#### Understanding the AGN population: X-ray surveys

arcmin











**Chandra Deep Field-South (CDF-S)** 

≈7Ms Chandra exposure (last obs. at March 2016)

≈3Ms XMM-*Newton* exposure

Deep multi-wavelength coverage

One of the legacy fields (no deeper field for the next 20 yrs)

*Chandra*: good on-axis PSF (i.e., excellent angular resolution) and low background → Sensitive to faint and distant AGN

**XMM-***Newton*: larger effective area (hence photon statistics), but much worse angular resolution and higher background

→ Better for X-ray spectroscopy of relatively bright AGN

#### The deepest X-ray field: CDF-S



Capable of probing the high-z Universe with some photon statistics

#### **This Lab Outline**

- Build the source catalog: Produce a mosaic using 4 long CDFS exposures and provide source detections with different setups. Visualize the outputs and cross-match sources with the official 7Ms source catalog.
- **2. Explore the source catalog**: For one of the newly produced catalogs, produce some relevant plots, and compare quantities with those reported in the 7Ms source catalog
- **3. Analyse the data products**: Fit the X-ray spectra of a few, particularly interesting sources.

a. Reprocess with chandra\_repro four different Chandra observations of the CDF-S. Generate all data products (event files, exposure maps...) that are needed to perform a source detection using the merge\_obs tool.

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punlearn merge obs
pset merge_obs infiles=@infile.lis
pset merge_obs outroot=CDFS_4obs
pset merge_obs asolfiles=@asol.lis
pset merge_obs badpixfiles=@bpix.lis
pset merge_obs maskfiles=@mask.lis
pset merge_obs parallel=yes
pset merge obs nproc=4
pset merge obs units=time
pset merge_obs bands=broad
pset merge obs xygrid=0.5:8192.5:1,0.5:8192.5:1
pset merge_obs psfecf=0.9
pset merge_obs psfmerge=exptime
merge_obs
```

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- b. Run the wavdetect tool to search sources in your observations, using different significance thresholds (i.e., your detections can be more or less reliable) and different maximum wavelet scales (important if there are extended sources and for objects in the external part of the field).

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- c. Cross-correlate the source lists generated in the previous steps with the official 7 Ms Chandra source catalog in the CDF-S (Luo et al. 2017), using varius cross-matching radii.
  - Compute the fraction of 7Ms sources found in the 4-observation mosaic using different thresholds (1E-6/1E-5/1E-4)/scales (5.6/8/11)/matching radii (1/2/3").
  - For your source list which has the largest number matches within 2" with the 7 Ms CDF-S catalog, compute the number of sources detected in the 4-observation mosaic and not in the 7Ms catalog, and visualize them: what are the possible explanations for their detection in the your shorter-exposure mosaic?

Cross-correlate the source lists generated in the previous steps with the official 7 Ms Chandra source catalog in the CDF-S (Luo et al. 2017), using varius cross-matching radii (e.g., 1,2,3 arcsec)

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a. Choose one of the catalogs you built (e.g., the one with largest number of matches with the CDF-S 7 Ms one) and produce some plots (number of counts vs. source significance, vs. exposure time, vs. positional uncertainty, etc.)

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- a. Choose one of the produced catalogs and produce some plots (number of counts vs. source significance, vs. exposure time, vs. positional uncertainty, etc.)
- b. For the sources associated with the 7 Ms source catalog, produce the redshift distribution histogram, Lx vs. z plot, etc.
- c. Repeat the operation done in b. after creating subsamples of sources from the 7 Ms source catalog (e.g., spec-z vs phot-z; low vs high band-ratio...). Are there any noticeable trends?

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- Repeat the operation done in b. after creating subsamples of sources from the 7 Ms source catalog (e.g., spec-z vs phot-z; low vs high band-ratio...). Are there any noticeable trends?
- b. The trends can also be quantified using the Topcat statistics tool.

a. Repeat the operation done in b. after creating subsamples of

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TOPCAT(5): Row Statistics

Row Statistics for 5: CDFS\_7Ms\_catalog.fits

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Name	Mean	SD Minim	um Max		
VLA_DEC	-5.7216	11.2414	-27.9885		
VLA_20_CM_MAG	3.84106	7.57468	0.		
SPECT_REDSHIFT	1.0809	0.784943	0.034		
SPECT_REDSHIFT_FLAG			INSECURE		
REF_SPECT_REDSHIFT	10.6576	6.54157	2		
PHOT_REDSHIFT_L10	0.542863	0.848864	0.		
PHOT_REDSHIFT_R11	1.03203	0.749643	0.		
PHOT_REDSHIFT_H14	1.07511	0.787236	0.		
PHOT_REDSHIFT_S14	0.82387	0.80083	0.		
PHOT_REDSHIFT_S15	0.809108	0.814808	0.		
PHOT_REDSHIFT_S16	0.936187	0.826658	0.		
REDSHIFT	1.08991	0.776239	0.038		
REF_REDSHIFT			H14		
REDSHIFT_NEG_ERR	0.002921	0.02576	0.		
REDSHIFT POS FRR	0.00354	0.026326	0.		
		All			
Subset for calculations: spec-z					
		phot-z			

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- b. For the sources associated with the 7Ms source catalog, produce the redshift distribution histogram, Lx vs. z plot, etc.
- c. Use the PIMMS Online tool (<u>https://cxc.harvard.edu/toolkit/</u> <u>pimms.jsp</u>) to compute the count rate-to-flux correction factor, using the photon index available in the catalog.

#### Lab Outline

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PIMMS v4.11a: with ACIS Pile up and Background Count Estimation

	Input	Output	
Ł	<u>Count Rate</u> <u>Flux</u> <u>Flux Density</u>	○ <u>Count Rate</u> ● <u>Flux</u> ○ <u>Flux Densi</u>	<u>ity</u>
(	Mission:Detector/Grating/Filter:CHANDRA-Cycle 11 ACIS-I/None/None	✓ Flux: Absorbed ✓	
	Input Energy: 0.5 to 2 keV	Output Energy: 0.5 to 2	
N F	Model:     Galactic NH:     Redshift(z):     R       Power Law     7E19     0     0       cm**-2     c	Redshifted NH:Photon Index:Count H01.71E-2cm**-2N=AE**-acts/s	<u>Rate</u>
	CALCULATE CLE	EAR HELP	
	PIMMS Predict 6.383E-14 erg/cm**2/s abso	ction: osorbed flux	

#### Lab Outline

#### 9) Evolara tha course astalad

PIMMS v4.11a: with ACIS Pile up and Background Count Estimation

8		Input			Output	
k	● <u>Count Rate</u> ○	<u>Flux</u> O <u>Flux D</u>	<u>ensity</u>	O <u>Count Ra</u>	ate • Flux •	Flux Density
(	Mission: CHANDRA-Cycle 11 ✓	What happens	changing the	e mission (	Cycle?	
	Input Energy: 0.5	to 2 k	eV	Output Ene	<u>rgy:</u> 0.5	to 2
<b>N</b>	<mark>∕Iodel:</mark> Power Law ✓	Galactic NH:Red7E190cm**-2	lshift(z): Reds	hifted NH: -2	Photon Index: 1.7 N=AE**-a	Count Rates
		CALCULA	TE CLEAR	HELP		
		PIMM 6.3 erg/cr	MS Prediction: 83E-14 n**2/s absorbed	l flux		

#### 3. Analyse the data products: spectral fitting

Fit *Chandra* spectra for sources whose properties suggest potential interesting outcome (e.g, high-z, high obscuration based on hardness ratio...).

XID_Luo17	Source coordinates	Z	Opt. Class + Info
551	03:32:29.85 -27:51:05.71	3.700	NL (Comastri+11)
746	03:32:39.66 -27:48:50.64	3.064	NL (Vito+13)
730	03:32:38.91 -27:57:00.48	0.298	NL
242	03:32:13.24 -27:42:40.96	0.605	NL

IDs reported in the spectral files we provide

All spectra and response matrices are provided

#### 3. Analyse the data products: spectral fitting

#### **Spectral analysis pipeline**

- 1. Choose one of the four sources
- 2. Group the spectra (grppha) accordingly to the quality of the data
- 3. Load spectra in XSPEC
- 4. Define a spectral model and fit it to the data
- 5. Once a physically justified model is obtained, save the X-ray spectral parameters (including errors) and produce confidence contours
- Check for further components (to lower the data/model residuals) Return to point 3

#### PLAN (III)

#### **OPTIONAL**

a. Re-run the procedure for a second source, better if at a different redshift range.



#### **Main publications**

- Xue Y.Q. et al. 2011, ApJS, 195, 10 4 Ms Chandra source catalog.
- Vito F. et al. 2013, MNRAS, 428, 354 High-redshift AGN population in the CDF-S.
- Luo B. et al. 2017, ApJ Suppl., 228, 2 The Chandra Deep Field-South

Survey: 7 Ms Source Catalogs.

#### **Command list: merge\_obs**

punlearn merge\_obs pset merge\_obs infiles=@infile.lis pset merge\_obs outroot=CDFS\_4obs pset merge\_obs asolfiles=@asol.lis pset merge\_obs badpixfiles=@bpix.lis pset merge\_obs maskfiles=@mask.lis pset merge\_obs parallel=yes pset merge\_obs nproc=4 pset merge\_obs units=time pset merge\_obs bands=broad pset merge\_obs xygrid=0.5:8192.5:1,0.5:8192.5:1 pset merge\_obs psfecf=0.9 pset merge\_obs psfmerge=exptime merge obs

#### **Command list: wavdetect**

punlearn wavdetect pset wavdetect infile=CDFS\_4obs\_merged\_057keV\_bin1.fits pset wavdetect outfile=CDFS\_4obs\_merged\_057keV\_wavdet\_1em6\_src.fits pset wavdetect scellfile=CDFS\_4obs\_merged\_057keV\_wavdet\_1em6\_cellimage.fits pset wavdetect imagefile=CDFS\_4obs\_merged\_057keV\_wavdet\_1em6\_reconstructed.fits pset wavdetect defnbkgfile=CDFS\_4obs\_merged\_057keV\_wavdet\_1em6\_normbakg.fits pset wavdetect regfile=CDFS\_4obs\_merged\_057keV\_wavdet\_1em6.reg pset wavdetect ellsigma=3.0 pset wavdetect sigthresh=1e-6 pset wavdetect scales="1 1.4 2 2.8 4 5.6 8 11" pset wavdetect expfile=CDFS\_4obs\_merged\_broad\_thresh.expmap pset wavdetect psffile=CDFS\_4obs\_merged\_broad\_thresh.psfmap wavdetect clobber+ verbose=3