Constraining Blazar emission models with the study of Waiting time between gamma-ray flares of FSRQ

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Flare recognition

- With a clustering method applied to gamma-ray collected within a suitable region around the source, we searched for gamma-ray flares of FSRQs in the energy range **0.1-300 GeV**.

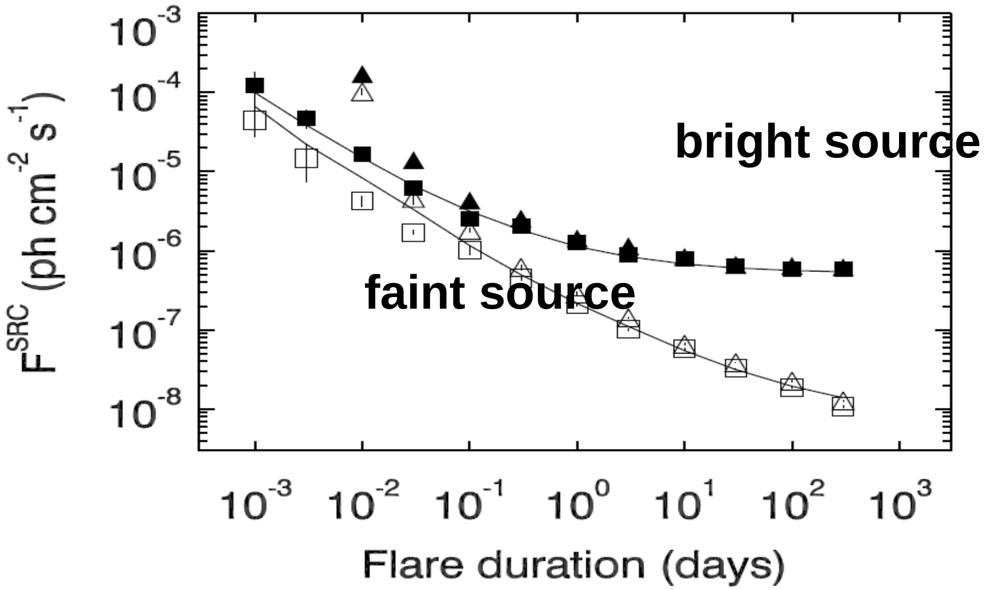
- **data set**: $\{X_{(i)}\}$ (gamma-ray events collected within an extraction region)
  where $X_{(i)}$ is the cumulative exposure (from the start of obs) of the collected event i.

- **clustering law**:

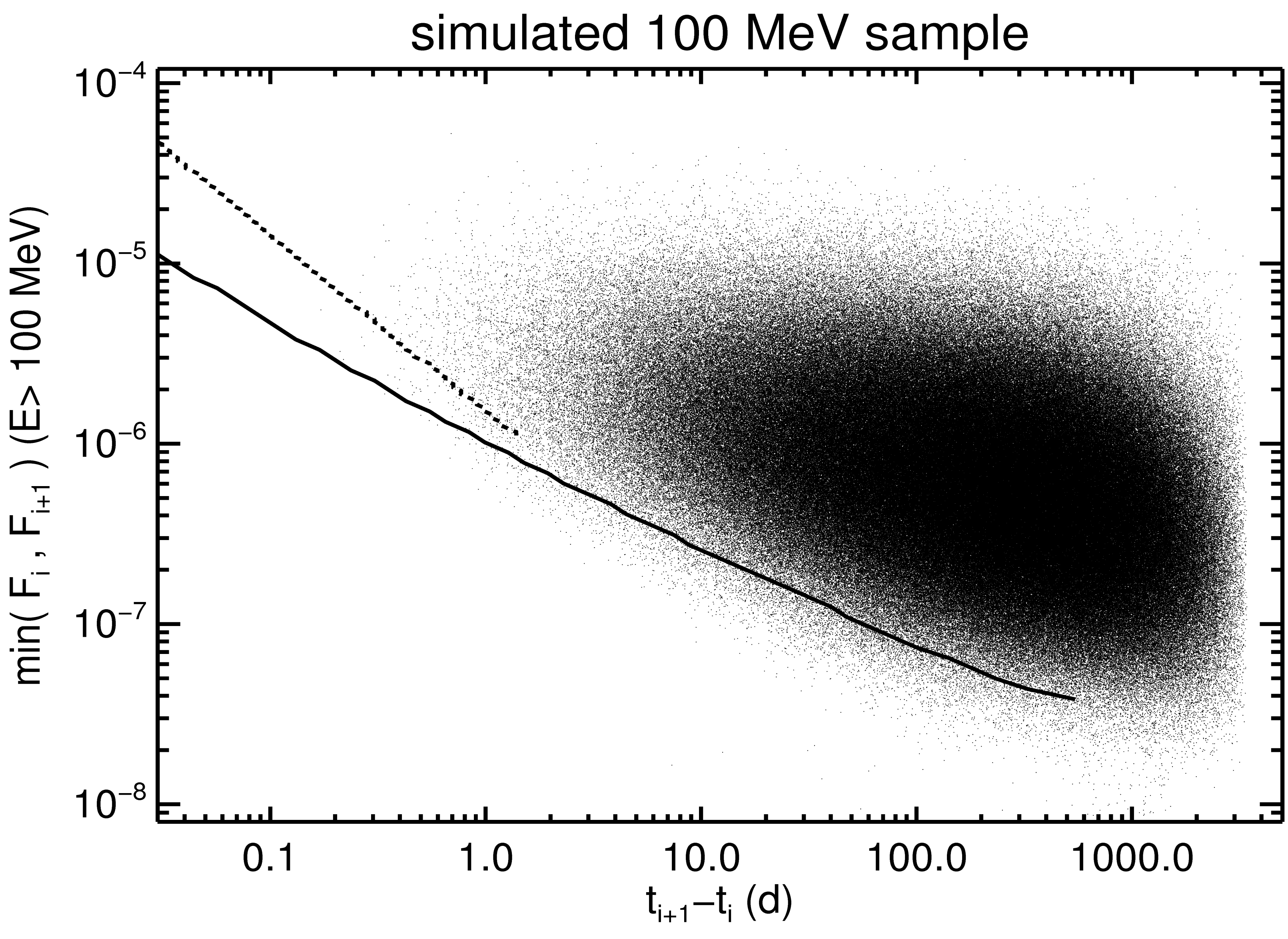
$$\begin{cases} X_{(i+k)} - X_{(l)} < k * \Delta_{thr} & (K < N_{tol}) \\ l \in [i, i+k] \end{cases}$$

- chance cluster probability is evaluated with a scan statistic related method (**maximum score scan statistic**, Glaz 2006, **conf. level set to $1*10^{-3}$)**.

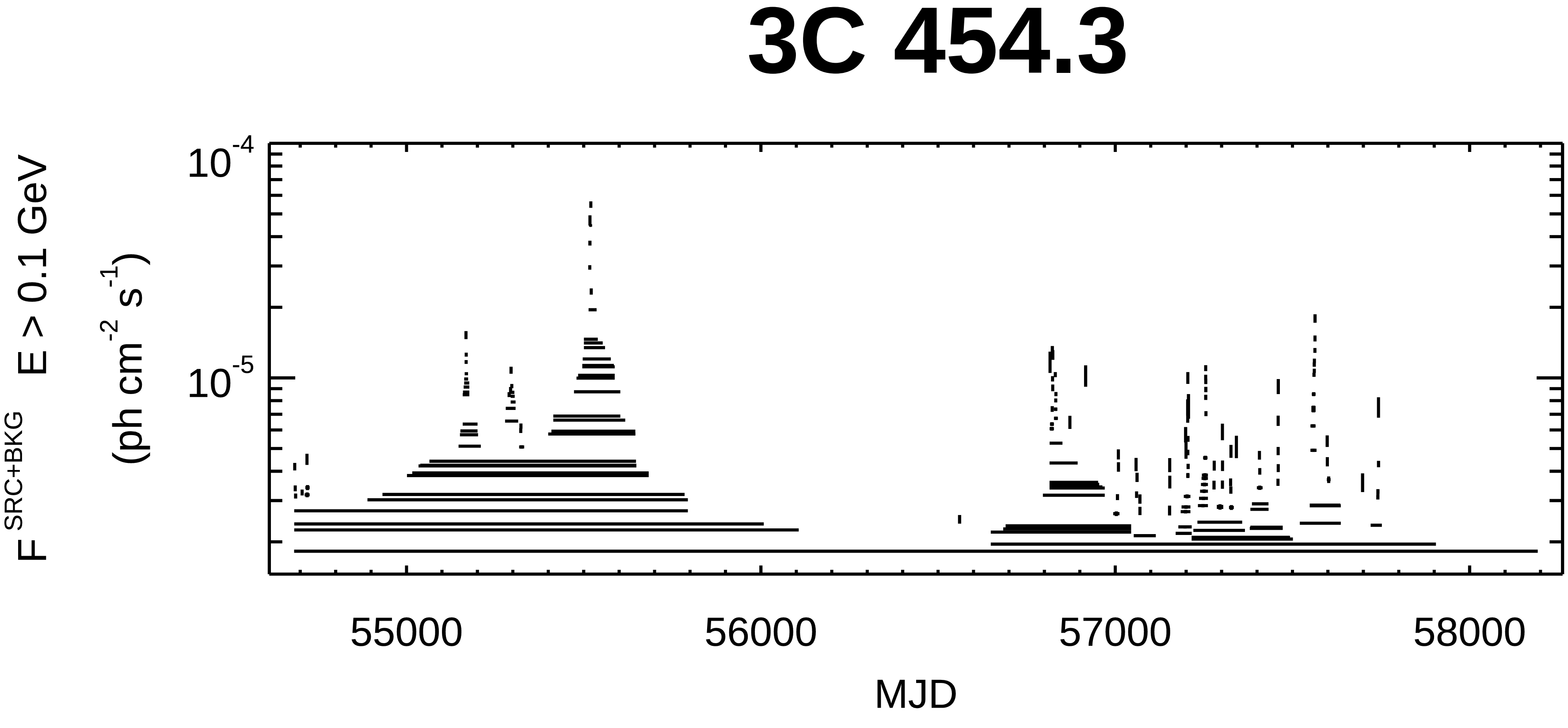
iSRS sensitivity



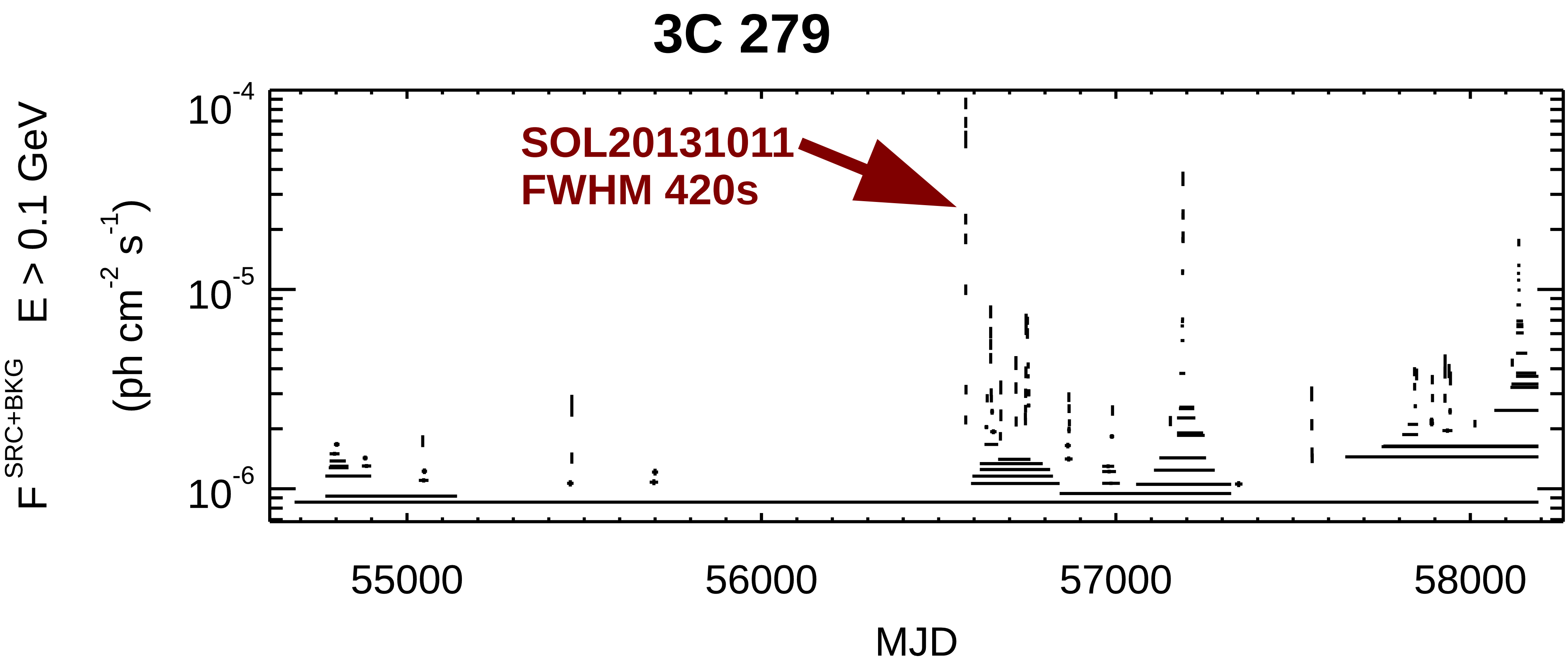
iSRS resolving power



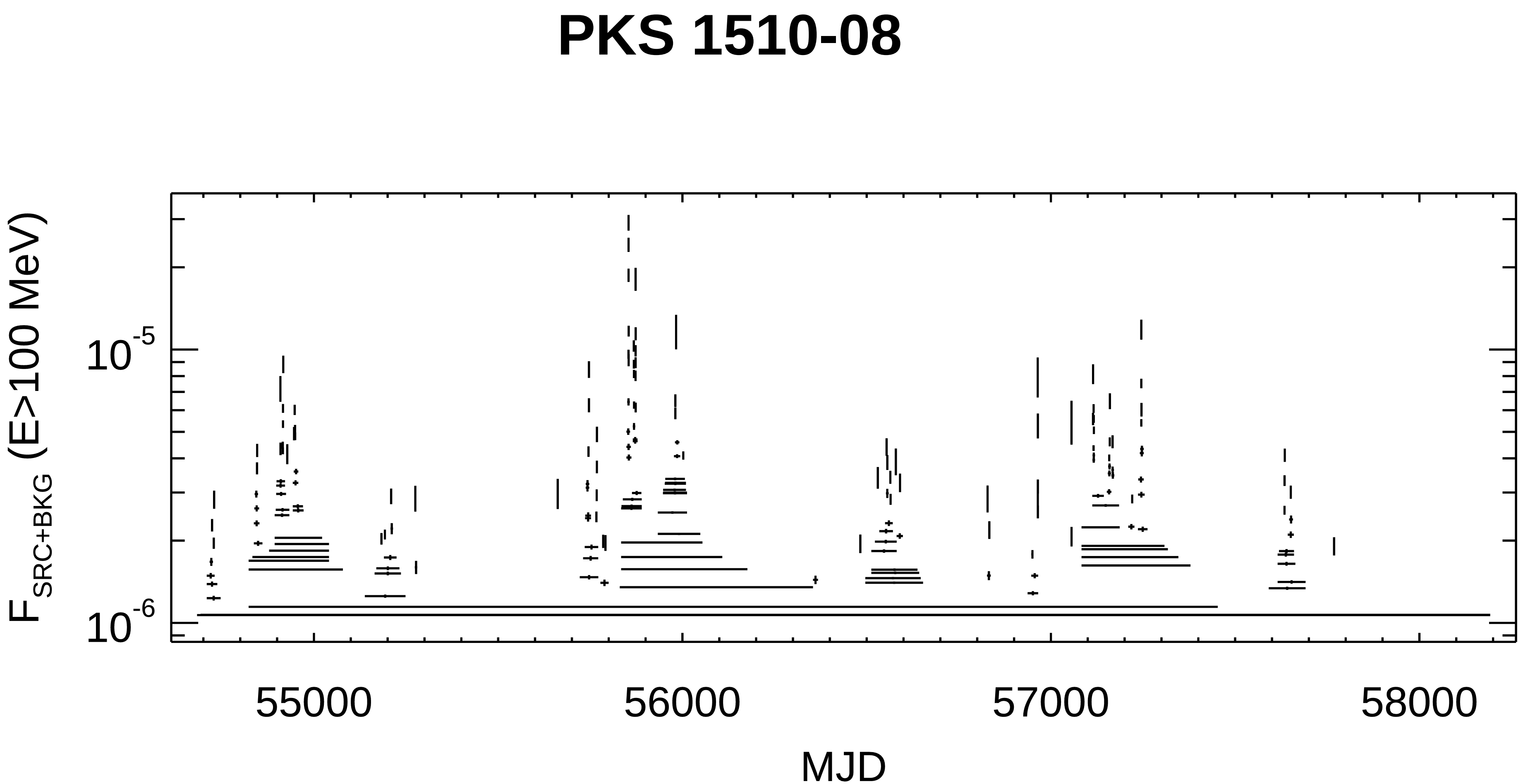
Unbinned light curves (NSIGMA=3)



Unbinned light curves (NSIGMA=3)



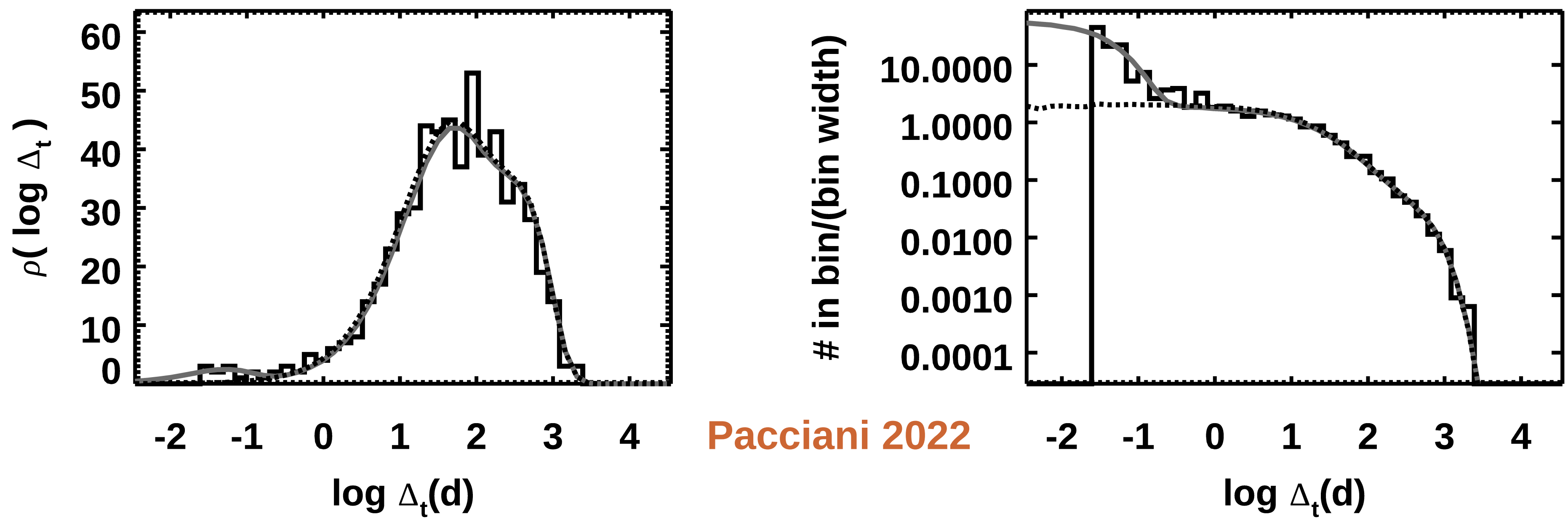
Unbinned light curves (NSIGMA=3)



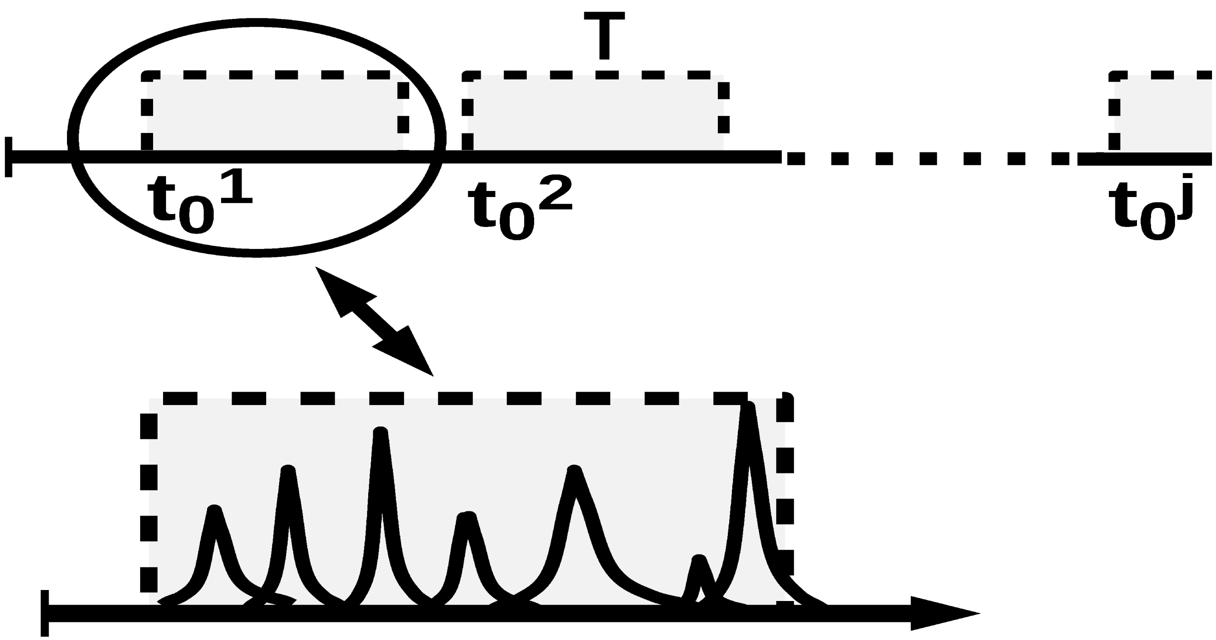
Sample selection

- FSRQs from the 3FGL catalog with TS > 49
- Only sources at least 15^0 above the galactic plane are chosen
- Peak fluxes evaluated with **photometry** are **compared with the full likelihood** analysis, and eventually the photometric flare is validated (TS flare > 9, Ratio of photomeric to likelihood based flare flux < 1.3)
- No other selections applied (e.g., Sun constrain)
- Comparison with simulations of time series taking into account of exposure variation with time
- esposure for each source evaluated with time bins of 1/1000 of a day
- 649 (713) flares from 147 (115) sources for E > 300 MeV (E> 100 MeV)

Waiting time distribution (all flares)



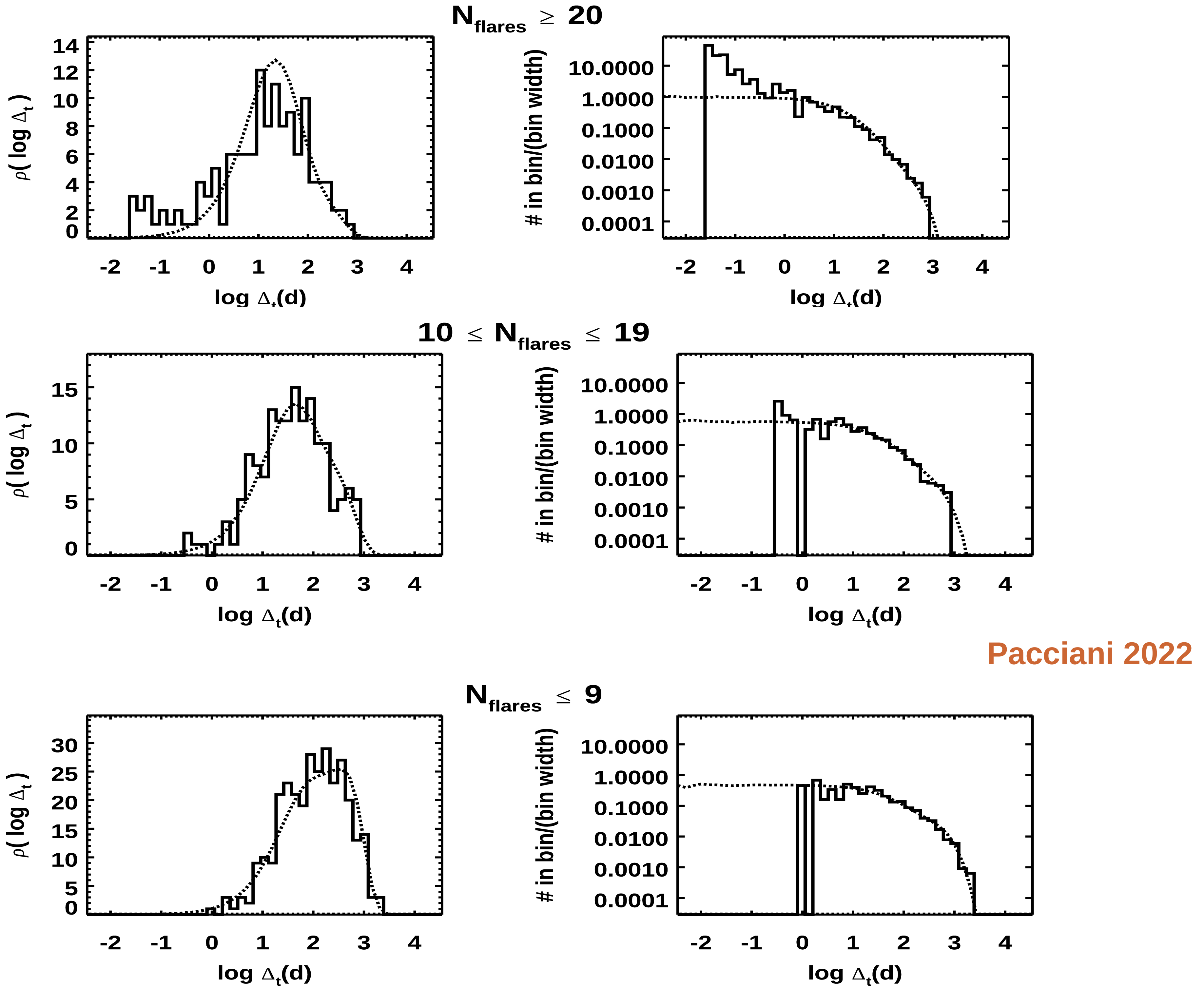
Pacciani 2022



overlapping bursts of flares with:

- burst duration **T ~0.6y**
- **burst rate 1.3y^{-1}**
- **flares uniformly distributed** within the burst

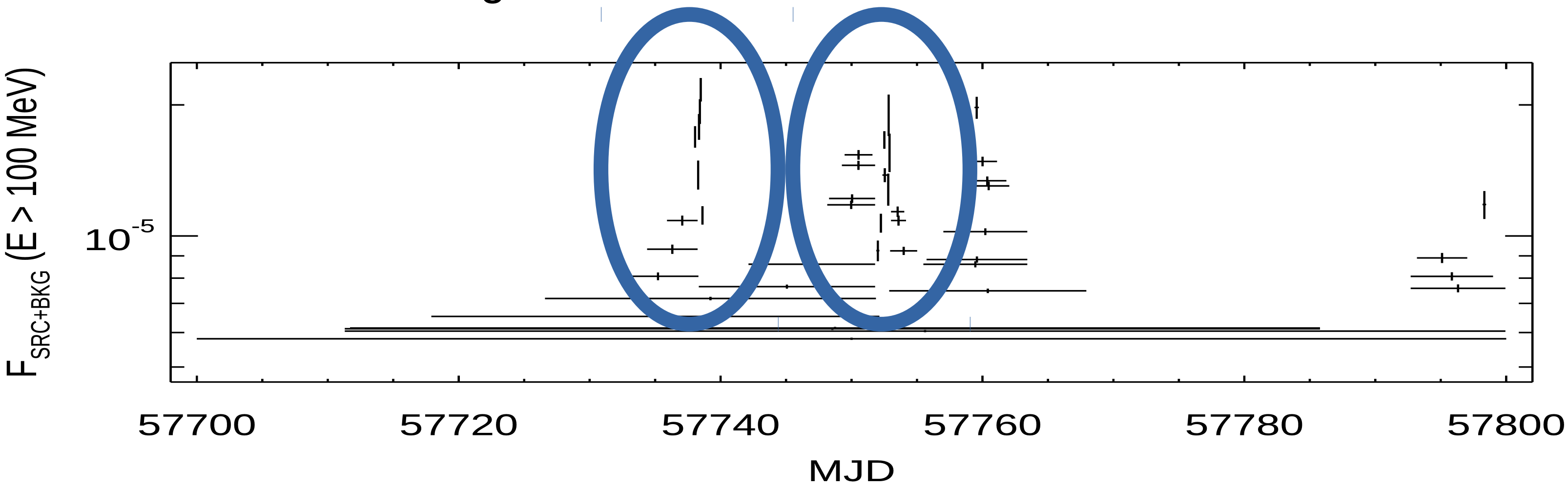
moreover a **fast component for Δt<1d**



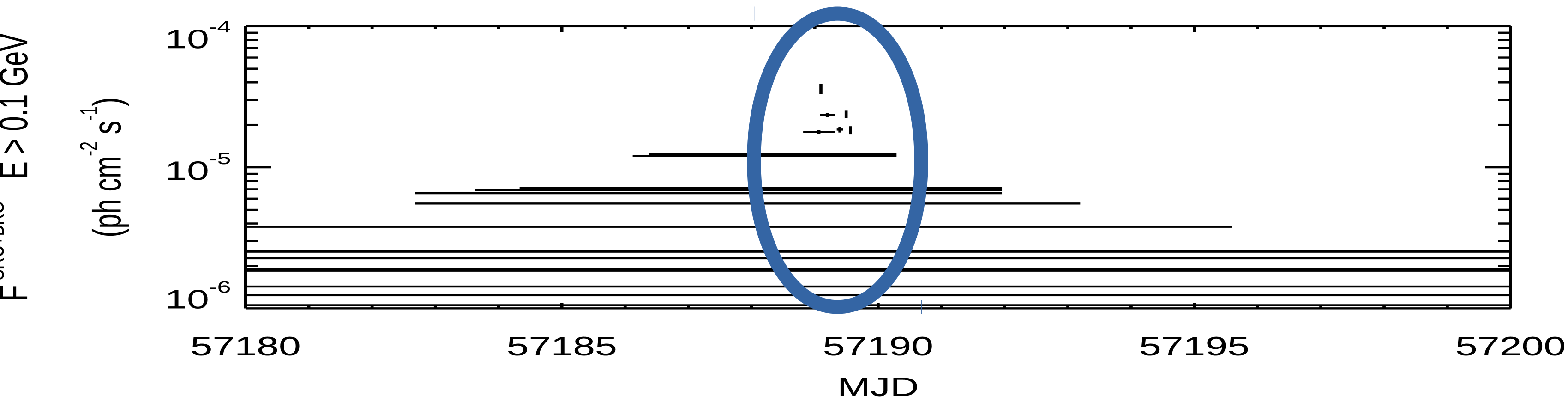
$N_{flares} \geq 20$
$10 \leq N_{flares} \leq 19$
$N_{flares} \leq 9$
ρ(log Δt)
in bin/(bin width)
log Δt(d)
Pacciani 2022

Short Waiting times and Fermi-LAT ToO

Apart from S4 0218+35, FERMI-LAT performed 5 ToO Campaigns on FSRQs during bright flares, and during 3 of theese campaigns we found the large majority of short waiting times:

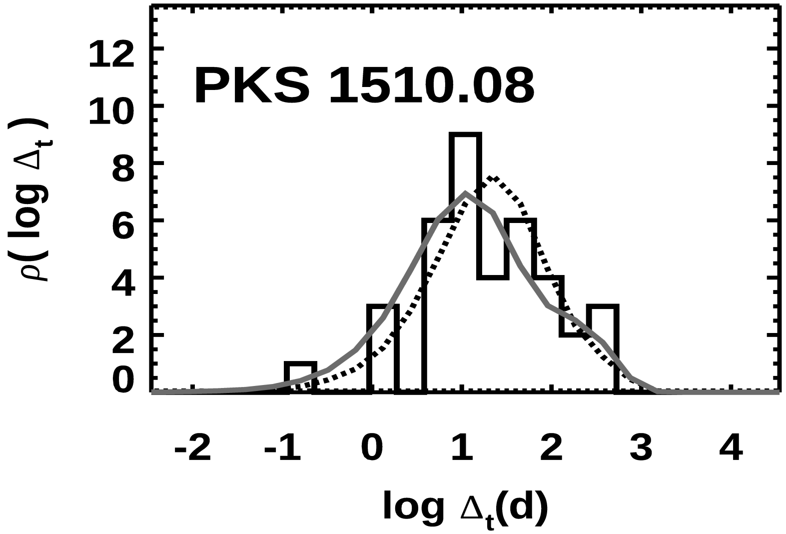


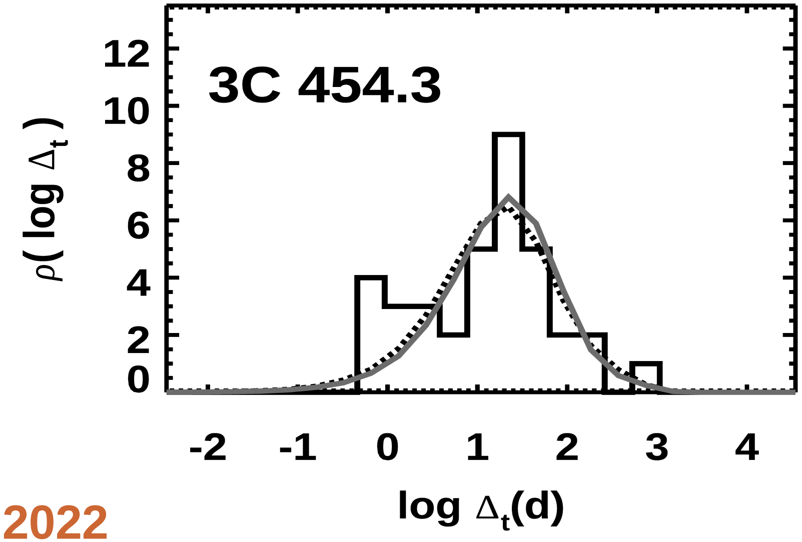
CTA 102



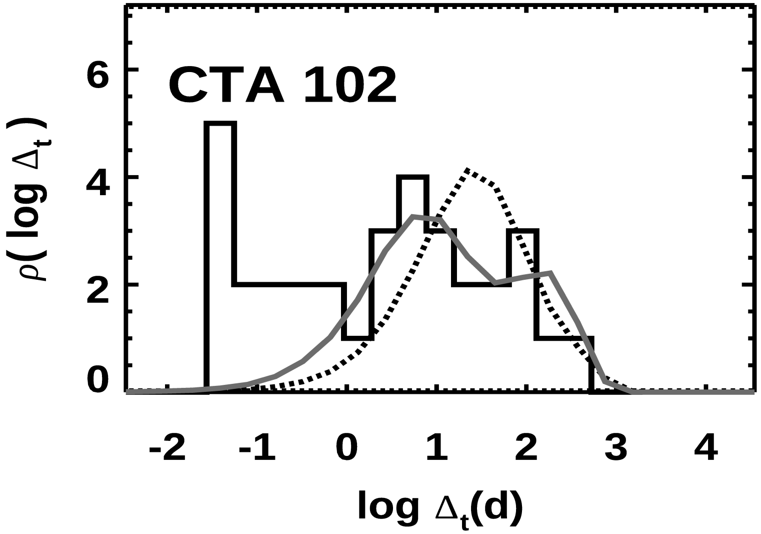
3C 279

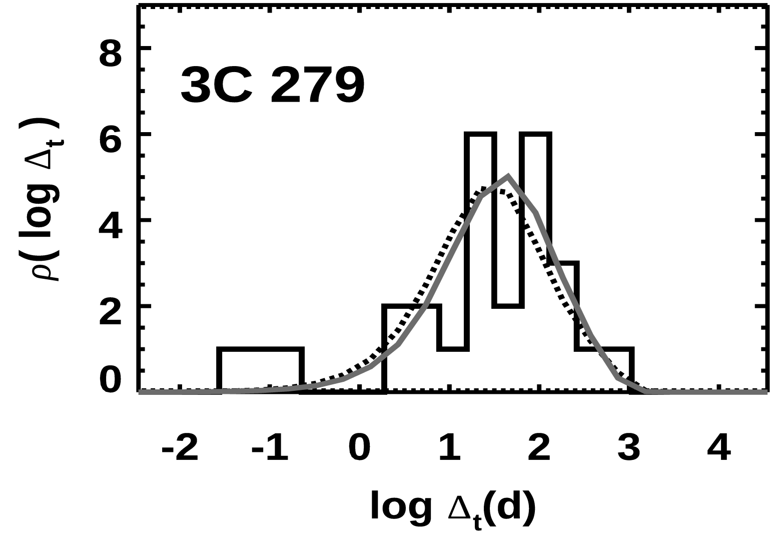
Fit for single sources



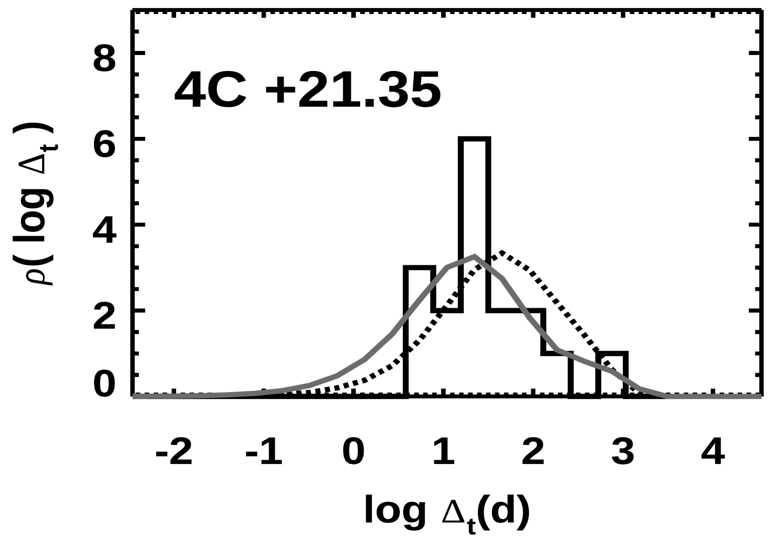


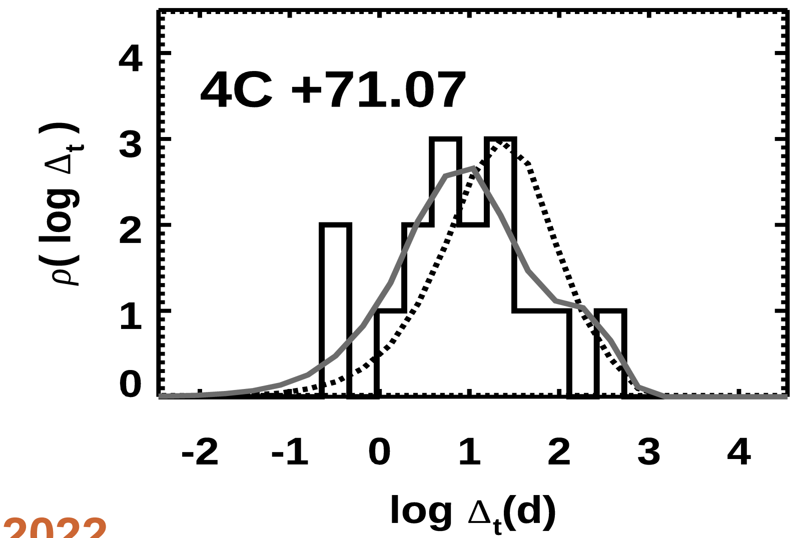
Pacciani 2022



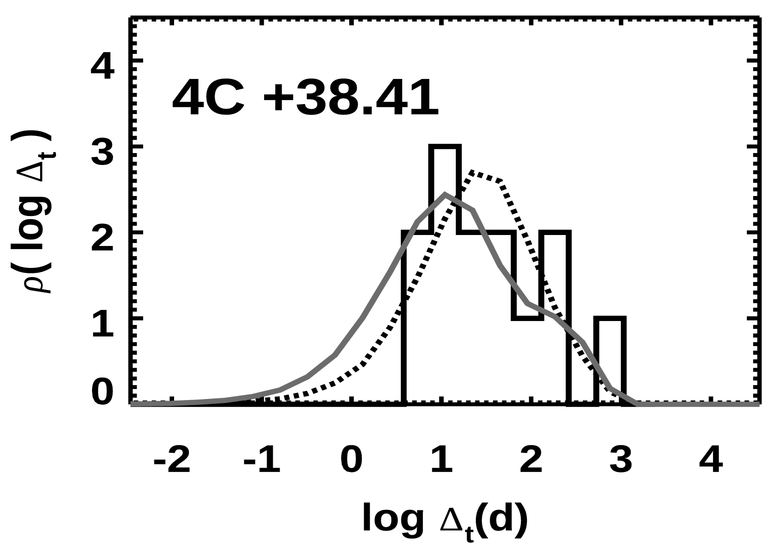


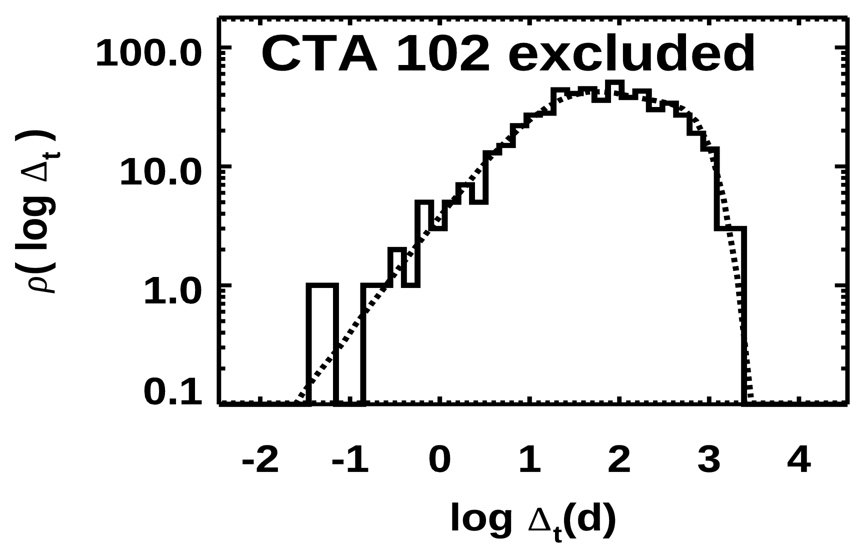
Fit for single sources



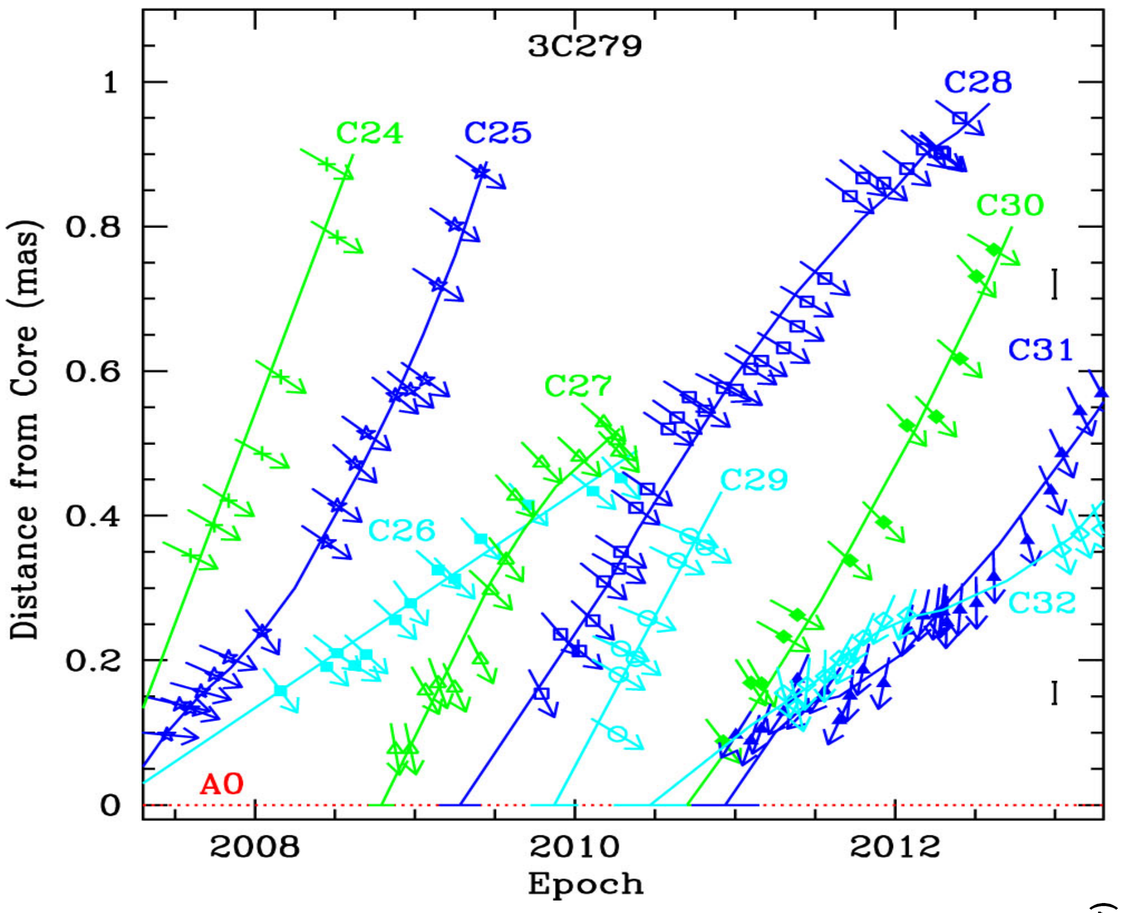


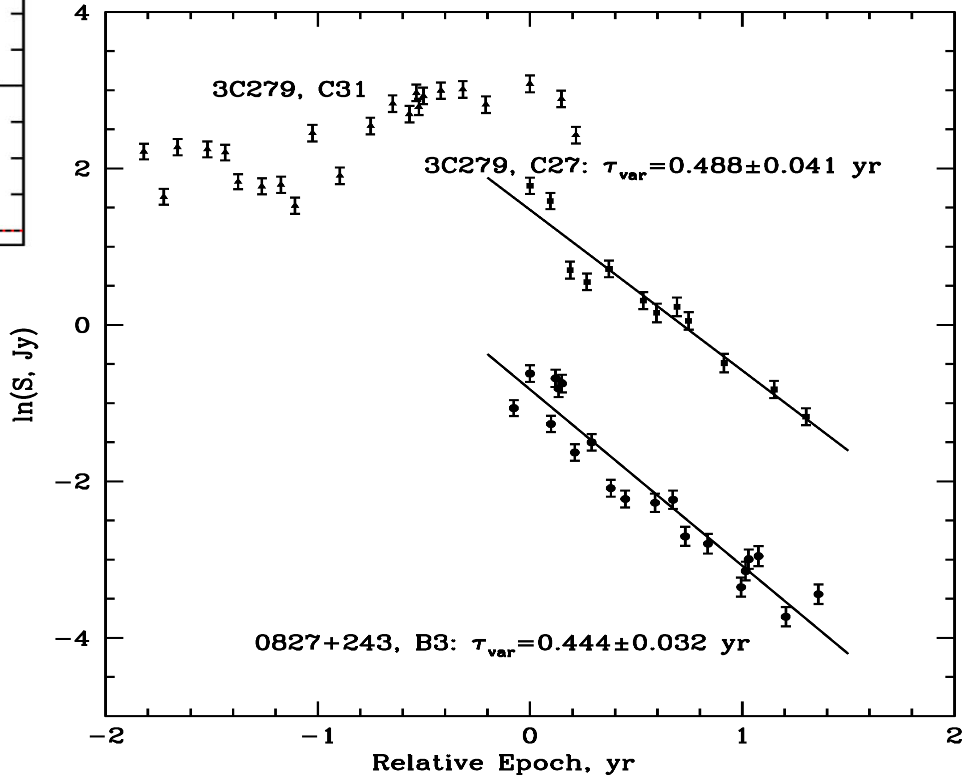
Pacciani 2022





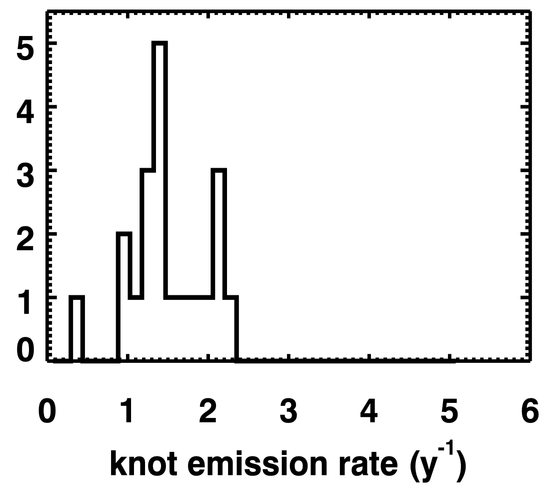
Radio data @ 43 GHz (Jorstad 2017)



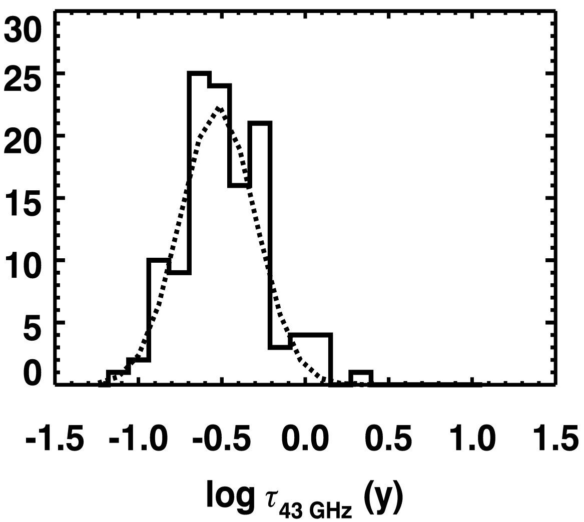


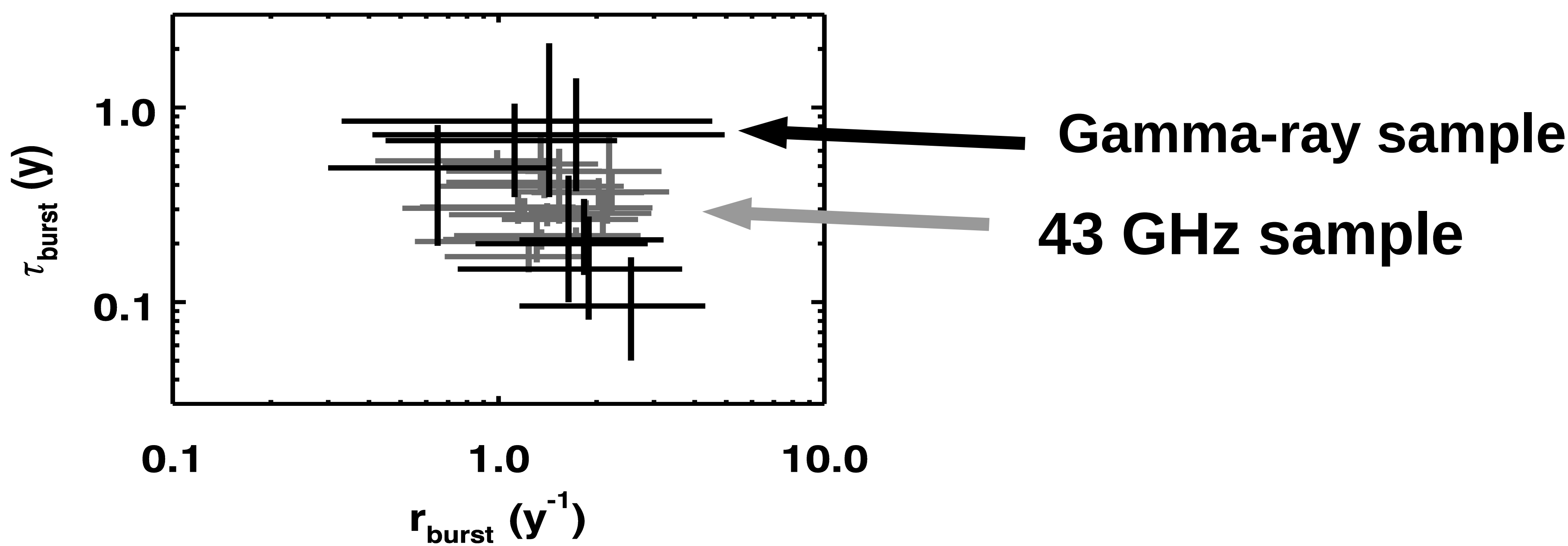
But we expect that the synchrotron power to decrease with 1/d.
Is the plasma accelerating along its path?

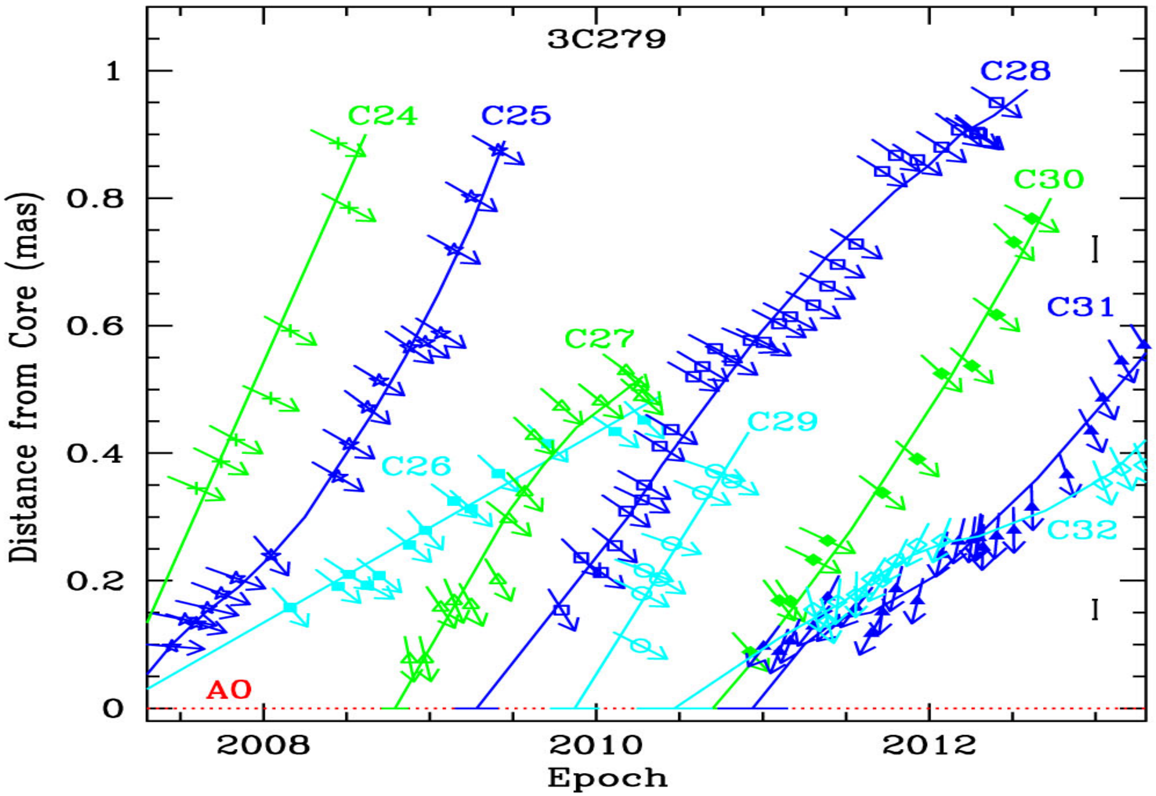
Comparison with Radio data



from Jortad 2017, table 7, and fig. 4



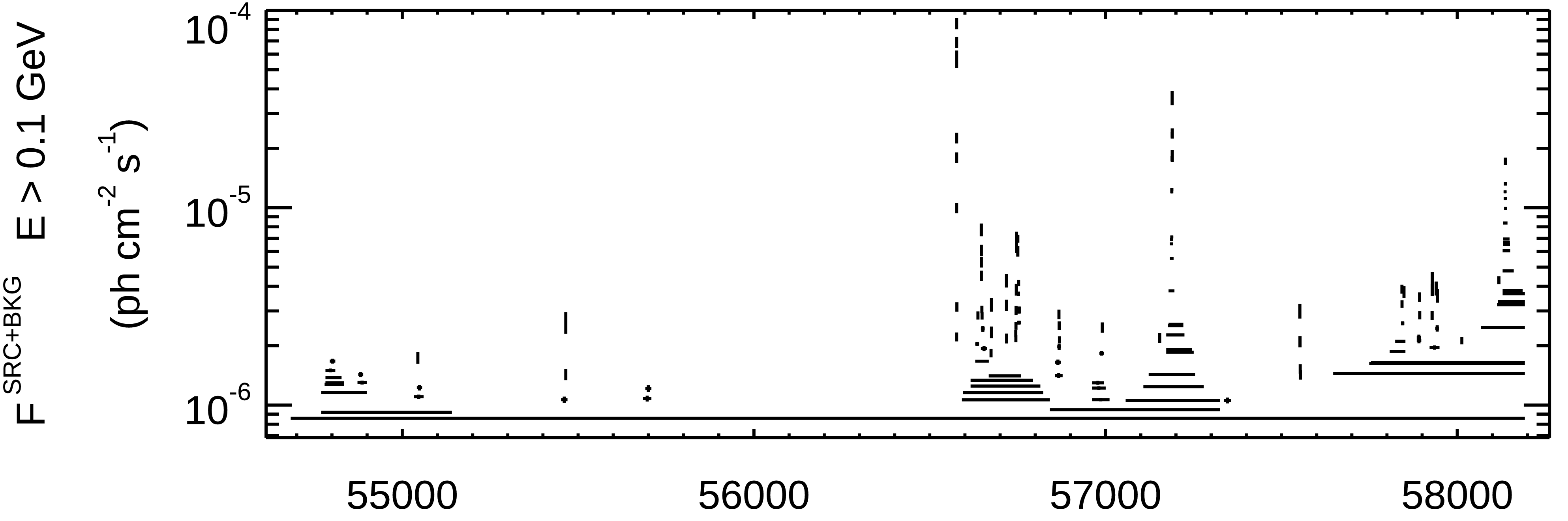


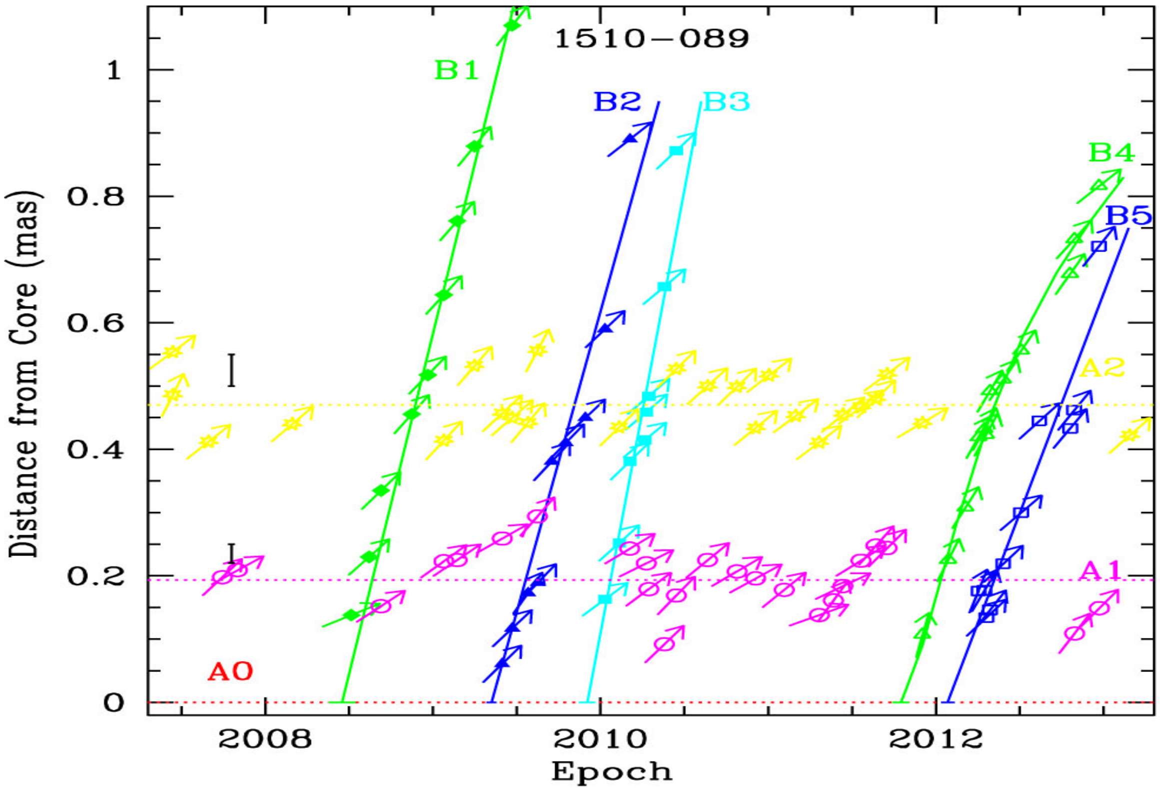


Radio data @ 43 Ghz and Gamma Ray data for 3C 279

2009-01-01: MJD:54832
2013-01-01: MJD 56293

Horizontal scales are aligned in radio and in Gamma-ray



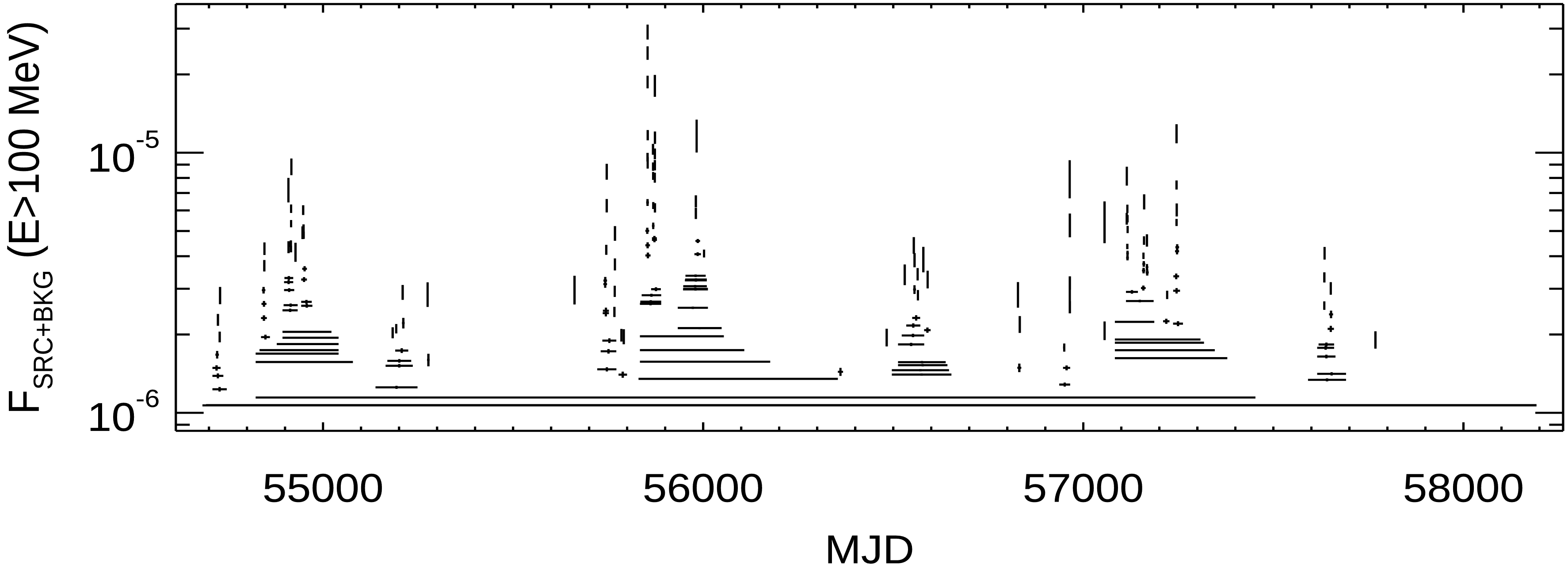


Radio data @ 43 Ghz and Gamma Ray data for PKS 1510-08

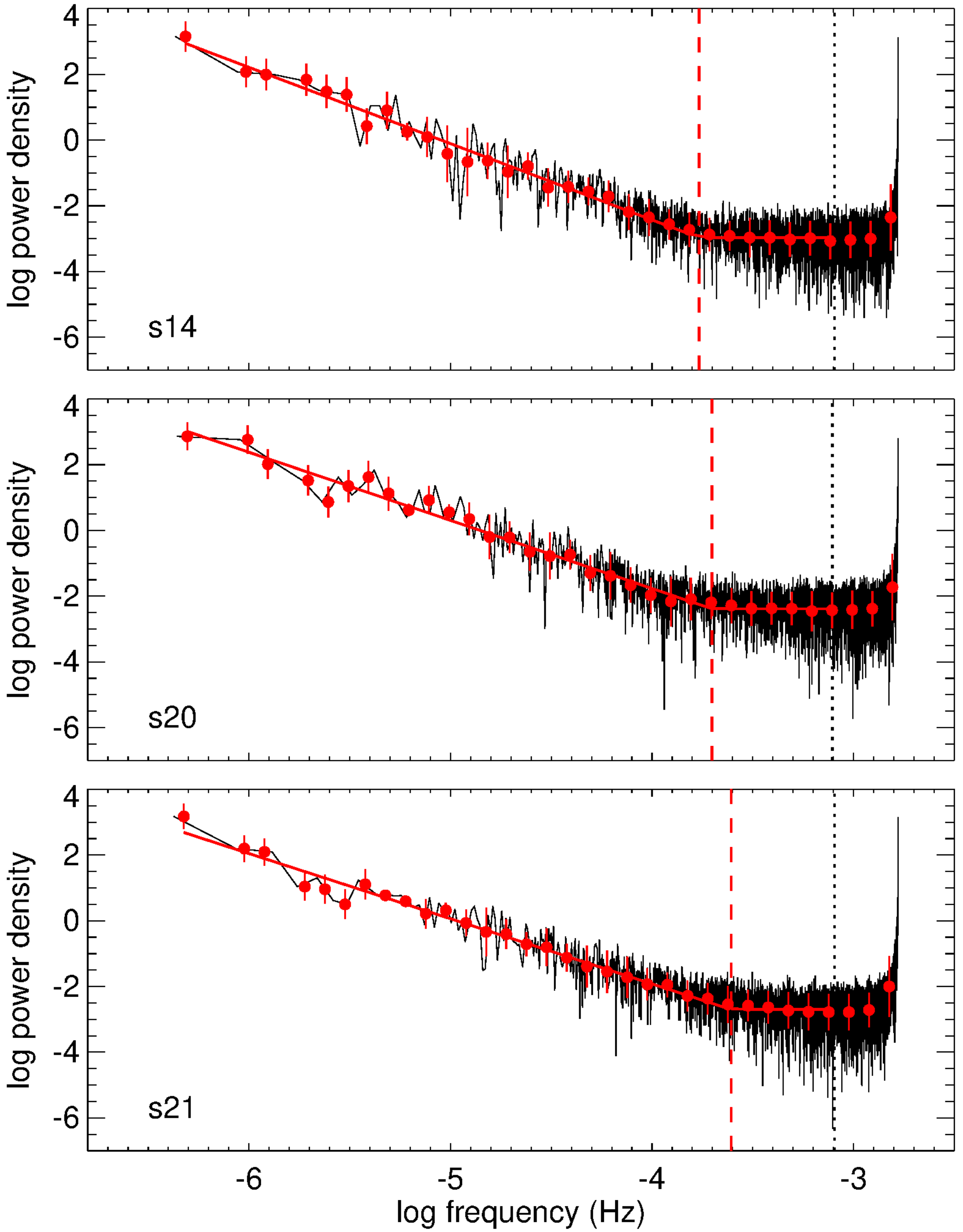
2009-01-01: MJD:54832

2013-01-01: MJD 56293

Horizontal scales are aligned in radio and in Gamma-ray



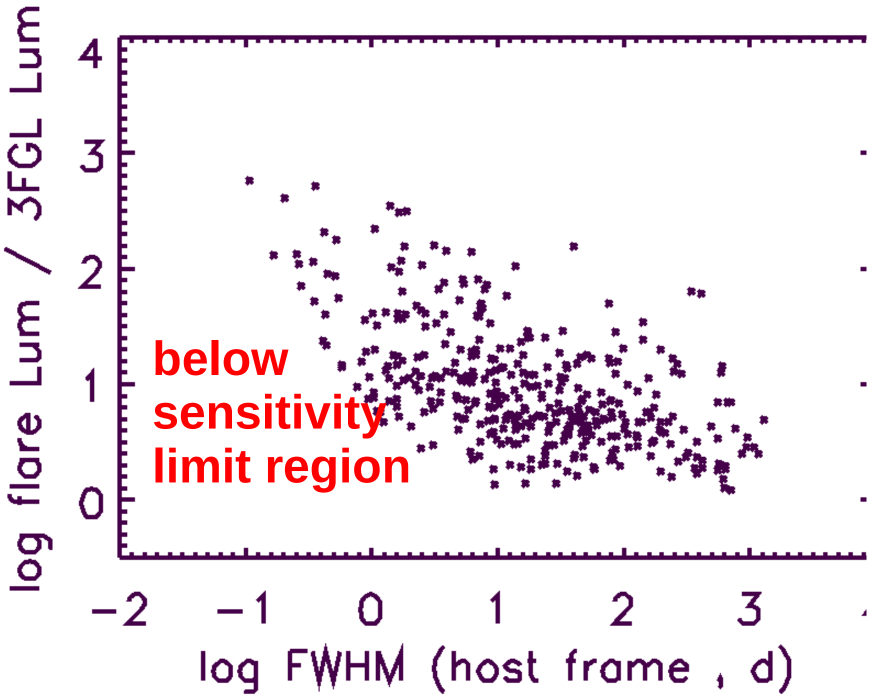
comparison with optical

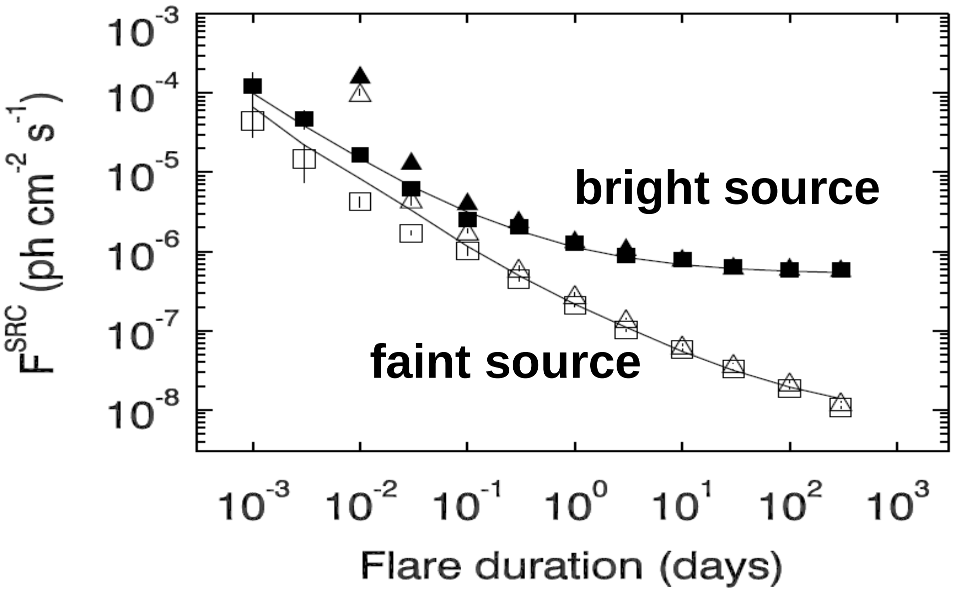


S4 0954+65
(a BL Lac object)
Raiteri2021

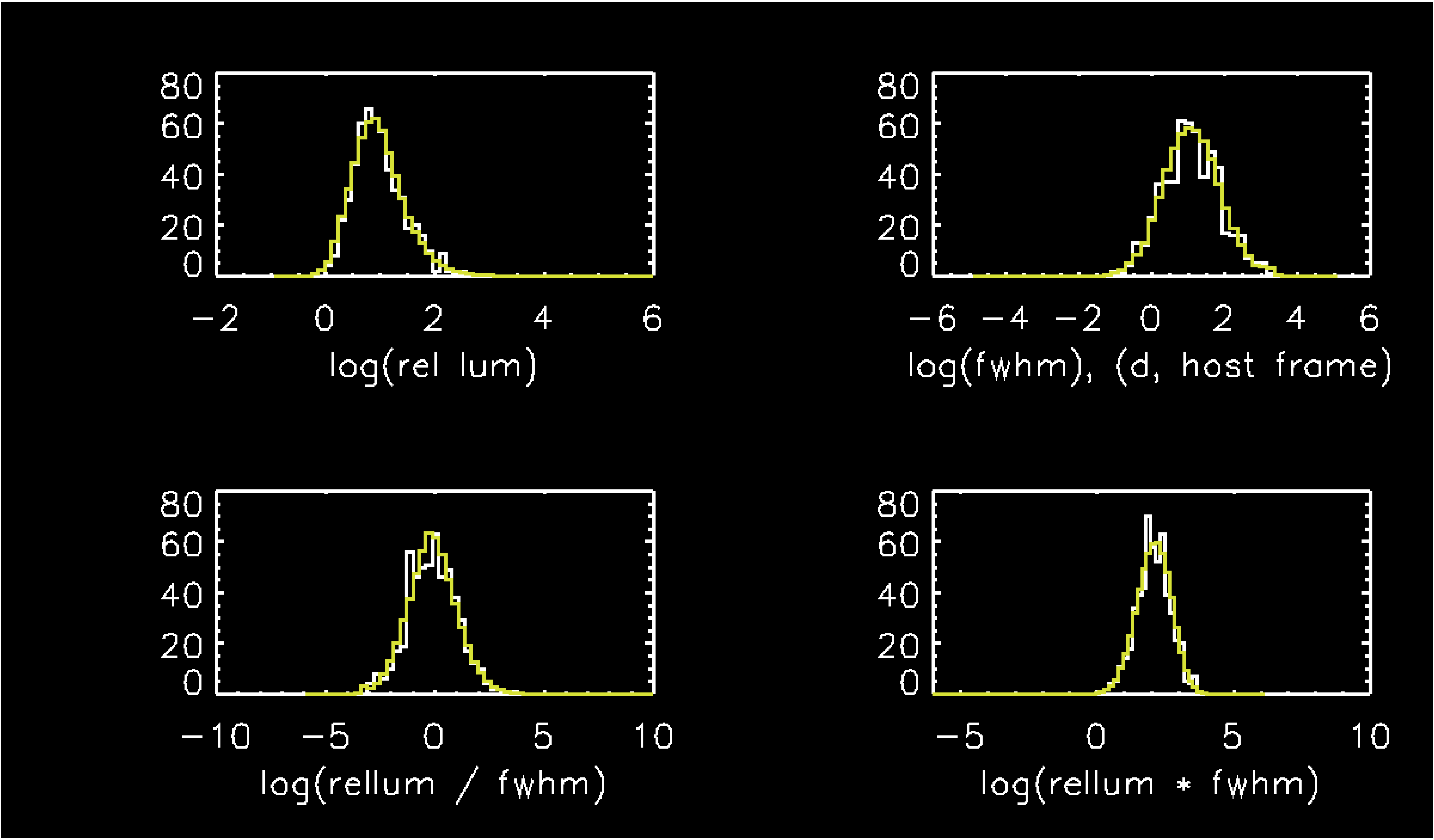
White noise appears for timescales < 1h

Flaring Luminosity
Vs
flare temporal
FWHM

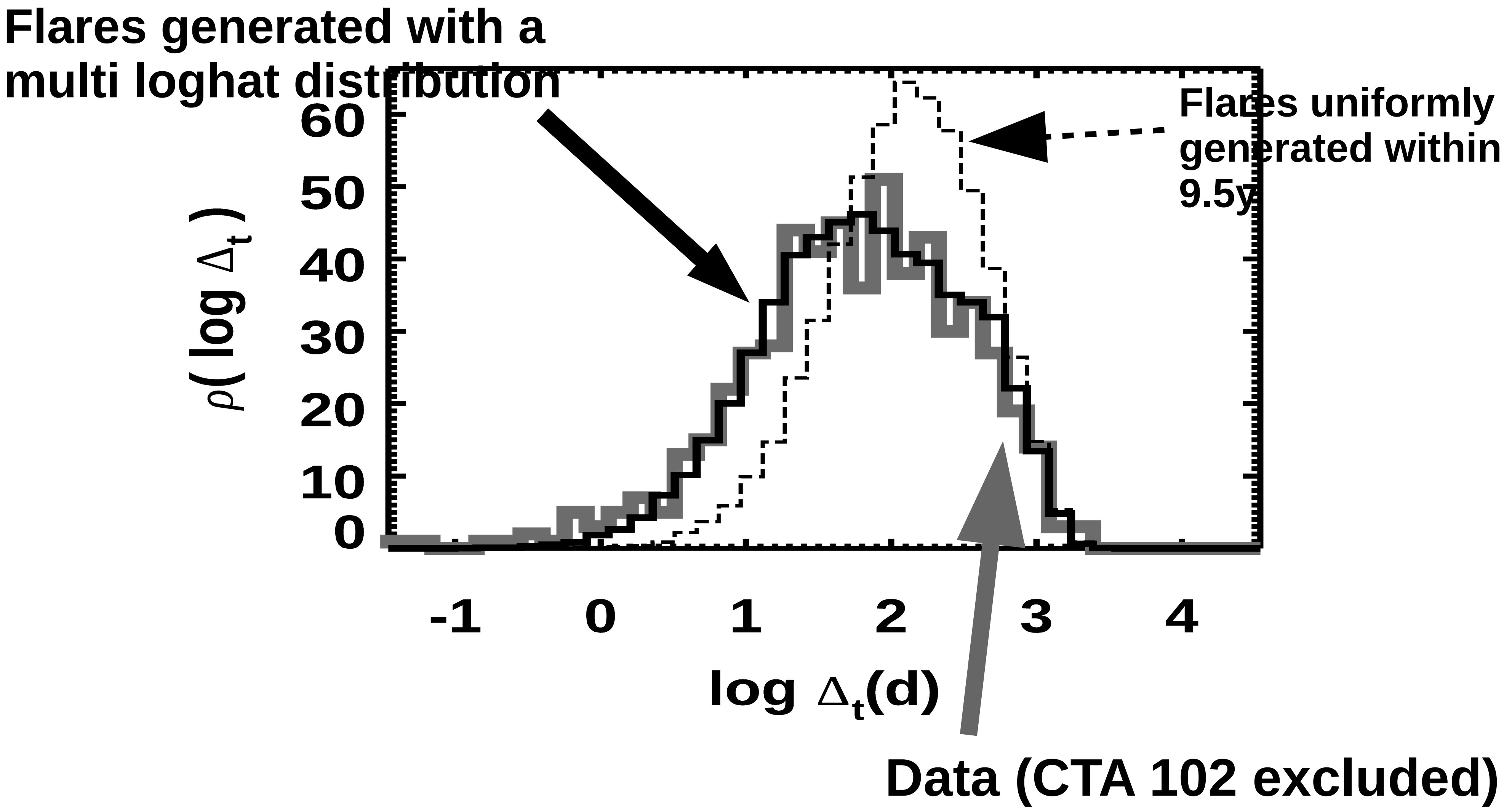




Flare Luminosity and duration distributions (300 MeV sample)



Fit taking into account for the exposure variation with time and simulating flares



recollimation shock scenario

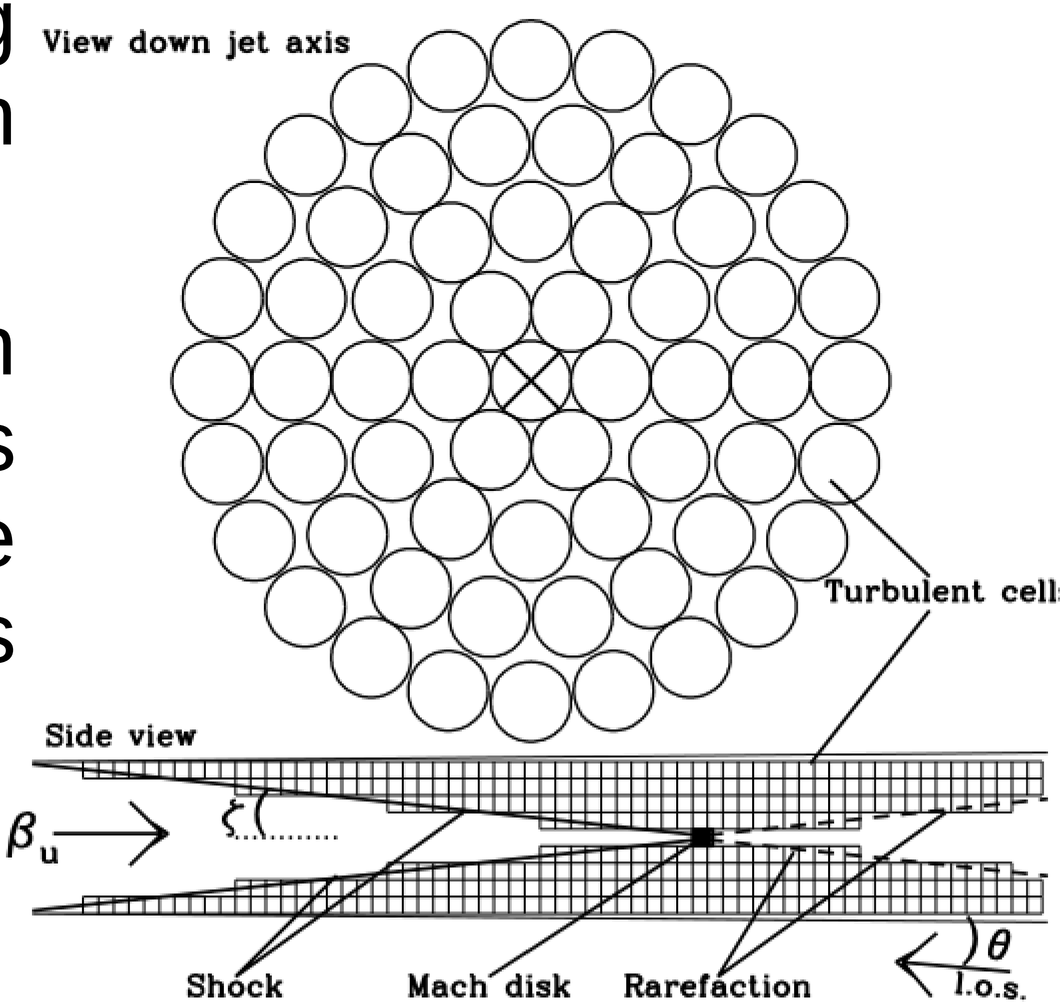
- If gamma-ray flares are produced while the traveling knots cross jet stationary features (recollimation shock scenario, Casadio2015, Casadio2019), from the burst period, knot size along the jet should be ~2pc (knot reference frame).

Turbulence in the jet

electron acceleration is caused by standing conical recollimation shocks.

Flux and polarization variability originates from turbulence in the flow, approximated as cilindrical cells

(Marscher 2014)



Flares emitted along the knot path

- If gamma-ray flares are produced along the path of traveling knots, from the burst duration, the travelled path should be ~30-50pc (assuming a bulk Lorentz factor of 10)
- the energy density of seed photon from BLR, and from the dusty torus should be

  $U_{ext} \sim 1/d^2$

- the magnetic field energy density should be reduced of the same amount
- $U_B \sim 1/d^2$
- Both radio knot emission, both gamma-ray emission should be extremely weak toward the path end
- Could acceleration of superluminal radio features compensate for the decrease of energy density along the path?

Magnetic reconnection scenario

at ~1 TeV:

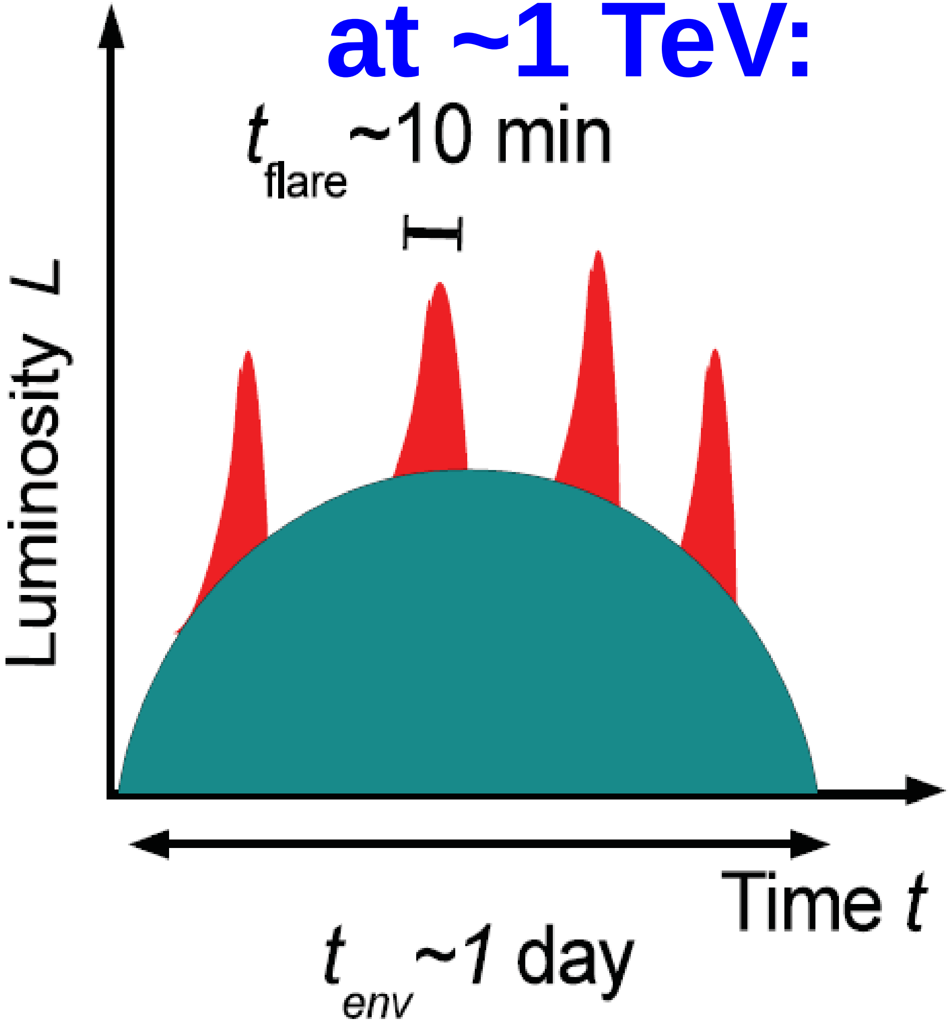


Figure 2. A sketch of the envelope-flare structure of the emission from a reconnection layer. The envelope duration corresponds to that of the reconnection event: $t_{env} = l'/\Gamma_j \epsilon c$. Monster plasmoids power fast flares which show exponential rise and last for $t_{flare} = 0.1 l'/\delta_p c$. For an envelope of ~1 d blazar flaring, the model predicts that monster plasmoids result in ~10-min flares.

Giannios 2013

Recent scenario for magnetic reconnections proposed for strongly magnetized jets (Giannios 2013) includes an envelope emission (lasting ~1 day) powered by plasmoids, together with fast flares (lasting ~10 min) generated by grown “monster plasmoids”.

In low magnetized plasma (such as at several parsec), reconnection time scales are longer, and longer flares (days to weeks) could arise (Giannios 2013).

“Monster plasmoids” contain energetic particles freshly injected by the reconnection event (Uzdensky et al. 2010)

Large variability observed when the reconnection layer is aligned with the jet axis and with the observer line of sight (Christie 2019)

Conclusions

- Waiting time distribution of FSRQs can be modeled with a set of overlapping bursts of flares with burst duration of ~0.6y, burst rate ~1.3y^{-1} + a fast component (for $\Delta t < 1d$)
- If flares are generated while the knot is crossing a stationary feature along the jet
  - the knot size along the jet should be ~2pc
- flares could be generated along the path of a superluminal knot traveling for ~30-50pc
  - but both the magnetic field energy density (why we see those fading radio features), both the external photon energy density decreases with $1/d^2$
  - Could acceleration of superluminal radio features compensate for the decrease of energy density along the path?
- magnetic reconnection could account for the observed waiting times if the magnetic field instabilities generating reconnection events (or the duration of plasma injection) lasts 0.6y, and if the generation rate of instabilities (or the rate of sporadic plasma injection) is ~1.3y
- While the short component represents a small subset of waiting times, it was found during 3/5 Fermi-LAT ToO campaings during bright FSRQs flares. It could witness **not rare cases of structured flares.**