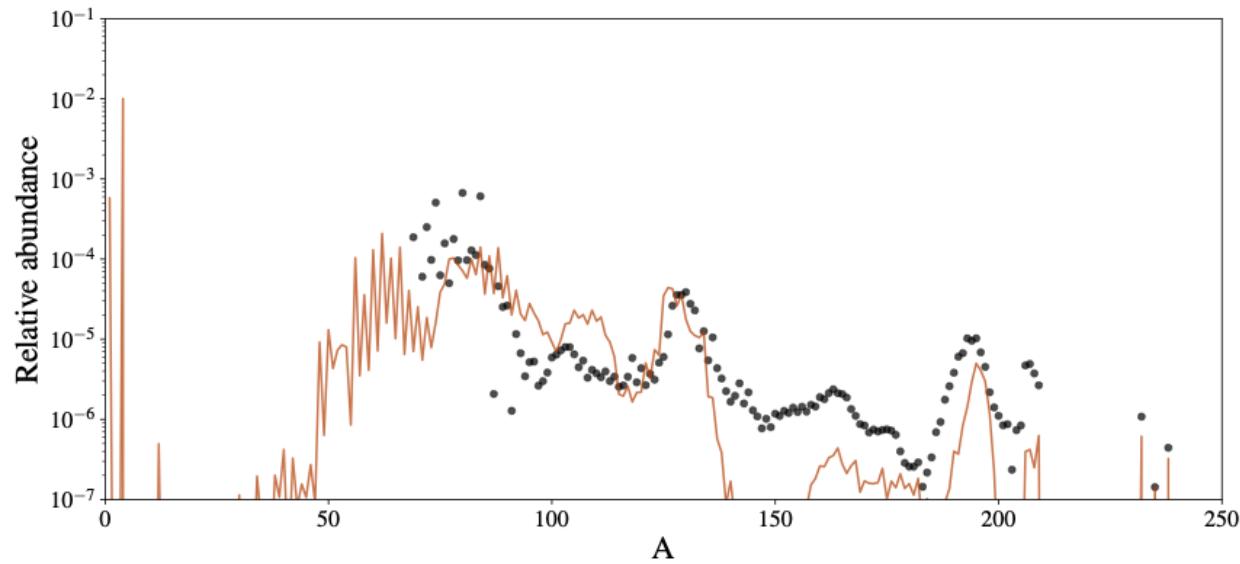
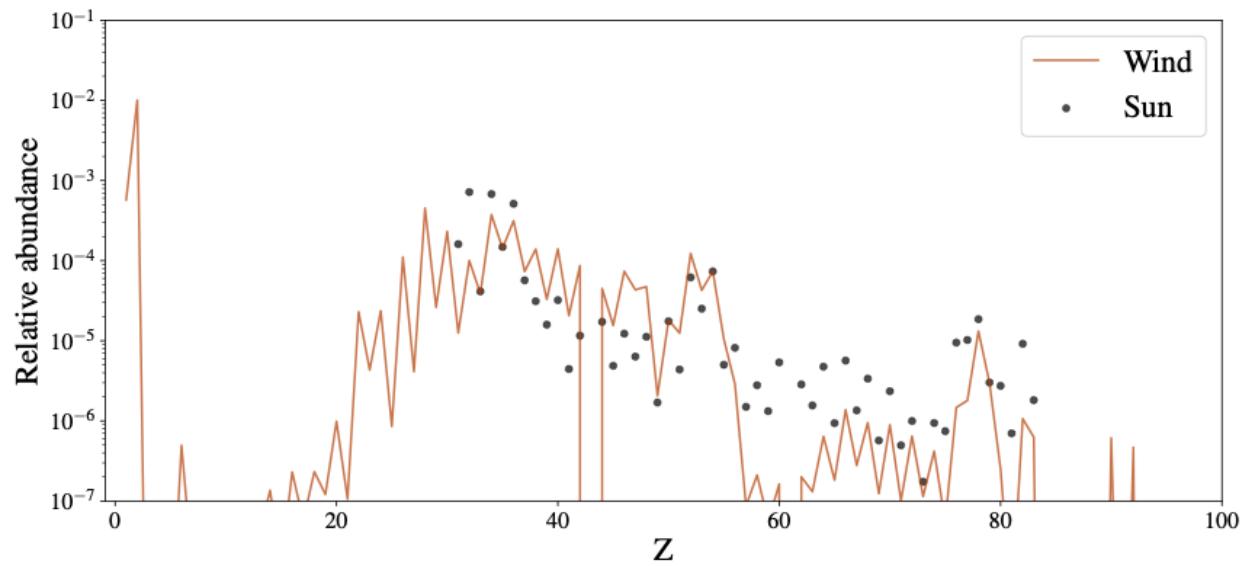
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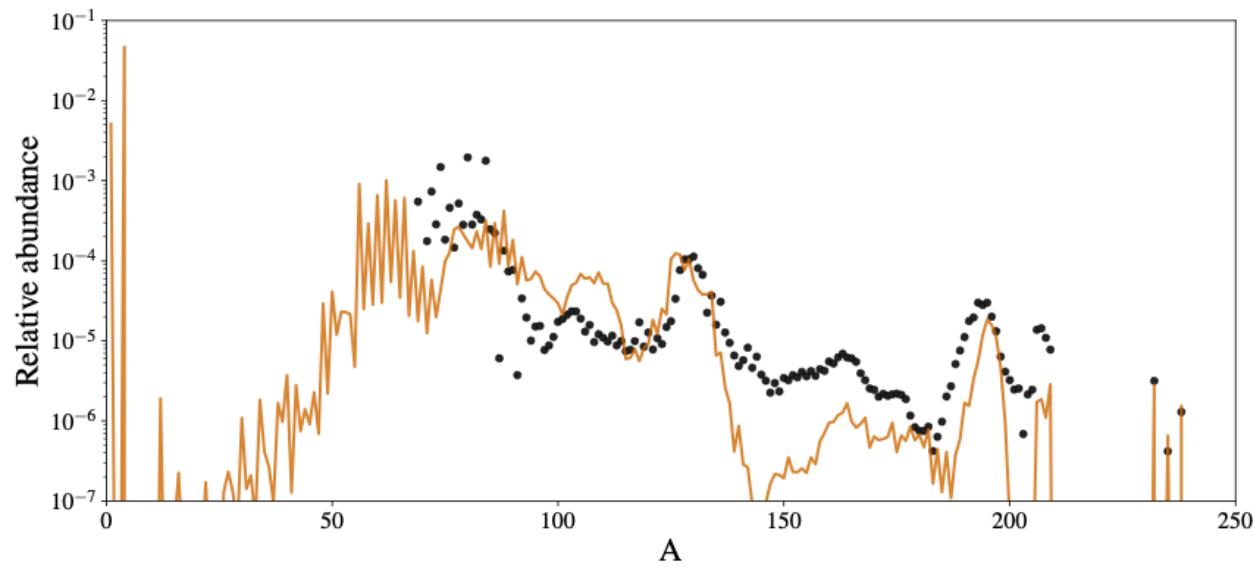
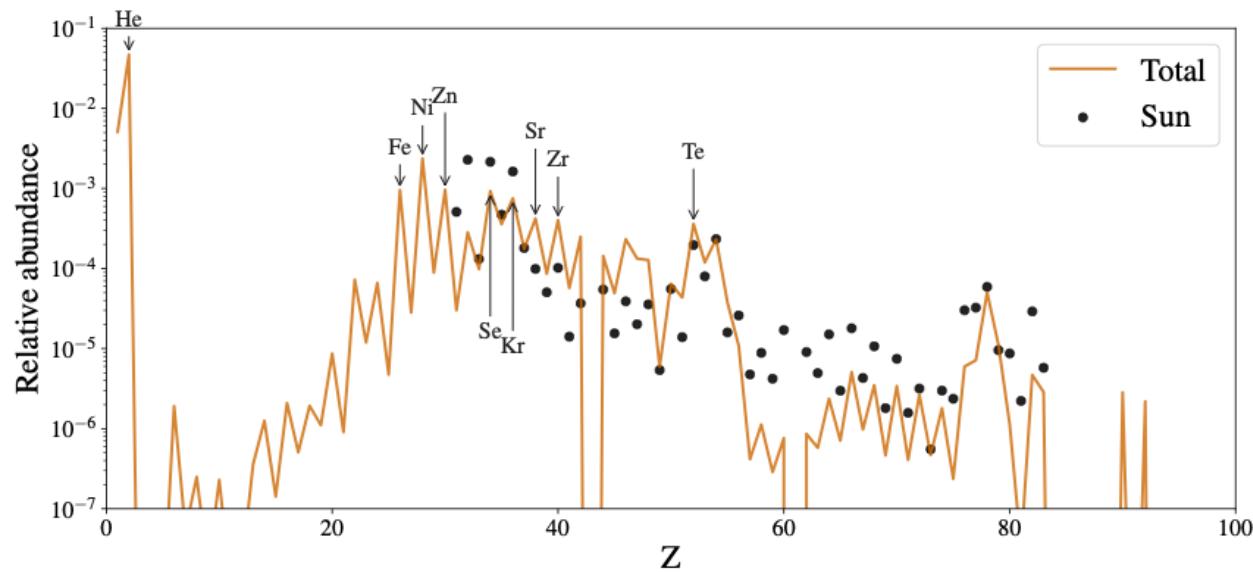


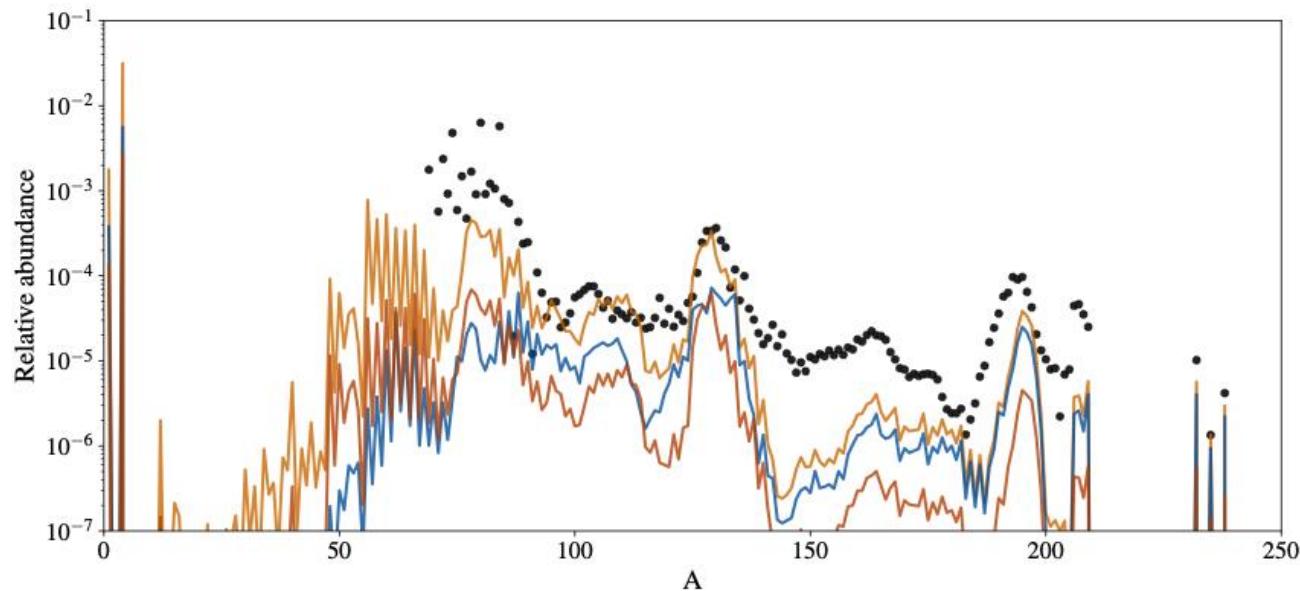
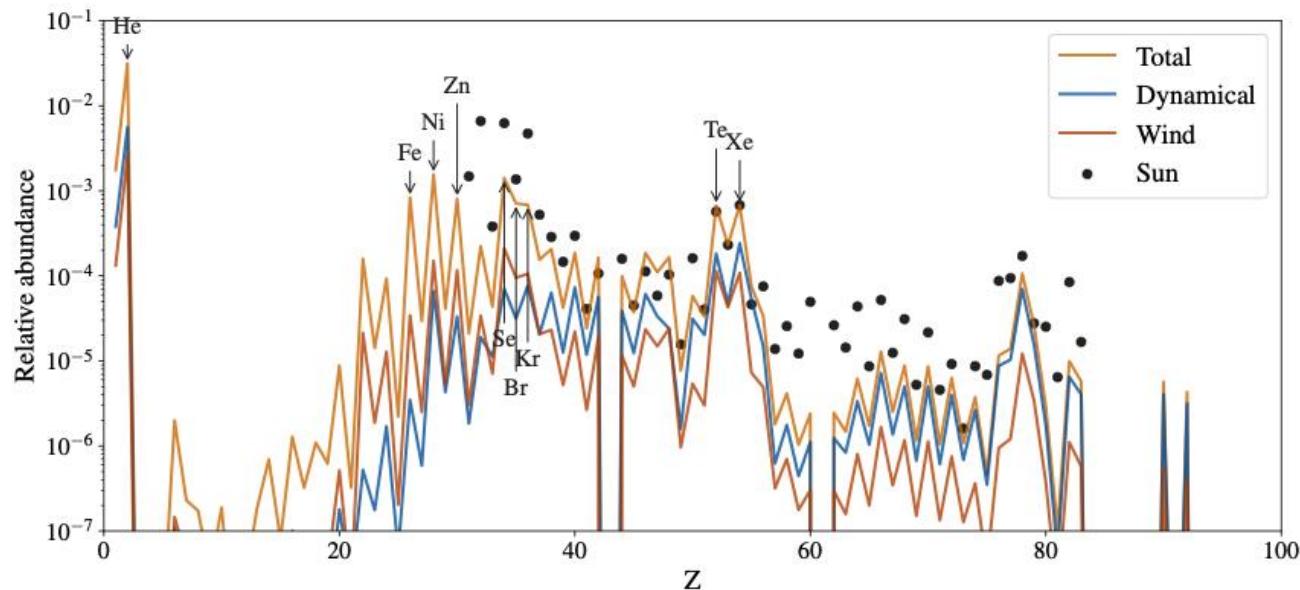
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See
Bezmalinovich's
talk



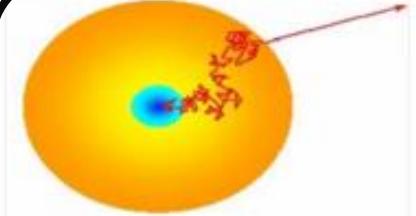
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Click on the button to download AGB yields.
[Go to data](#)



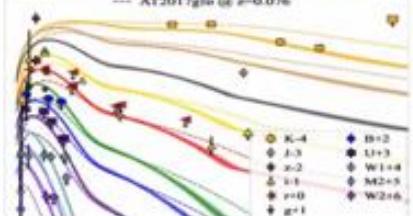
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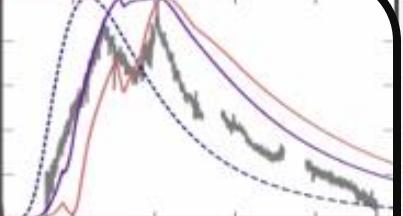
Dust-AGB
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Atomic-Opacities
Click on the button to download Element Atomic Opacities.
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--- AT2017gfo @ z=0.076
KNe-lightcurves
Click on the button to visualize Kilonovae Lightcurves.
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KNe-spectra
Click on the button to visualize Kilonovae Spectra.
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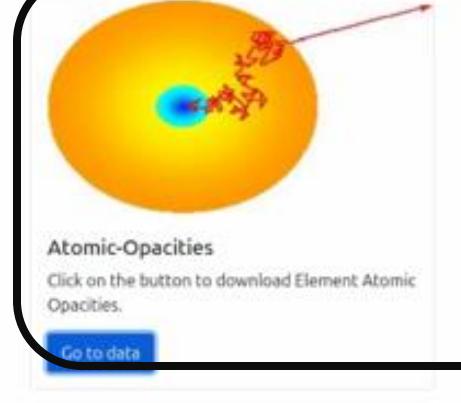
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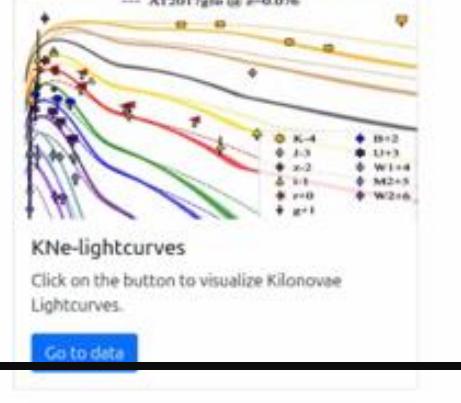
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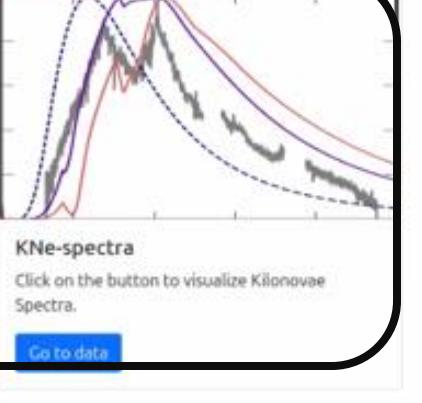
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KNe-spectra
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Conclusions

- Dynamical ejecta: production of elements in 2nd and 3rd peak.
- Neutrino wind: production of lighter elements.
- Iron group elements among the most abundant irrespective of the EOS.
- Nuclear input physics?
- Asymmetric mass ratios?
- Role of oscillations?