Institute of Space Sciences

EXCELENCIA MARÍA DE MAEZTU

Concurrent spectral and light curve synchro-curvature modelling of pulsars

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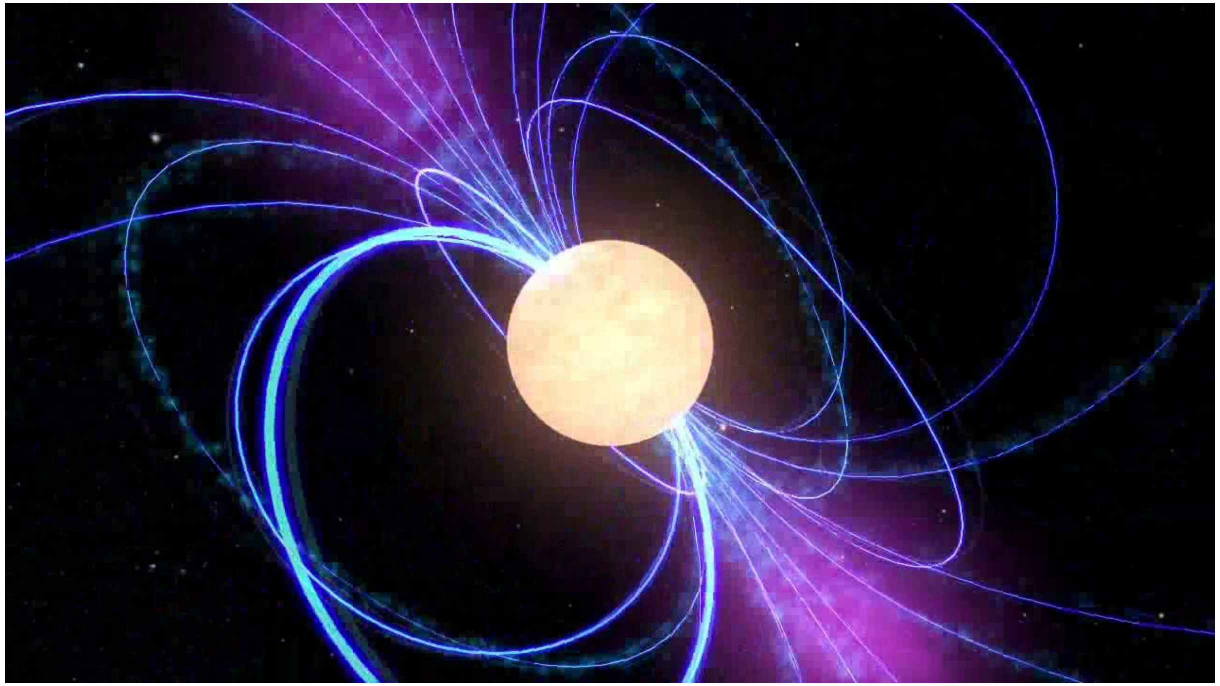
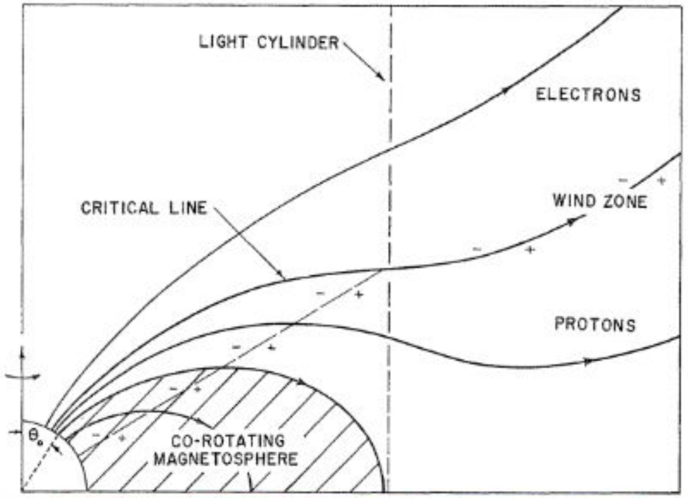
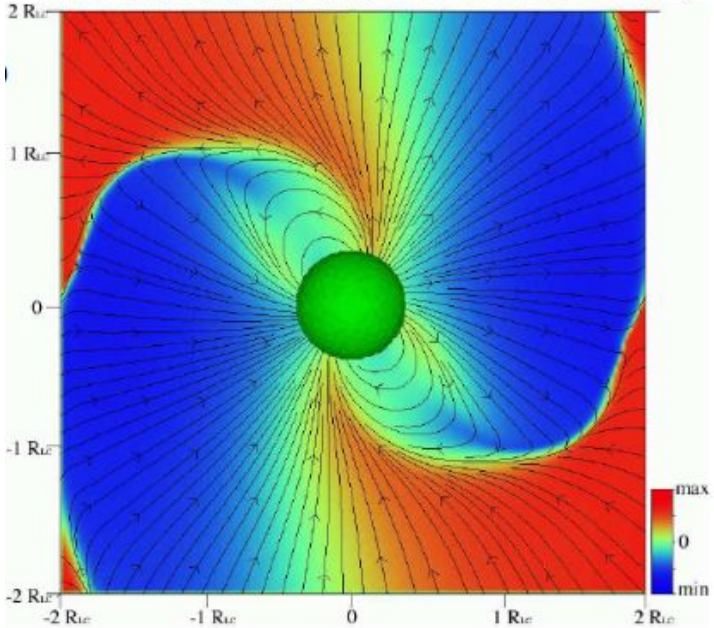
Motivation of the project

- High-energy pulsars are those with detected pulsated high-energy emission, gamma-rays and X-rays.



Credit: NASA/Fermi

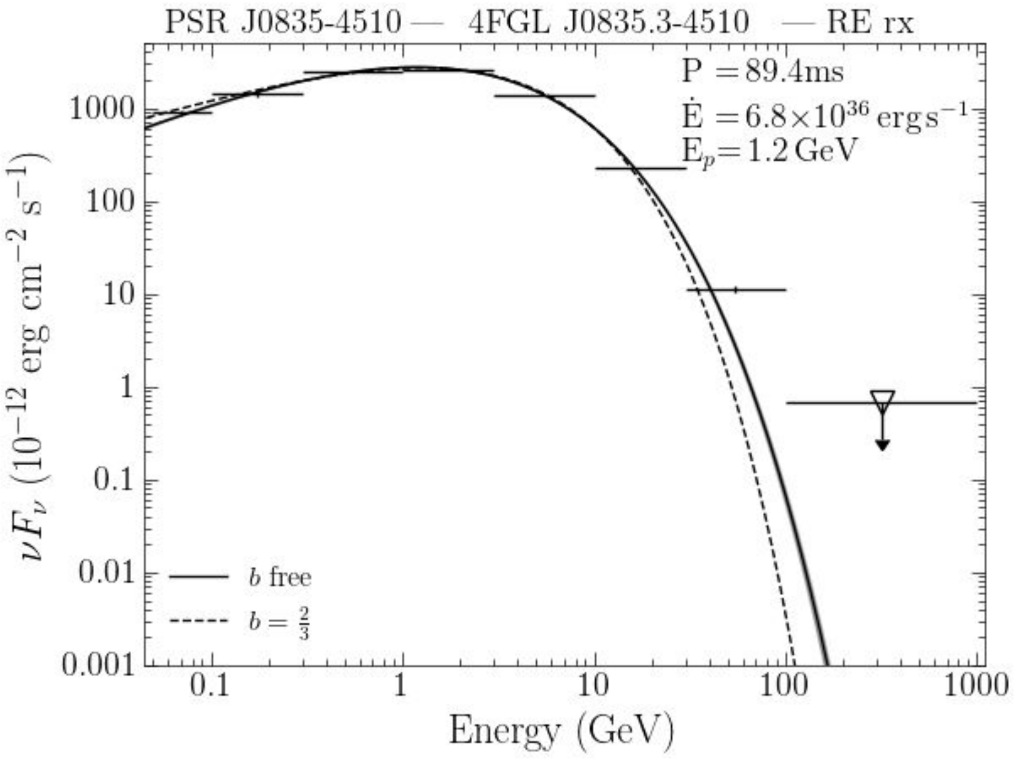


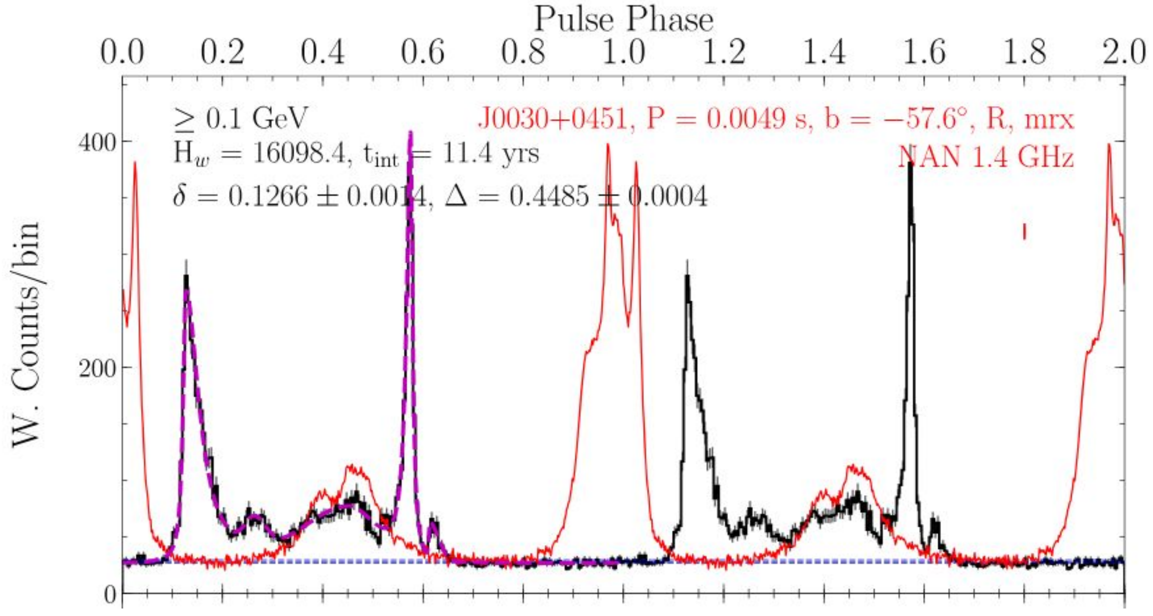


[Goldreich & Julian (1969); Spitkovsky (2006)]

Motivation of the project

- Models can be tested with e.g. the data of the gamma-ray emitting pulsars released in the 3PC (Smith et al. (2023))





- Our goal is to reproduce observational data with a model that contains simple but realistic physics and is computationally affordable

Spectral model: particle dynamics

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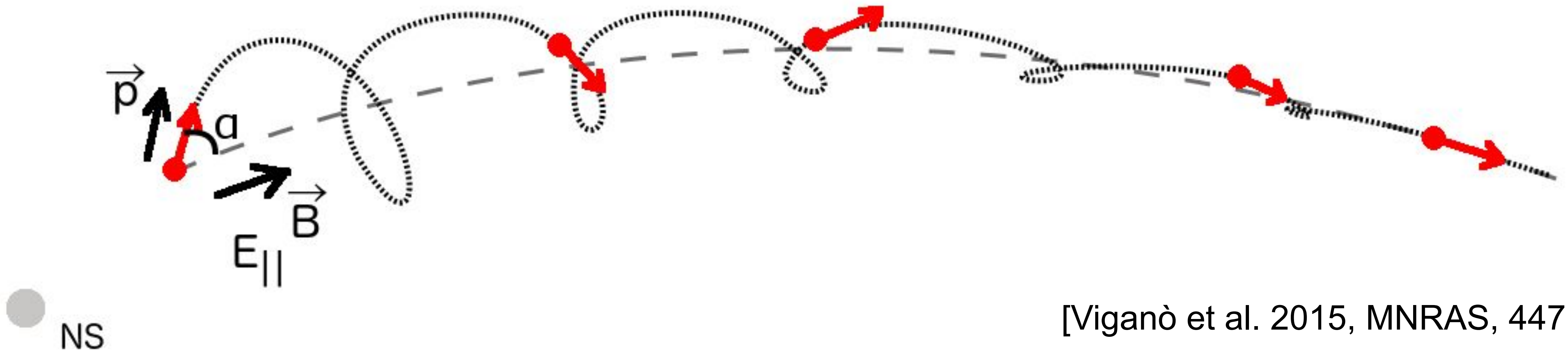
- We follow the dynamics of the emitting particles, ruled by electric acceleration and synchro-curvature losses and with two free parameters involved: $E_{||}$, b

- Solving the equation of motion gives the evolution of the relativistic momentum and of the Lorentz factor Γ and pitch angle α



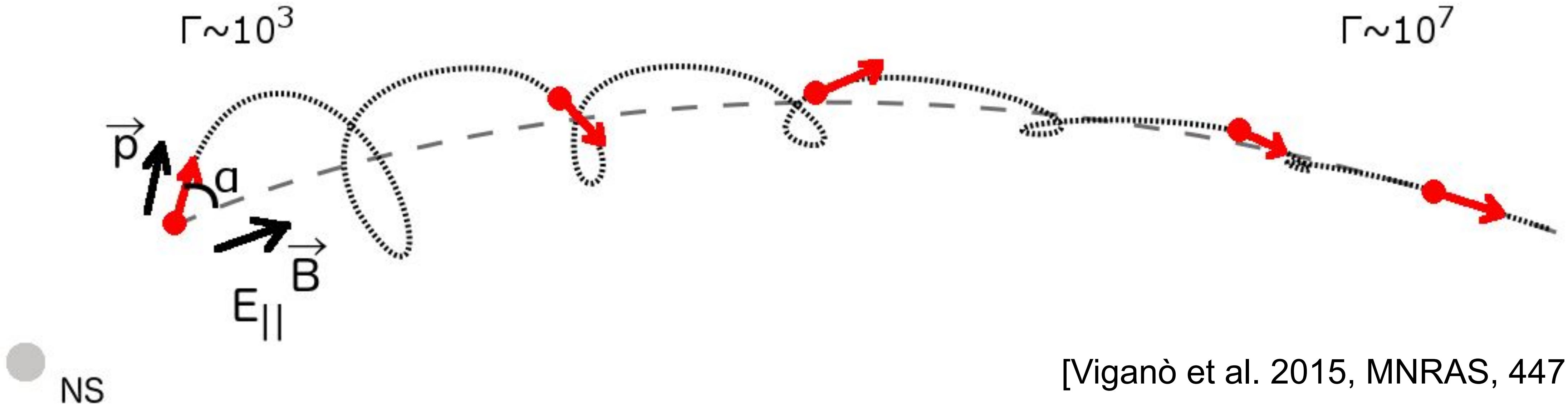
[Viganò et al. 2015, MNRAS, 447, 1164]

Spectral model: particle dynamics

- Power of synchrotron and curvature radiations depend on the Lorentz factor Γ of particles differently

$$P_{syn} = \frac{2}{3}\frac{(Ze)^4 B^2 (\Gamma^2 - 1)\sin^2\alpha}{m^2 c^3}$$

$$P_c = \frac{2}{3}\frac{(Ze)^2 c \Gamma^4}{r_c^2}$$

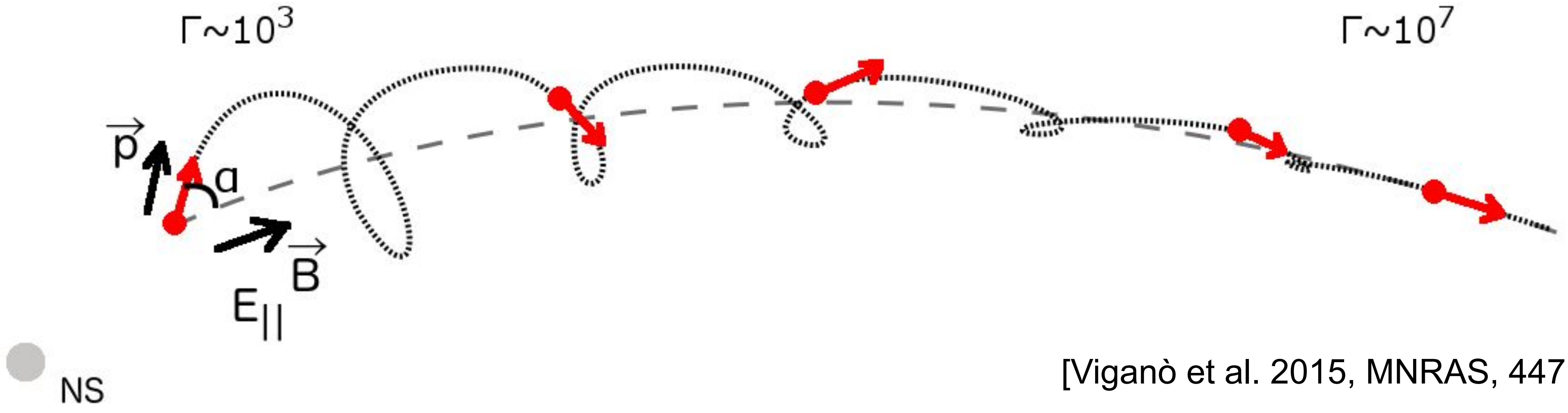


[Viganò et al. 2015, MNRAS, 447, 1164]

Spectral model: emission

- Synchro-curvature formulae gives the emission of the particles all along the trajectory, which convolved with an effective particle distribution gives the total radiation from the emission region

- We produce theoretical spectra with just three free parameter ($E_{||}$, b, x_0) and a normalization factor



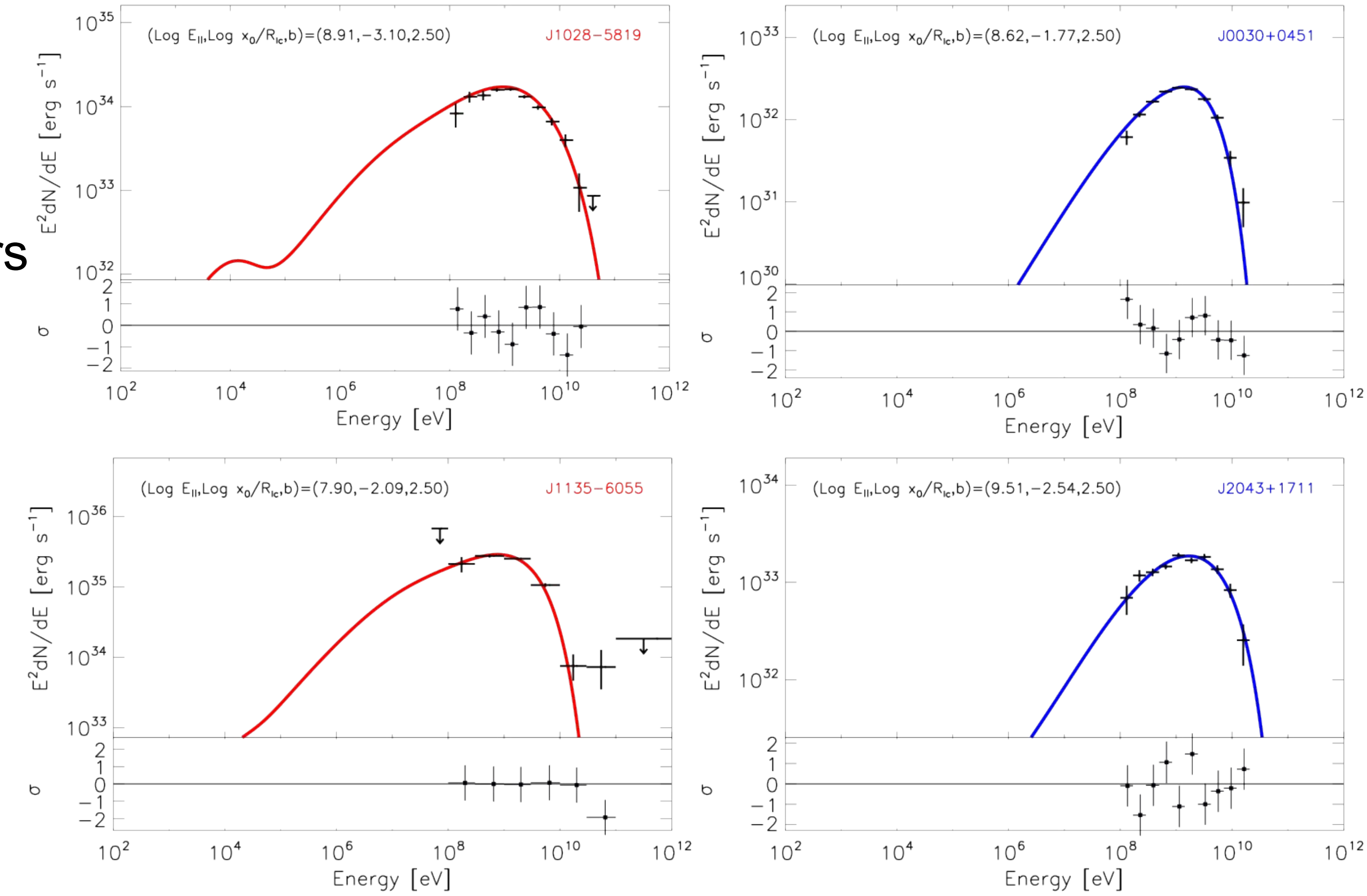
[Viganò et al. 2015, MNRAS, 447, 1164]

Spectral fitting

[Viganò et al. 2015, MNRAS, 453, 2599; Íñiguez-Pascual et al. (in prep.)]

- The model successfully fitted the 117 gamma ray-pulsars on the 2PC (Abdo et al. (2013)) and the ~300 gamma-ray pulsars on the 3PC (Smith et al. (2023))

- A relevant synchrotron contribution is needed to match gamma-ray spectra, implying that synchro-curvature radiation is an appropriate mechanism to explain the emission from these objects



Spectral fitting

[Torres 2018, Nat. Astron., 2, 247;
Torres et al. 2019, MNRAS, 489, 5494;
Íñiguez-Pascual, Viganò & Torres 2022, MNRAS, 516, 2475]

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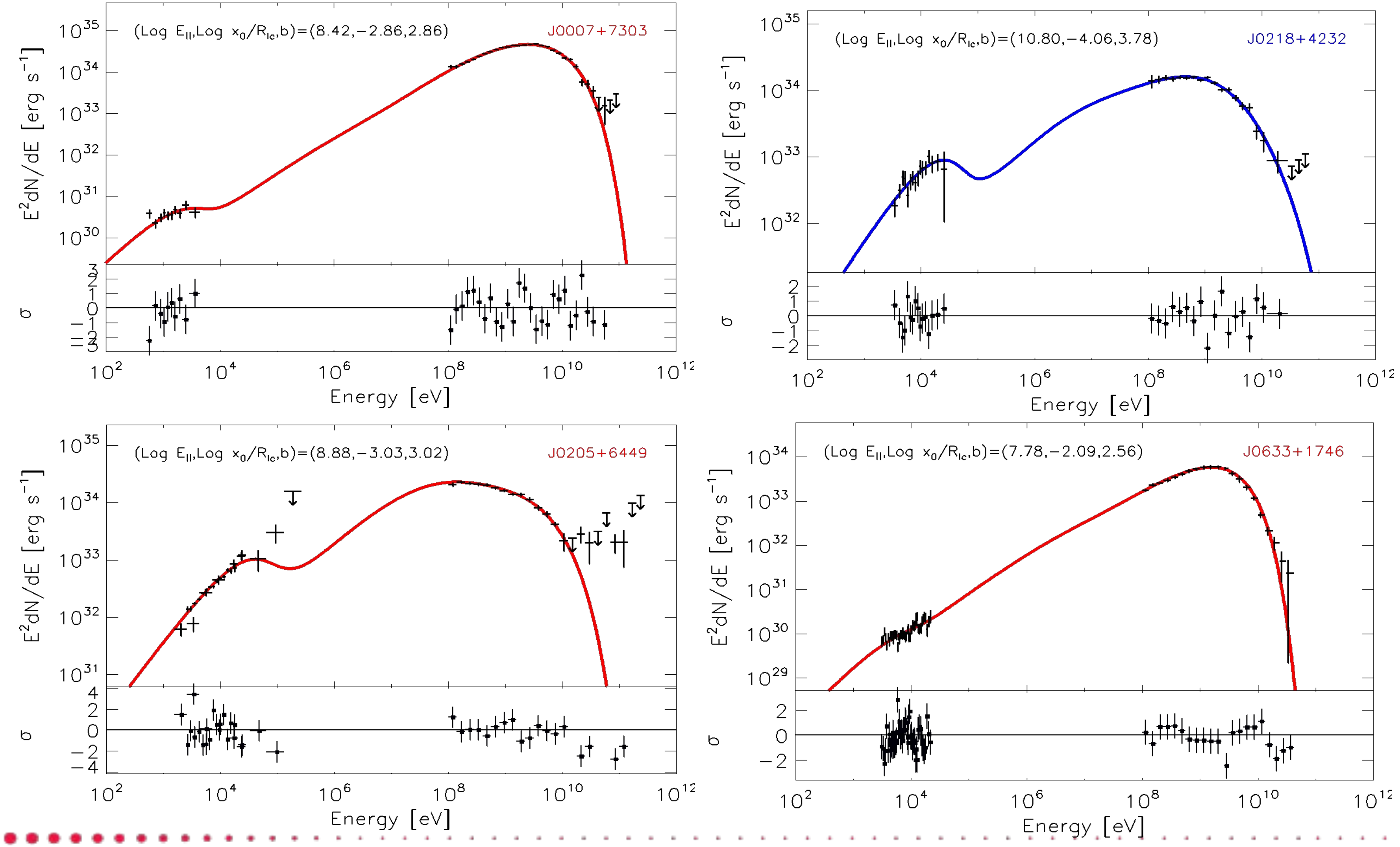
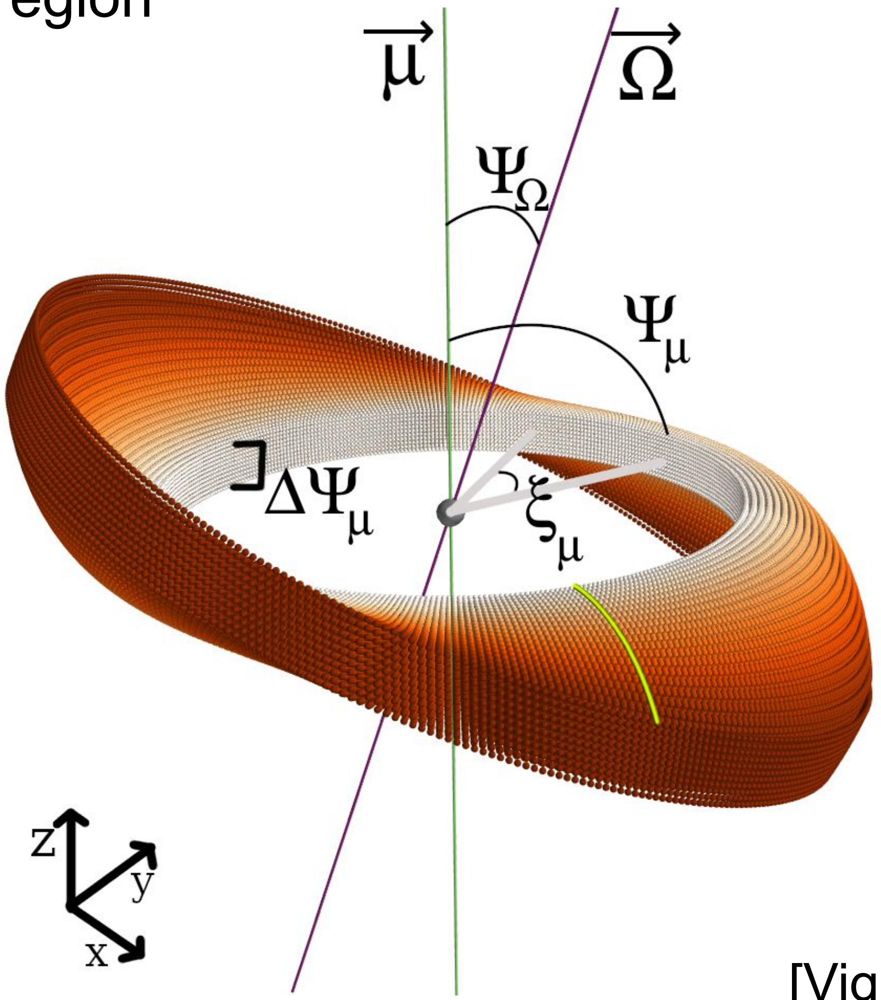
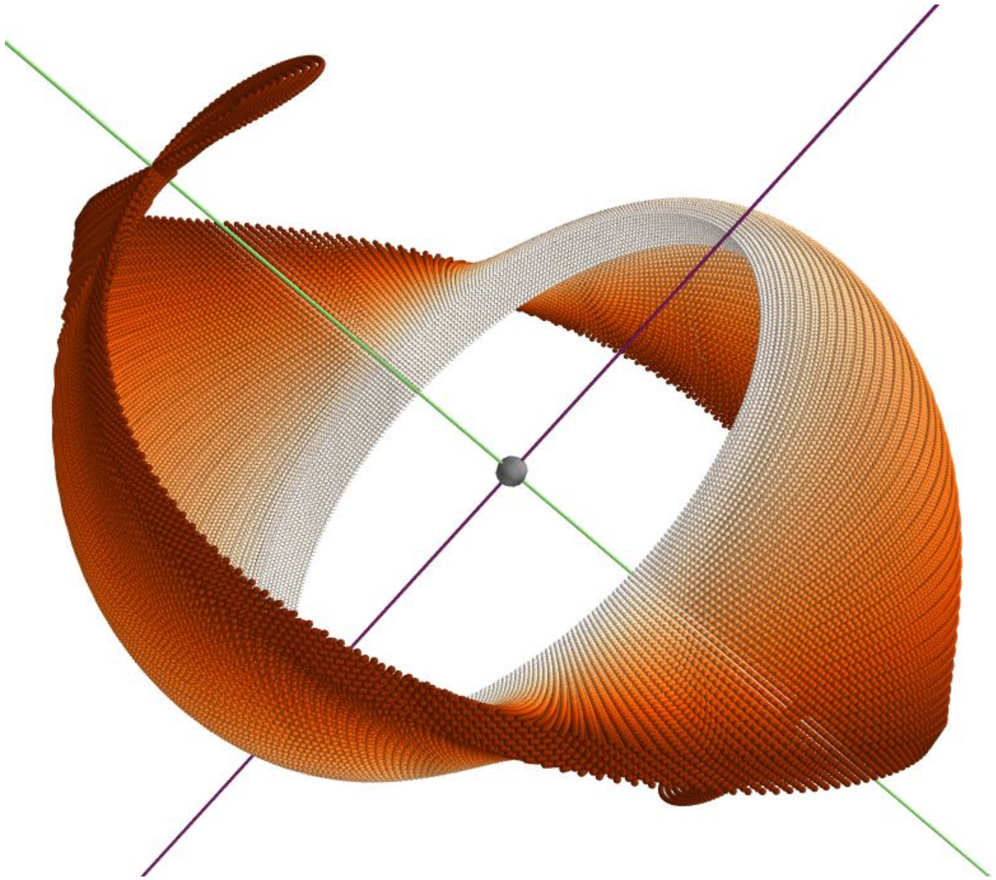


- The model can be extended to resemble the X-ray regime too, getting to do it in a majority of the ~40 high-energy pulsars

Geometrical model

- The inclination angle and the meridional extent define the geometry of the emission region

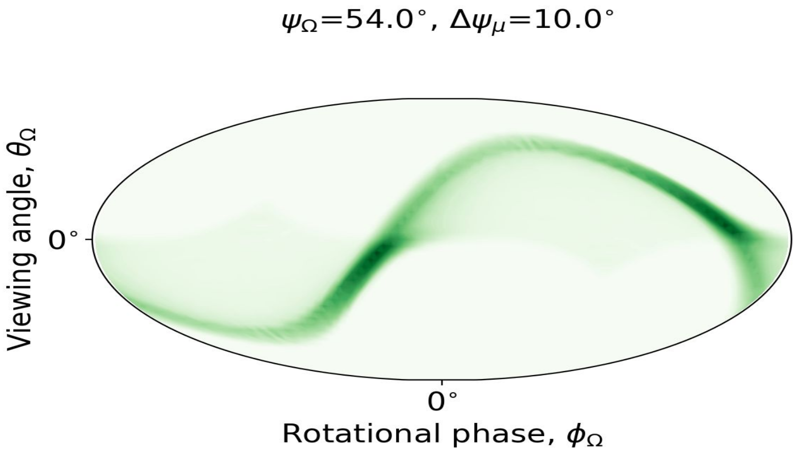


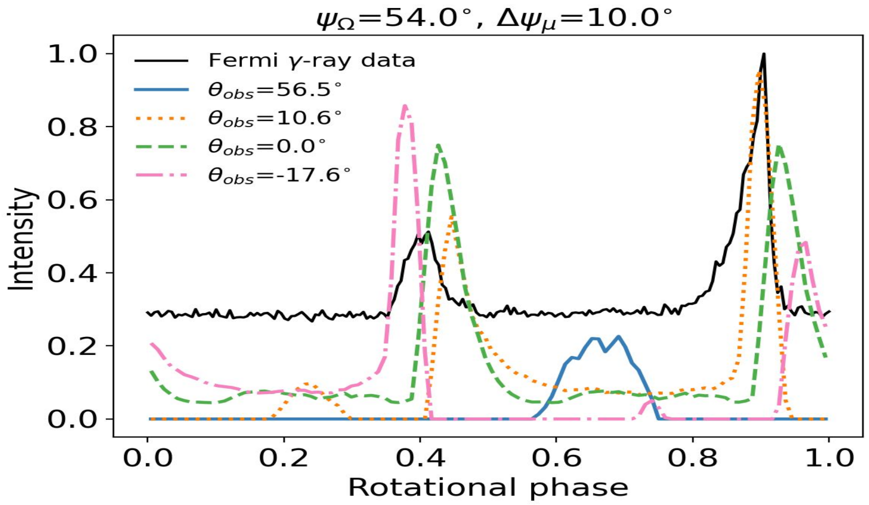


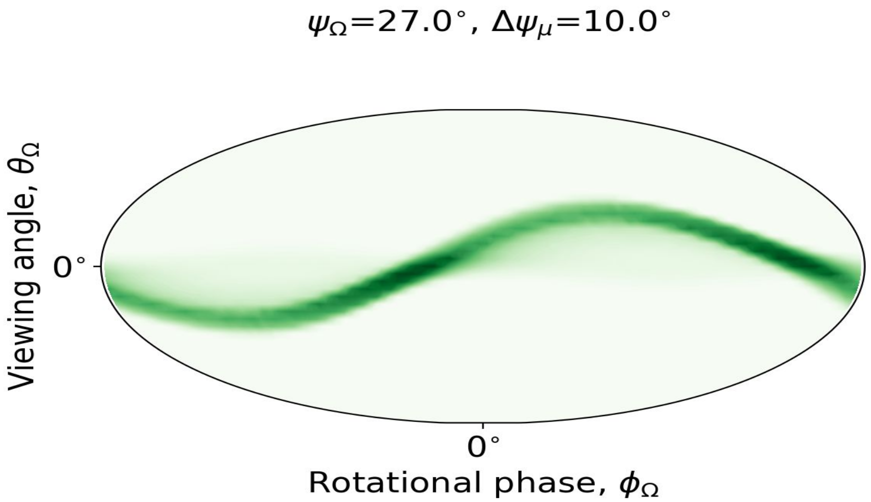
[Viganò & Torres (2019); Íñiguez-Pascual, Torres & Viganò (2024)]

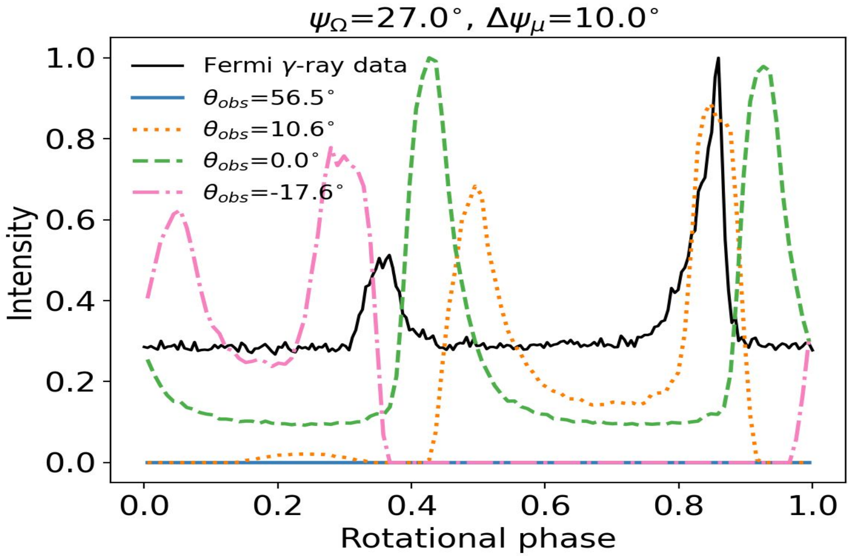
Emission maps and light curves

- We can build emission maps (skymaps), from which light curves are obtained





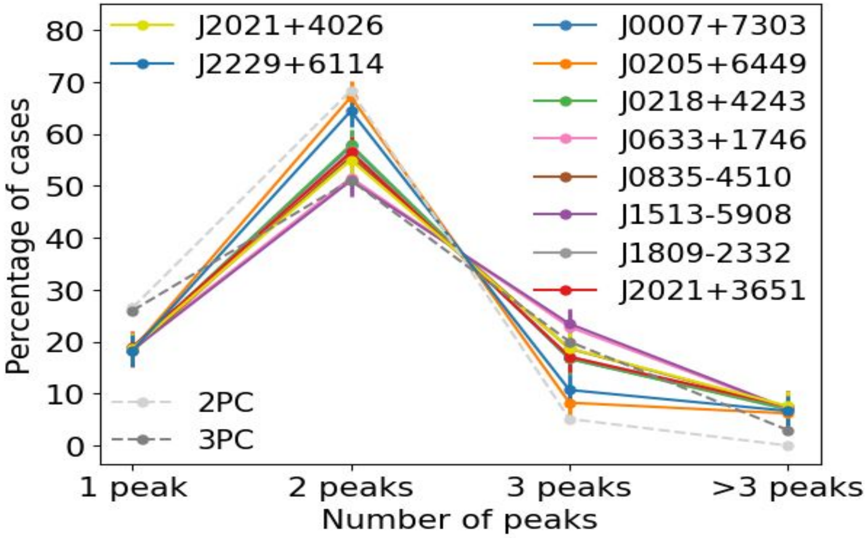


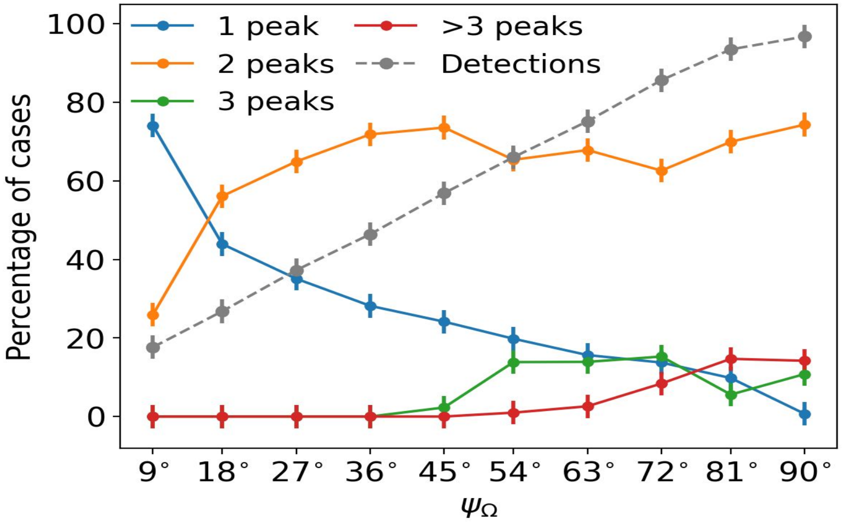


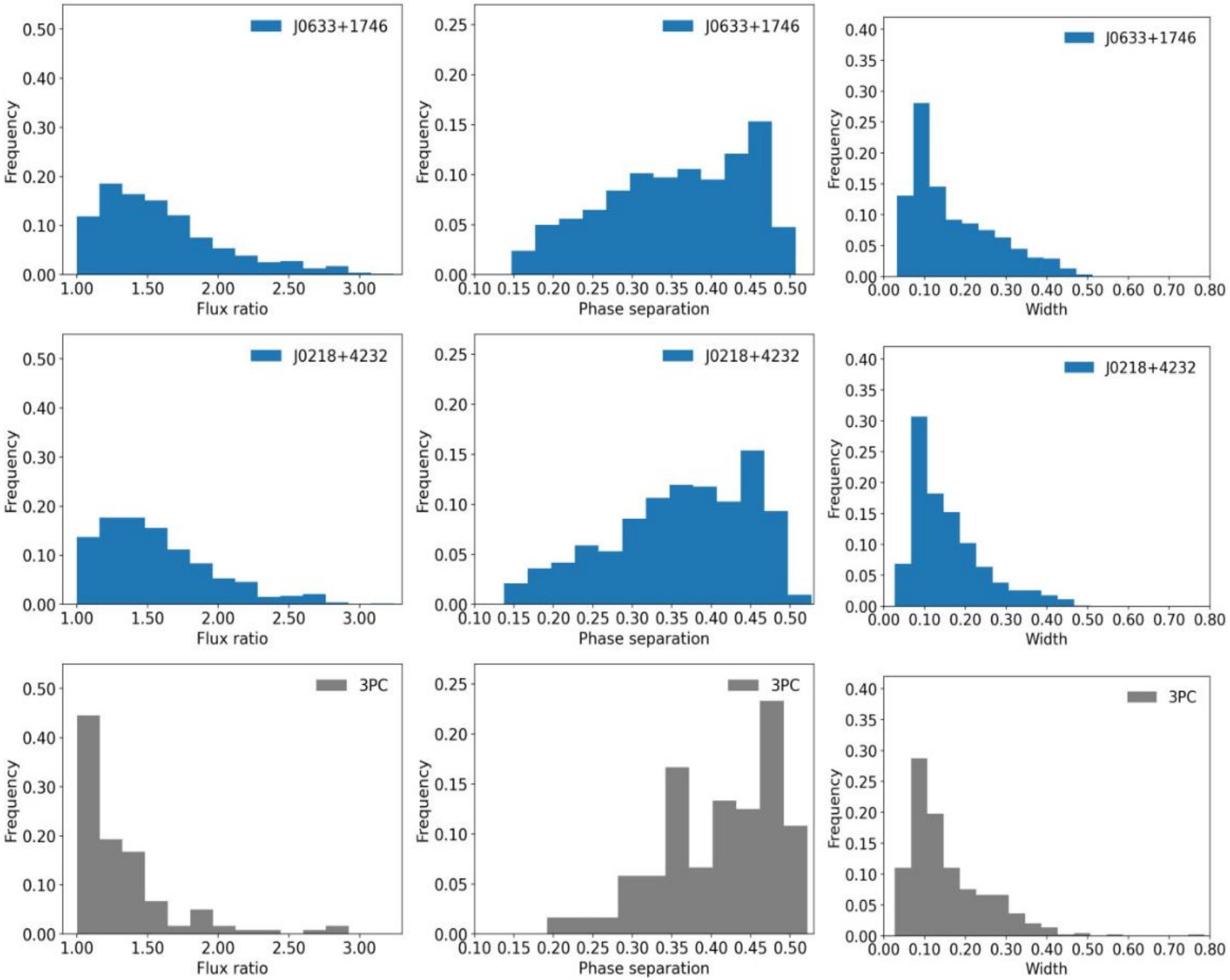
[Íñiguez-Pascual, Torres & Viganò 2024, MNRAS, 530, 1550]

Light curves statistics

[Íñiguez-Pascual, Torres & Viganò 2024, MNRAS, 530, 1550]

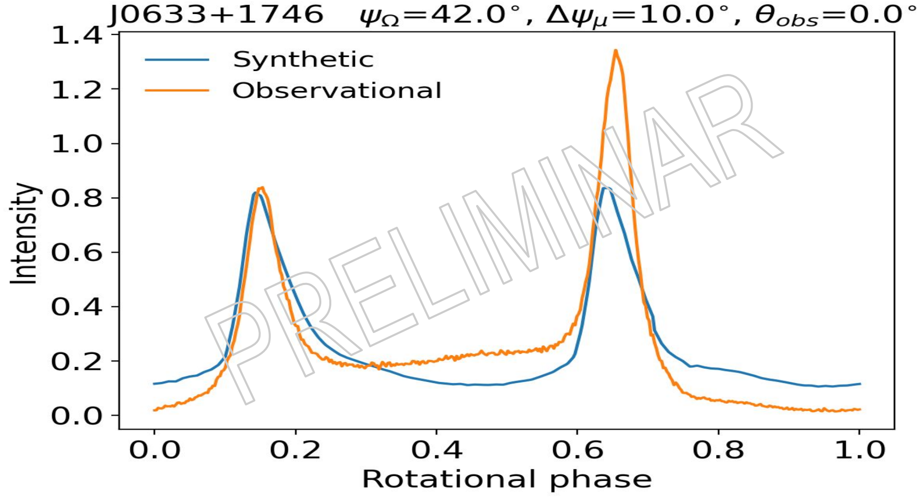


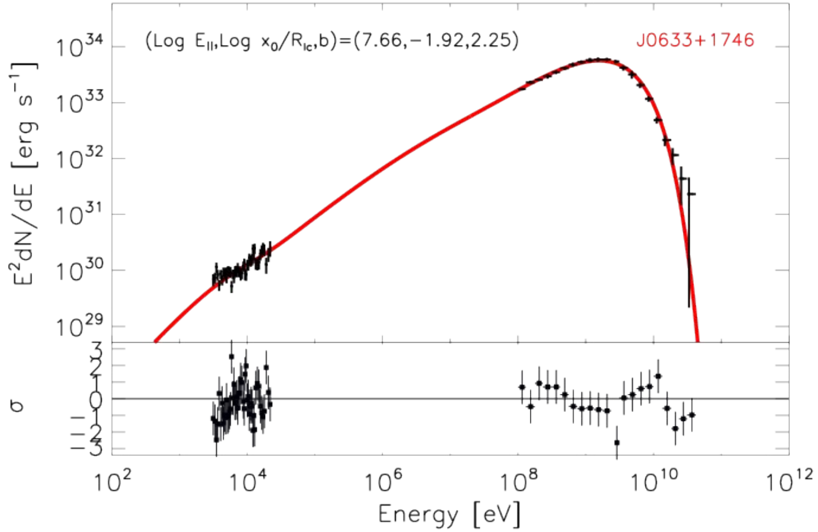


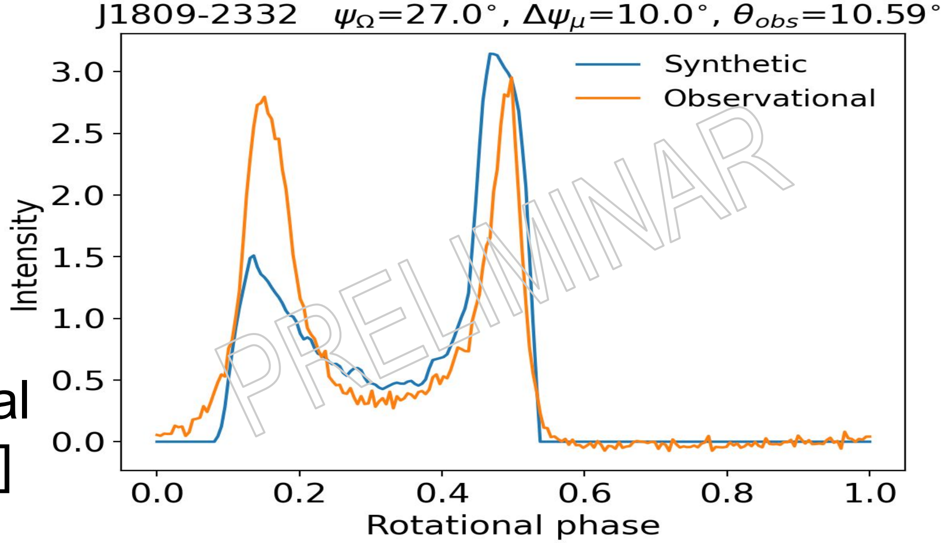


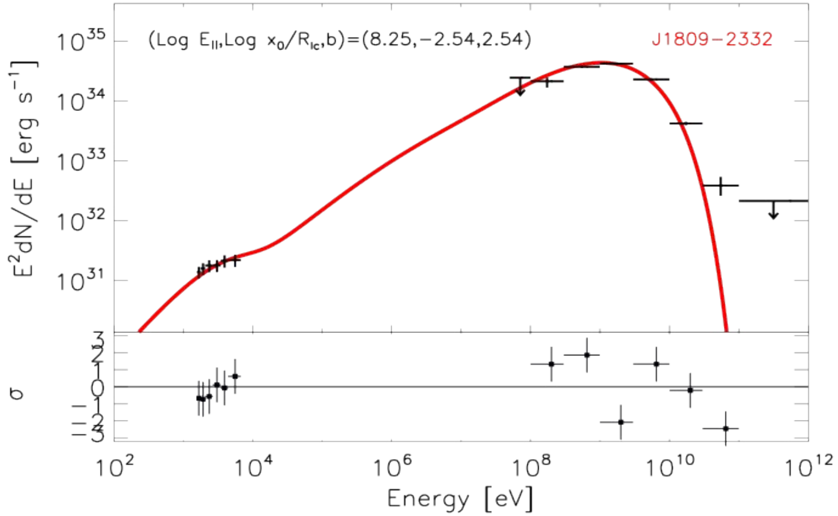
Light curves fitting

- Fitting synthetic light curves to observational ones, concurrently to the spectral fitting









[Íñiguez-Pascual et al. (in prep.)]

Conclusions

- Our spectral model is able to properly fit the population of high-energy emitting pulsars, showing that synchro-curvature radiation is an appropriate mechanism to explain the emission from these objects

  [*Íniguez-Pascual D., Viganò D., Torres D. F., 2022, MNRAS, 516, 2475 (2208.05549)*]

- The geometrical model reproduces the variety of observational gamma-ray light curves but still cannot capture all their small scale features

  [*Íniguez-Pascual D., Torres D. F., Viganò D., 2024, MNRAS, 530, 1550 (2404.01926)*]

Future prospects

- Improve the spectral and geometrical models by including more realistics physics while keeping our effective approach

- Possess a fully working concurrent spectral and light curves fitting

Thank you

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Back up slide: particle dynamics formulae

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- The equation of motion of charged particles balances electric acceleration and synchro-curvature losses

$$\frac{d\vec{p}}{dt} = ZeE_{\parallel}\hat{b} - (P_{sc}/v)\hat{p} \qquad \vec{p} = \Gamma m \vec{v}$$

[Viganò et al. (2014)]

- We solve it numerically, considering separately the components parallel and perpendicular to the trajectory

$$\frac{d(p\cos\alpha)}{dt} = ZeE_{\parallel} - \frac{P_{sc}}{v}\cos\alpha \qquad \frac{d(p\sin\alpha)}{dt} = -\frac{P_{sc}}{v}\sin\alpha$$

- Local magnetic field strength and curvature radius are parametrize in an effective way:

$$B = B_{\star}\left(\frac{R_{\star}}{x}\right)^{b} \qquad r_c = R_{lc}\left(\frac{x}{R_{lc}}\right)^{\eta}$$

Back up slide: synchro-curvature formulae

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- Single-particle synchro-curvature power spectra:

$$\frac{dP_{sc}}{dE} = \frac{\sqrt{3}(Ze)^2\Gamma y}{4\pi\hbar r_{eff}}\left[(1+z)F(y) - (1-z)K_{2/3}(y)\right]$$

- Synchro-curvature power: $P_{sc} = \frac{2(Ze)^2\Gamma^4 c}{3r_c^2} g_r$

[Cheng & Zhang (1996), Viganò et al. (2014)]

$$r_{eff} = \frac{r_c}{\cos^2\alpha}\left(1+\xi+\frac{r_{gyr}}{r_c}\right)^{-1}$$

$$z = (Q_2 r_{eff})^{-2}$$

$$y = \frac{E}{E_c}$$

$$Q_2^2 = \frac{\cos^4\alpha}{r_c^2}\left[1+3\xi+\xi^2+\frac{r_{gyr}}{r_c}\right]$$

$$E_c = \frac{3}{2}\hbar c Q_2 \Gamma^3$$

$$F(y) = \int_y^\infty K_{5/3}(y')dy'$$

$$\xi = \frac{r_c}{r_{gyr}}\frac{\sin^2\alpha}{\cos^2\alpha}$$

Back up slide: synchro-curvature formulae

- Convolving the single-particle power spectra with an effective particle distribution,

$$\frac{dN}{dx} = N_0 \frac{e^{-(x-x_{min})/x_0}}{x_0(1-e^{-(x_{max}-x_{min})/x_0})}$$

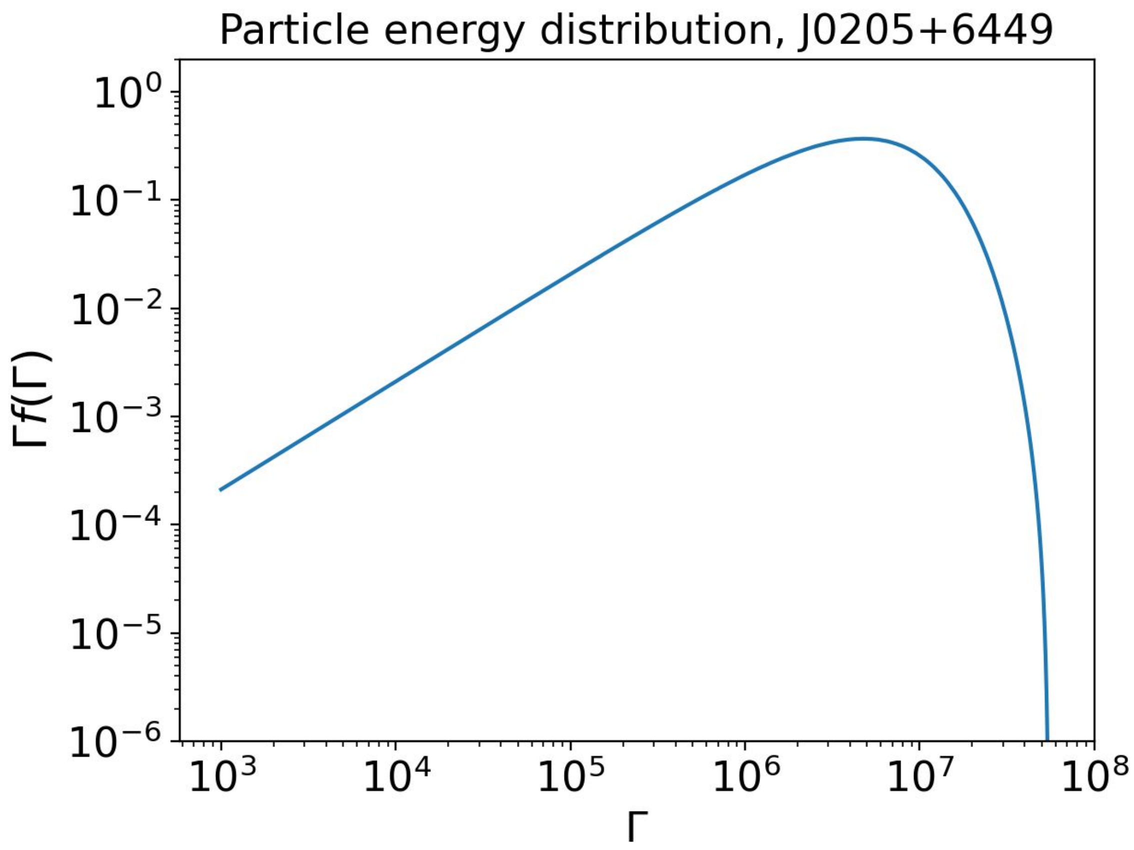
We obtain the total emission from the region:

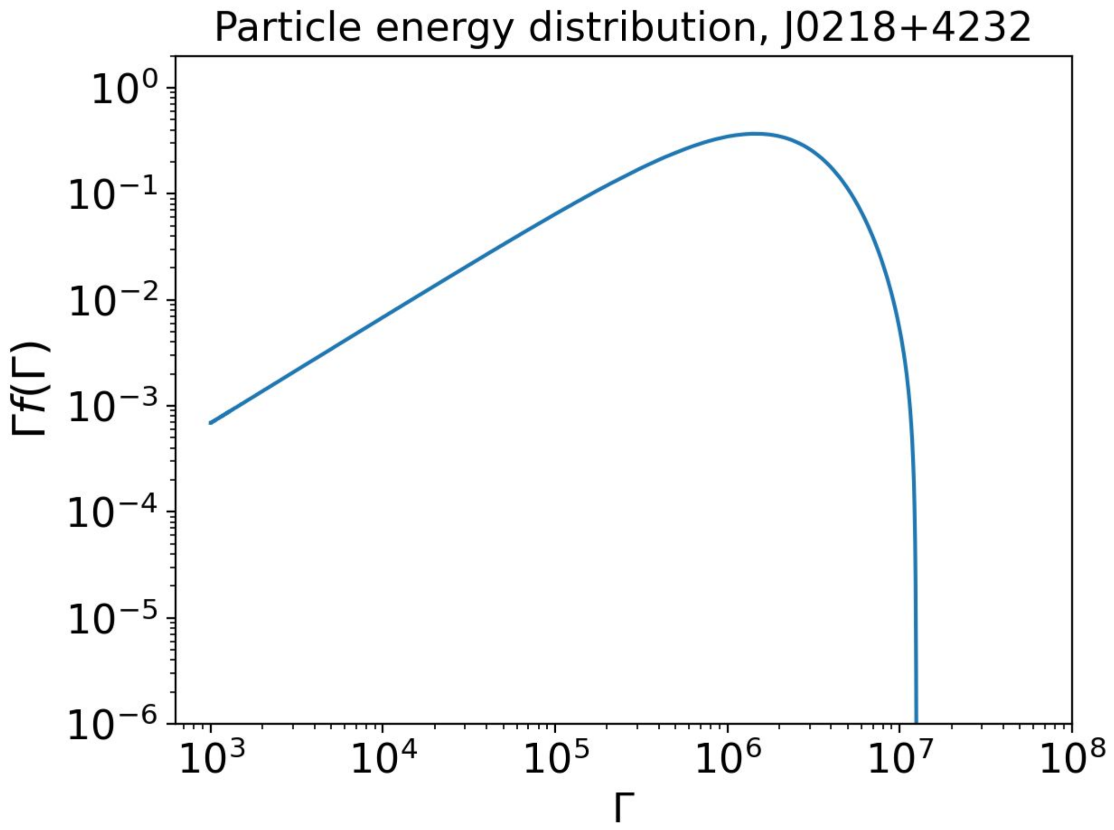
$$\frac{dP_{tot}}{dE} = \int_{x_{min}}^{x_{max}} \frac{dP_{sc}}{dE}\frac{dN}{dx}dx$$

[Cheng & Zhang (1996), Viganò et al. (2014)]

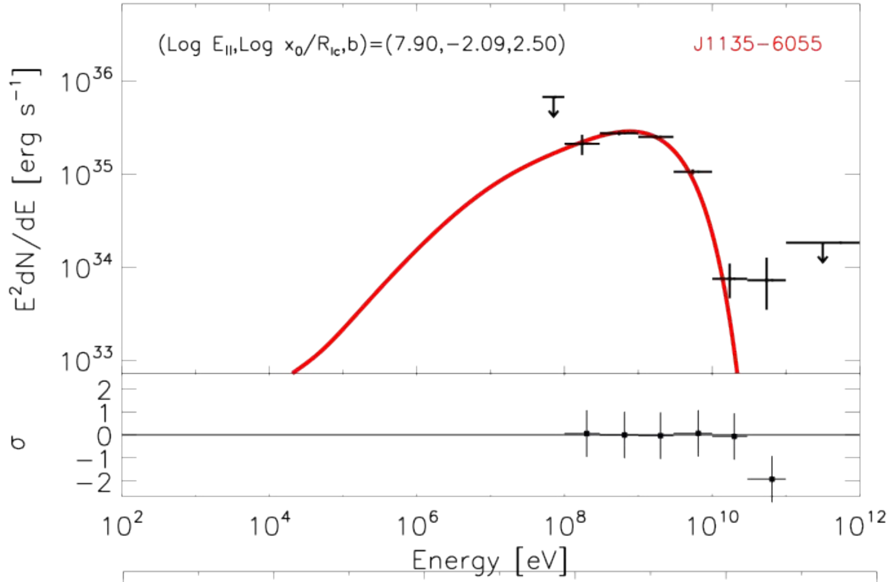
Back up slide: particle energy distribution

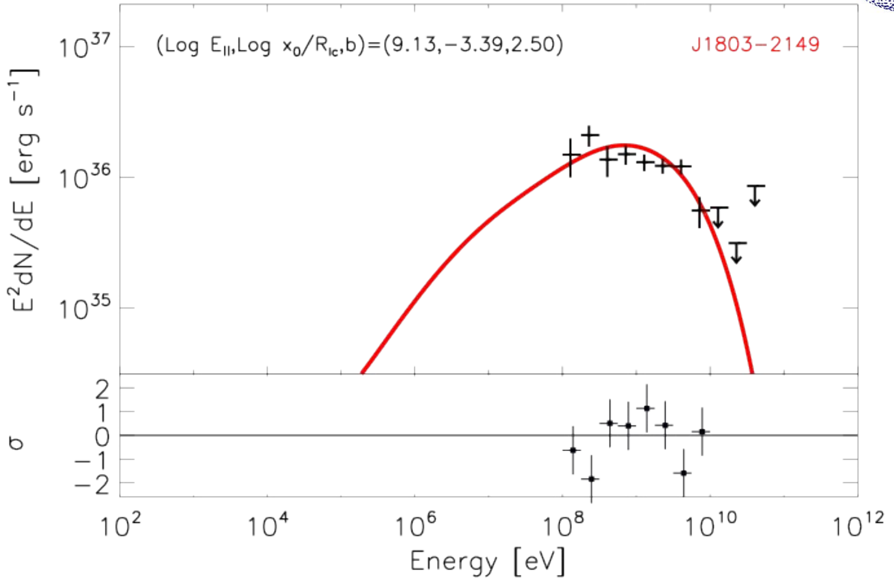
- Lorentz factors Γ typically range from 10^3 to 10^7.

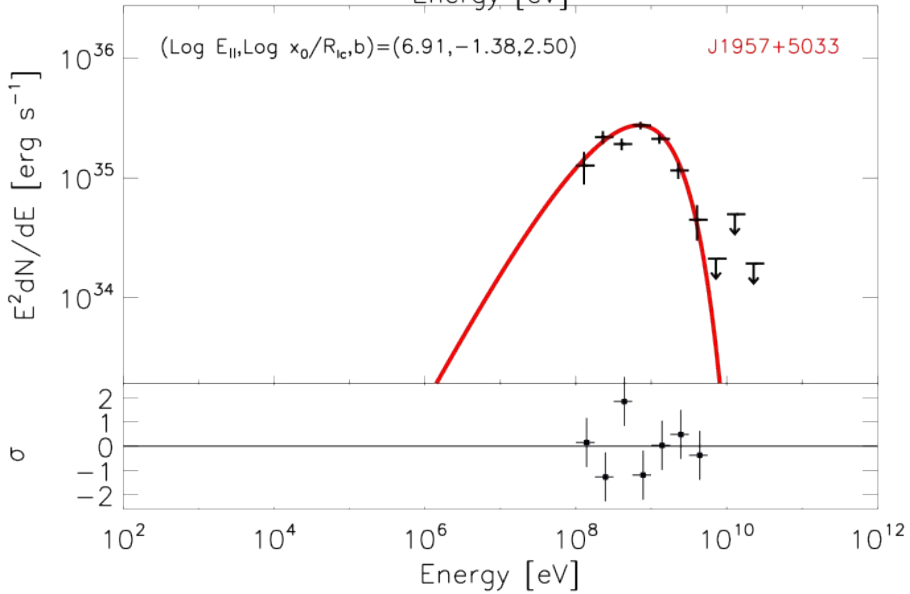


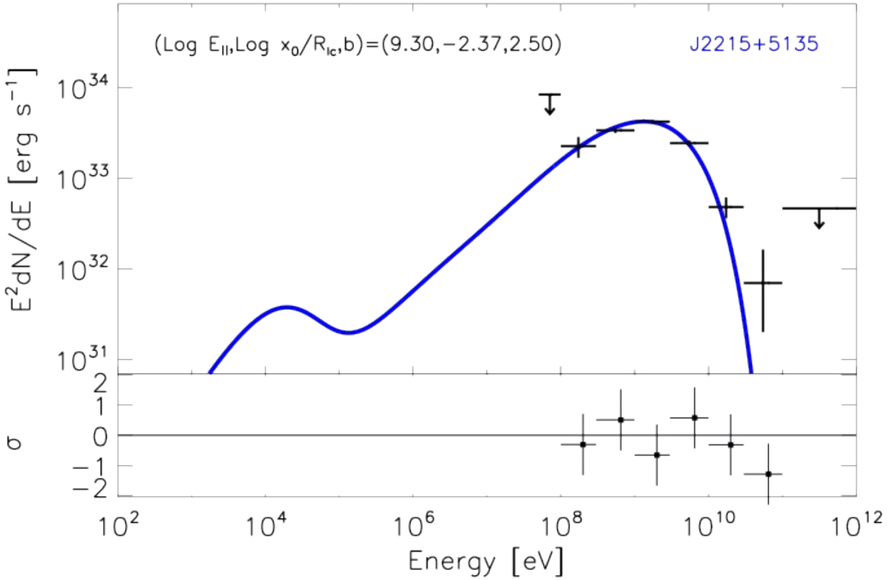


Back up slide: more spectral fits

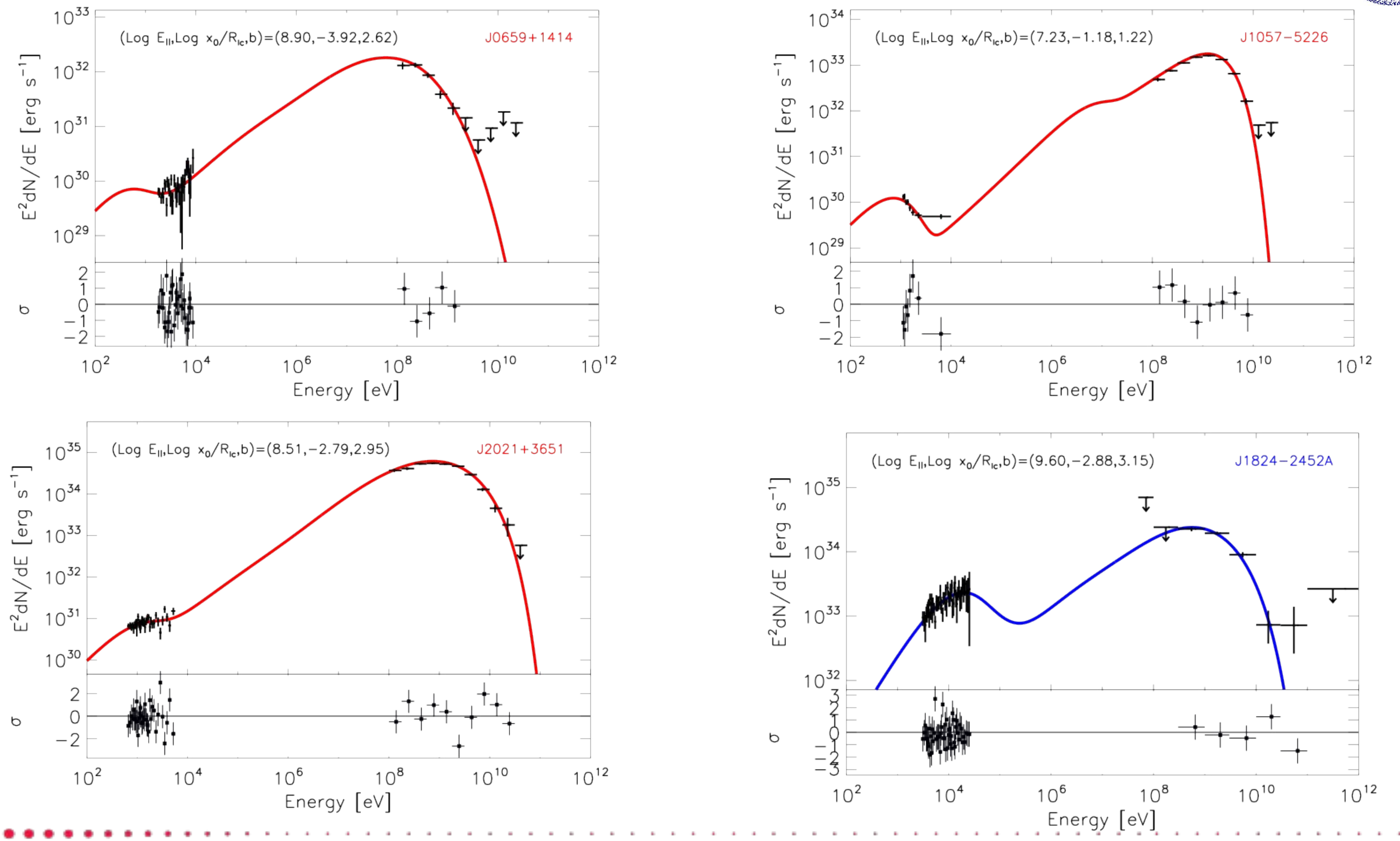








Back up slide: more spectral fits



Back up slide: geometrical model

- Frenet-Serret equations allow to geometrically describe a curved trajectory with torsion

$$\frac{d\hat{t}}{d\lambda} = \frac{1}{r_c}\hat{n} \,,$$

$$\frac{d\hat{n}}{d\lambda} = -\frac{1}{r_c}\hat{t} + \tau\hat{b} \,,$$

$$\frac{d\hat{b}}{d\lambda} = -\tau\hat{n} \,.$$

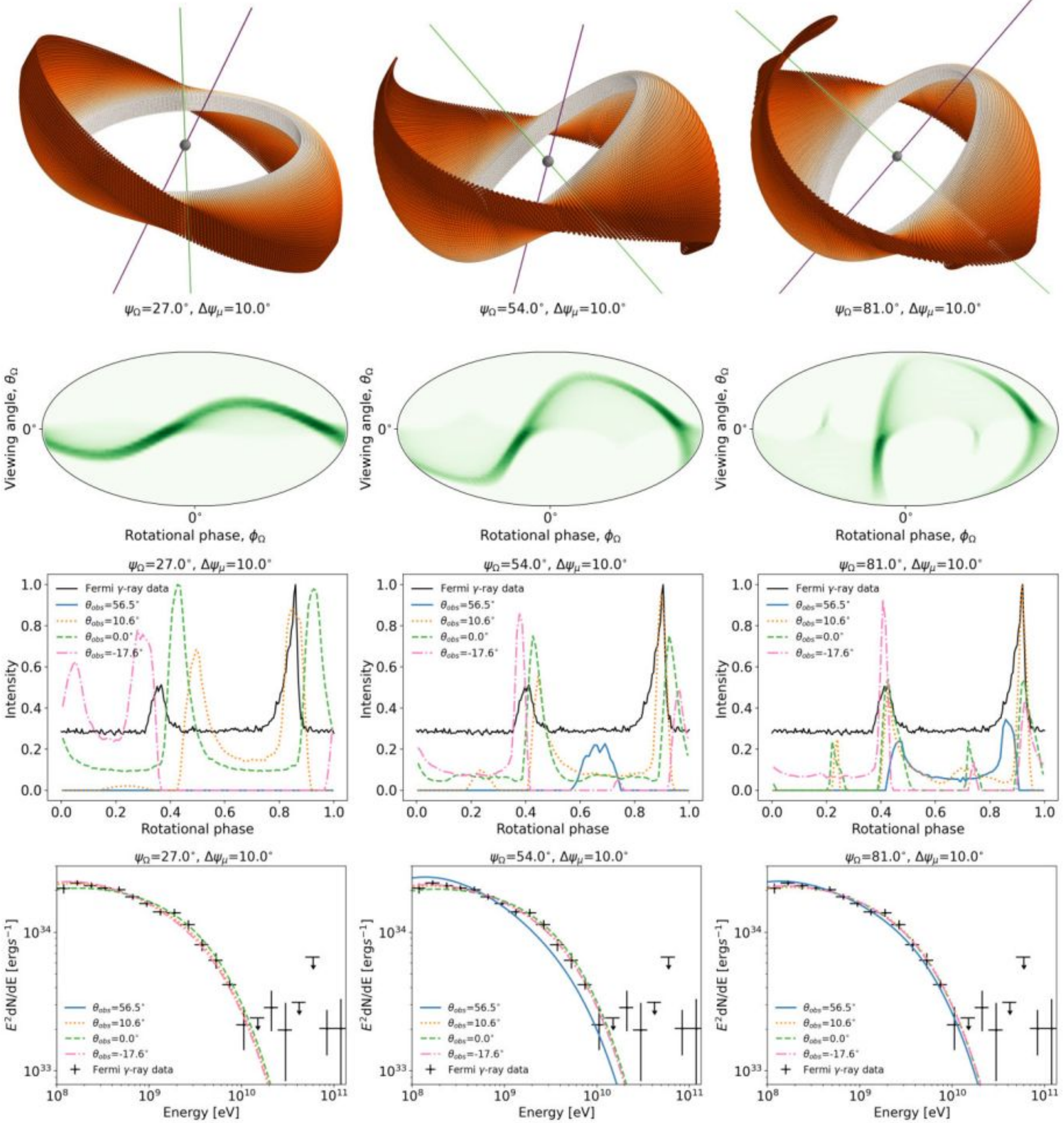
- The emission region is build around a centered value of the magnetic colatitude:

$$\Psi_\mu^c(\xi_\mu, R, \Psi_\Omega) = \frac{\pi}{2} + A(R, \Psi_\Omega)\sin(\xi_\mu - \pi/2)$$

with

$$A(R, \psi_\Omega) = K\ \Psi_\Omega\ (R/R_{lc} - R^0/R_{lc})^2$$

Back up slide: geometrical model



Back up slide: light curve fits

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